

A solar-like magnetic cycle on the mature K-dwarf 61 Cygni A (HD 201091)

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ABSTRACT

Context. The long-term monitoring of magnetic cycles in cool stars is a key diagnostic in understanding how dynamo generation and amplification of magnetic fields occur in stars similar in structure to the Sun.

Aims. We investigated the temporal evolution of a possible magnetic cycle of 61 Cyg A. The magnetic cycle is determined from 61 Cyg A's large-scale field over its activity cycle using spectropolarimetric observations and compared to the solar large-scale magnetic field.

Methods. We used the tomographic technique of Zeeman Doppler imaging (ZDI) to reconstruct the large-scale magnetic geometry of 61 Cyg A over multiple observational epochs spread over a time span of nine years. We investigated the time evolution of the different components of the large-scale field and compared it with the evolution of the star's chromospheric activity by measuring the flux in three different chromospheric indicators: Ca II H&K, H α and Ca II infrared triplet lines. We also compared our results with the star's coronal activity using *XMM-Newton* observations.

Results. The large-scale magnetic geometry of 61 Cyg A exhibits polarity reversals in both poloidal and toroidal field components, in phase with its chromospheric activity cycle. We also detect weak solar-like differential rotation with a shear level similar to that of the Sun. During our observational time span of nine years, 61 Cyg A exhibits solar-like variations in its large-scale field geometry as it evolves from minimum activity to maximum activity and vice versa. During its activity minimum in epoch 2007.59, ZDI reconstructs a simple dipolar geometry which becomes more complex when it approaches activity maximum in epoch 2010.55. The radial field flips polarity and reverts back to a simple geometry in epoch 2013.61. The field is strongly dipolar and the evolution of the dipole component of the field is reminiscent of solar behaviour. The polarity reversal of the large-scale field indicates a magnetic cycle that is in phase with the chromospheric and coronal cycle.

Key words. dynamo – stars: activity – stars: chromospheres – stars: magnetic field – stars: solar-type

1. Introduction

High resolution observations of the Sun, including over 30 yr of synoptic magnetic field maps, have revealed a coherent picture of the way in which the large-scale magnetic field evolves over the course of the 22-yr solar magnetic cycle. At minimum, the field is concentrated at the polar regions and is mostly axisymmetric with respect to the rotation axis. At activity maximum, the field is no longer concentrated at the polar regions and is restricted to lower latitudes (Hathaway 2010). When the next minimum occurs, the polar field reappears but with the opposite sign to that observed in the previous minimum at the start of the cycle. This cyclic process repeats every 11 yr and it takes 22 yr for the polar magnetic fields to revert back to the same polarity, forming a 22-yr magnetic cycle. During the course of the solar magnetic cycle the geometry of the large-scale field of the Sun also changes dramatically. The Sun varies from an almost dipolar geometry during cycle minimum to a more complex geometry during

cycle maximum (Sanderson et al. 2003; DeRosa et al. 2012). A recent study by DeRosa et al. (2012) has revealed that during cycle minimum the dipolar component of the large-scale solar field dominates the other components. During polarity reversals the quadrupolar and octopolar component of the large-scale field dominates the dipolar component.

When it comes to other solar-type stars, observations do not have even a fraction of the spatial and temporal resolution of the solar observations. Magnetic cycles in other cool stars have been investigated by using two well known proxies of magnetic activity: chromospheric activity (Baliunas et al. 1995; Hall 2008) and coronal activity (Pevtsov et al. 2003; Güdel 2004). The first long-term monitoring of chromospheric activity in solar-type dwarfs was carried out at the Mount Wilson observatory (Wilson 1978; Duncan et al. 1991; Baliunas et al. 1995), where emission in the line cores of Ca II H&K lines was measured. These observations revealed that different solar-type

stars tend to exhibit different levels of activity variation: irregular activity variations in fast rotating young stars, cyclic activity in comparatively older slowly rotating solar-type stars, and a possible Maunder-minimum in some stars (Baliunas et al. 1995). Our investigation of chromospheric activity cycles of Mount Wilson and HARPS stars has also confirmed that solar-like cycles are more likely for slow rotating stars than for fast rotating stars (Boro Saikia et al., in prep.). A further analysis of the Maunder-minimum candidates from the Mt Wilson survey by Schröder et al. (2013) reveals them to be slightly more evolved than the Sun. Coronal activity of cool stars, on the other hand is not intensively monitored compared to their chromospheric activity. Analysis by Hempelmann et al. (1996) produced weak statistical evidence of coronal cycles in cool stars, and since then coronal activity cycles have been detected for a few selected G and K dwarfs. For example, Favata et al. (2008) reported the presence of a coronal activity cycle for the binary system HD81809. Cyclic X-ray activity was also detected for our target star 61 Cyg A (Robrade et al. 2012) and the K1 dwarf α Cen B (Robrade et al. 2012; Ayres 2015). The shortest coronal activity cycle of 1.6 yr was reported for the young active Sun ι Horologii (Sanz-Forcada et al. 2013), where the coronal activity cycle is in phase with its short chromospheric activity cycle.

Although these indirect proxies are reliable indicators of magnetic activity in cool stars, they do not provide any direct information on the strength or orientation of the large-scale magnetic field. Direct measurements of the surface magnetic field rely on the Zeeman effect. By measuring the broadening of spectral lines caused by the Zeeman effect acting on unpolarised spectra (Robinson 1980; Saar 1996; Reiners & Basri 2006), the integrated unsigned magnetic flux, averaged over the entire disk of the star, can be measured. One advantage of this technique is that the small-scale magnetic features also contribute to the total field measurements. This technique does not provide information on the large-scale magnetic field geometry. However the tomographic technique Zeeman Doppler imaging (ZDI; Semel 1989; Donati et al. 1997) uses spectropolarimetric observations to reconstruct the stellar surface magnetic field geometry. Due to cancellation effects this technique is insensitive to small-scale multipolar features, as the overall polarisation signatures of the multipolar field are cancelled out and only the large-scale field is reconstructed. As the ZDI technique provides information about the vector magnetic field, it can provide invaluable insights into the large-scale field geometry as well as the temporal evolution and polarity reversals of the large-scale field.

The large-scale magnetic field of several solar-type stars reconstructed using ZDI over multiple epochs, has revealed their strongly varying magnetic geometry. For the G7 dwarf ξ Bootis A ($0.86 M_{\odot}$, $T_{\text{eff}} = 5551$ K; Morgenthaler et al. 2012), the K2 dwarf ϵ Eridani ($0.856 M_{\odot}$, $T_{\text{eff}} = 5146$ K; Jeffers et al. 2014), and the G0 dwarf HN Peg ($1.085 M_{\odot}$, $T_{\text{eff}} = 5974$ K; Boro Saikia et al. 2015) rapidly varying field geometry was detected, with surprising appearances and disappearances of the azimuthal field for HN Peg and ϵ Eridani. A two-year magnetic cycle was revealed for the planet hosting F7 dwarf τ Bootis ($1.42 M_{\odot}$, $T_{\text{eff}} = 6360$ K; Donati et al. 2008; Fares et al. 2009). τ Boo is a hot Jupiter host where the planet is orbiting at 0.049 AU. No correlation was detected between the large-scale polarity switch and its chromospheric activity (Fares et al. 2009). Further monitoring of τ Boo over three epochs by Fares et al. (2013) has confirmed the previously determined two-year magnetic cycle. Recent new analysis of τ Boo polarised spectra (Mengel et al. 2016) shows a 3:1 ratio between the magnetic and chromospheric cycles. A magnetic cycle was also detected for the G0

Table 1. Summary of the physical parameters of 61 Cyg A.

Parameters	HD 201091	References
Effective temperature, T_{eff} (K)	4545 ± 40	this work
Spectral type	K5V	1
$\log g$	4.75 ± 0.10	this work
Radius, $R_{\odot}$	0.665 ± 0.005	1
Mass, $M_{\odot}$	0.66	1
Rotation period, days	35.7 ± 1.9	this work
..	34.2 ± 3.7	this work
Inclination	70°	this work
$B - V$	1.069	2
Age, Gyr	6.0	3
..	3.6	4
..	2.0	5
..	1.331	6

References. (1) Kervella et al. (2008); (2) Perryman et al. (1997); (3) Kervella et al. (2008); (4) Mamajek & Hillenbrand (2008); (5) Barnes (2007); (6) Marsden et al. (2014). The rotation period of 35.7 ± 1.9 is derived from chromospheric activity measurements and the period of 34.2 ± 3.7 is derived from ZDI (see text for more details).

dwarf HD 78366 ($1.34 M_{\odot}$, $T_{\text{eff}} = 6014$ K; Morgenthaler et al. 2011), where the radial component of the magnetic field exhibits polarity reversals indicating a possible three-year cycle. Polarity reversals were observed in the large-scale field of the G2 dwarf HD 190771 ($0.96 M_{\odot}$, $T_{\text{eff}} = 5834$ K; Petit et al. 2009; Morgenthaler et al. 2011). However, the variability of the large-scale field of HD 190771 is more complex as the polarity reversal in the azimuthal field reported by Petit et al. (2009) is not detected in subsequent epochs. Instead polarity reversal is reported in the radial component of the magnetic field in the subsequent epochs as shown by Morgenthaler et al. (2011). Polarity reversals of the large-scale field have been detected only for a select few cool stars, out of which none exhibit a magnetic cycle period equivalent to the star's chromospheric cycle period.

This paper investigates the variability of the large-scale magnetic field geometry of the solar-type K5 dwarf 61 Cyg A, using a time series of spectropolarimetric observations over nine years. We also investigate the correlation between its mean magnetic field, chromospheric activity, and coronal activity. The stellar parameters of 61 Cyg A are discussed in Sect. 2, followed by instrumental setup and data reduction in Sect. 3. Section 4 covers both direct and indirect field detection. The large-scale magnetic geometry is discussed in Sect. 5, followed by the long-term evolution of the magnetic and activity cycle in Sect. 6. The results are discussed in Sect. 7 and finally the summary in Sect. 8.

2. Physical properties of 61 Cyg A

A solar-type dwarf of spectral type K5, 61 Cyg A forms the well known visual binary 61 Cyg together with the K7 dwarf 61 Cyg B. The binary 61 Cyg is a northern system with a semi-major axis of approximately $24''$ (Malkov et al. 2012). It is the first stellar object to have its parallax measured (Bessel 1838). 61 Cyg A has a parallax of 287.13 ± 1.51 mas (Perryman et al. 1997) and is at a distance of approximately 3.5 pc. It is a slow rotator with a rotation period of 35.7 ± 1.9 days as shown in Table 1, where the rotation period is obtained from long-term chromospheric activity measurements. We also derive a rotation period of 34.2 ± 3.7 days using ZDI (see Sect. 5.3 for more details). 61 Cyg A has a radius of $0.665 \pm 0.005 R_{\odot}$ (Kervella et al. 2008), where the radius was determined using interferometric

data, and is in agreement with the radius of $0.62 R_{\odot}$ (Takeda et al. 2007) calculated by using a stellar evolution code. Using the same code and stellar evolutionary tracks, Takeda et al. (2007) also determined a mass of $0.660 M_{\odot}$. The angular diameter of 1.775 ± 0.013 mas (Kervella et al. 2008) and the bolometric flux of 0.3844×10^{-9} (Mann et al. 2013) can be used to determine a T_{eff} of 4374 K. Applying to our highest S/N spectra our automatic spectral classification tool¹, we find that the photosphere temperature, logarithmic gravity, and metallicity of 61 Cyg A are respectively equal to 4545 ± 40 K, 4.75 ± 0.10 , and -0.18 ± 0.05 . The T_{eff} from our analysis agrees with the literature value of 4525 ± 140 K (Takeda et al. 2007) with improved accuracy. Using the radius and rotation period from Table 1, we determine an equatorial v_e of 0.94 km s^{-1} . This suggests that the $v \sin i$ of 61 Cyg A should be $\leq 0.94 \text{ km s}^{-1}$. A chromospheric age of 1.331 Gyr was determined by Marsden et al. (2014) from the activity age relationship of Wright et al. (2004). Using chromospheric activity and rotation relation, Mamajek & Hillenbrand (2008) determine an age of 3.6 Gyr. In contrast, Kervella et al. (2008) predicted an age of approximately 6 Gyr using evolutionary models combined with interferometric radius measurements. The age estimates of 61 Cyg A from gyrochronology puts its age at approximately 2 Gyr (Barnes 2007). The stellar parameters of 61 Cyg A are summarised in Table 1.

61 Cyg A is a moderately active star and was included in the long-term Mount Wilson survey (Duncan et al. 1991). It was discovered to exhibit an activity cycle of 7.3 ± 0.1 yr (Baliunas et al. 1995). 61 Cyg A was also included in the Lowell Observatory long-term survey, where it was classified as a star with solar-type variability (Hall et al. 2007; Lockwood et al. 2007). A coronal activity cycle of approximately seven years has also been detected for 61 Cyg A (Robrade et al. 2012) using X-ray observations, where the coronal activity cycle is in phase with the chromospheric activity cycle.

3. Instrumental setup and data reduction

3.1. Optical data

The spectropolarimetric data was collected as part of the BCool collaboration² using the NARVAL spectropolarimeter, at the 2.0 m Telescope Bernard Lyot at Pic du Midi Observatory. NARVAL is equipped with a cross dispersed echelle spectrograph, which is a twin of the EsPaDonS spectropolarimeter (Donati 2003) at the 3.6 m Canada-France-Hawaii telescope (CFHT). It covers the full optical wavelength range from 370 nm to 1000 nm, with a resolving power of approximately 65 000 (Aurière 2003). The spectropolarimetric observations of 61 Cyg A were obtained by combining four successive sub-exposures, taken with different half wave rhomb angles. The data is reduced on-site by using a fully automated reduction package, LIBRE-ESPRIT, which generates both intensity (Stokes I) and circularly polarised (Stokes V) spectra. LIBRE-ESPRIT is based on the reduction pipeline ESPRIT developed by Donati et al. (1997). The reduction package also calculates a diagnostic Null profile, which is a control profile, by combining the four sub-exposures in such a way that any polarisation signal is cancelled out. The output from the reduction pipeline is continuum normalised. Table A.1 shows the journal of observations of 61 Cyg A.

Seven data sets comprising seven epochs spanning nine years of both Stokes I and V spectra, were obtained covering epochs

2007.59, 2008.64, 2010.55, 2012.54, 2013.61, 2014.61, and 2015.54. Each epoch contains between nine and sixteen spectra. The spectropolarimetric data from the 2012.54 epoch was discarded due to polarisation anomalies in one of the polarisation rhombs at NARVAL during 2011–2012³. The azimuthal angle of the rhomb was wrong during that epoch, leading to incorrect polarisation measurements.

For active cool stars, Zeeman-induced circular polarisation has too low an amplitude to be detected in individual spectral lines. Hence, the technique of least square deconvolution (LSD) is applied to the spectra of 61 Cyg A. LSD is a multiline technique which assumes a similar line profile for all magnetically sensitive lines in a spectra and generates an averaged line profile by deconvolving the stellar spectra to a line mask (Donati et al. 1997; Kochukhov et al. 2010). We used a K5 mask consisting of approximately 12 000 lines to compute the LSD Stokes I and Stokes V profiles of 61 Cyg A. The LSD technique used in this paper is detailed in Marsden et al. (2014).

3.2. X-ray data

61 Cyg A is a strong X-ray source and has been continuously monitored with *XMM-Newton* since 2002 with a cadence of two observations a year (see e.g. Hempelmann et al. 2006; and Robrade et al. 2012). Therefore, in order to check for periodicity in X-ray luminosity we obtained all public data sets from the *XMM-Newton* Science Archive. Data reduction was performed with the *XMM-Newton* Science Analysis System version 10.0, and XSPEC 12.7.1. In order to avoid contamination from 61 Cyg B, source photons were extracted with a $15''$ aperture, while background levels were estimated with a $45''$ aperture in source-free regions. Time intervals during which obvious flaring occurred were omitted. The spectra of all three EPIC detectors (MOS 1&2 and PN) were then fitted simultaneously in the 0.2–2.0 keV range with a three-temperature APEC model.

4. Magnetic field detection: direct and indirect approach

4.1. Mean longitudinal magnetic field

We measure the line-of-sight component of the magnetic field, averaged over the stellar surface for each observation using Stokes V and Stokes I LSD profile. The mean longitudinal field is measured from the Stokes V and I profile (Donati et al. 1997) as shown in Eq. (1),

$$B_l(\text{G}) = -2.14 \times 10^{11} \frac{\int v V(v) dv}{\lambda_0 g c \int (I_c - I(v)) dv} \quad (1)$$

where B_l represents the mean longitudinal magnetic field of 61 Cyg A in Gauss, $\lambda_0 = 630$ nm is the central wavelength of the LSD profile, $g = 1.22$ is the Landé factor of the line list, c is the speed of light in km s^{-1} , v is the radial velocity in km s^{-1} , and I_c is the normalised continuum intensity. A velocity range of $\pm 20 \text{ km s}^{-1}$ from the line centre was used. The velocity window was selected so that the entire signal could be included. If a smaller integration window is used, then the wings of the signal are ignored which may result in an underestimated value of B_l . We also calculated the magnetic field from the Null profile (N_l). The error estimates were made by propagating the uncertainties calculated by the reduction pipeline, as mentioned in

¹ Our spectral classification tool is inspired from that of Valenti & Fischer (2005) and is discussed in a previous paper (Donati et al. 2012).

² <http://bcool.ast.obs-mip.fr/Bcool>

³ <http://spiptbl.bagn.obs-mip.fr/Actualites/Anomalies-de-mesures>

Marsden et al. (2014). Table A.2 shows the values of B_l , N_l , and the associated uncertainties for each observation.

4.2. Chromospheric activity as a proxy of magnetic activity

61 Cyg A is known to exhibit cyclic chromospheric activity of approximately seven years (Baliunas et al. 1995). In order to investigate the chromospheric activity of 61 Cyg A, we monitored the fluxes in three chromospheric indicators: Ca II H&K, H α , and the Ca II infrared triplet. These three chromospheric indicators are well known proxies of stellar activity, where the different indicators have different sensitivities to the various chromospheric features such as plages, networks, flares, and filaments (Meunier & Delfosse 2009). The Ca II H&K lines are resonance lines which are dependent on the local chromospheric temperature and electron density (Jefferies & Thomas 1959). Contrary to the Ca II H&K lines, the H α line is a Balmer line which is insensitive to the local chromospheric conditions but more sensitive to the radiation temperature (Jefferies & Thomas 1959). Hence, it is important to explore the different chromospheric tracers to obtain as complete a picture as possible of the chromosphere of an active cool star.

4.2.1. S-index

The S-index has been one of the most widely used proxies of magnetic activity since its introduction in the long-term Mount Wilson survey (Wilson 1978; Duncan et al. 1991; Baliunas et al. 1995). We measure the S-index of 61 Cyg A by using two triangular band passes with a FWHM of 0.1 nm, centred at the line cores of the Ca II K and H line at 393.3663 nm and 396.8469 nm respectively. The fluxes in the nearby continuum are measured by using two 2 nm wide rectangular band passes at the blue and red side of the K and H lines at 390.107 nm and 400.107 nm respectively (Duncan et al. 1991; Wright et al. 2004). In order to calibrate our measurements to the original Mount Wilson scale, we use the calibration coefficients calculated by Marsden et al. (2014), which is used in the S-index formula as shown in Eq. (2),

$$\text{S-index} = \frac{\alpha F_H + \beta F_K}{\gamma F_R + \delta F_V} + \phi \quad (2)$$

where F_H , F_K , F_R , and F_V are the flux in the H, K, R, and V bands and $\alpha = 12.873$, $\beta = 2.502$, $\gamma = 8.877$, $\delta = 4.271$, and $\phi = 1.183\text{e-}03$ are the calibration coefficients respectively. The uncertainties associated with the S-index measurements are obtained by carrying out error propagation and are tabulated in Table A.2.

Figure 1 shows the rotational variability of both the B_l and the S-index as a function of the rotational phase for each epoch, allowing us to study variability on the rotation timescale. Overall, 61 Cyg A exhibits a stronger rotational variability in its longitudinal magnetic field during epochs 2007.59, 2010.55 and 2015.54 compared to epochs 2008.64, 2013.61 and 2014.61. During our observational time span the longitudinal field B_l shows a change in sign which indicates that the field is non-axisymmetric, except in epochs 2013.61 and 2014.61 which show signs of a more axisymmetric field. The S-index exhibits low rotational dependence, indicating a simple magnetic geometry. However, one should be careful in interpreting the results because the rotational variations might be hidden due to poor phase coverage in certain epochs such as 2012.54. The longitudinal field could not be measured for epoch 2012.54, due to the

unavailability of Stokes V observations as already mentioned in the previous section.

61 Cyg A was part of the original Mount Wilson survey and was observed for three decades at the Mount Wilson observatory from 1961 to 1991 (Baliunas et al. 1995). A continuation of the Mount Wilson survey is being currently carried out at the Lowell observatory (Hall et al. 2007) and all available S-index data of 61 Cyg A together with our measurements are shown in Fig. 2. It shows the temporal variation of the S-index of 61 Cyg A calibrated to the Mount Wilson scale over a time span of 48 yr. At the beginning of our observational data in epoch 2007.59, 61 Cyg A exhibits low activity and it increases to its maximum during epoch 2010.55. The chromospheric activity is low in epoch 2013.61, which indicates it is close to activity minimum. The activity in epoch 2014.61 is also very low but it increases again during epoch 2015.54. An activity cycle period of 7.2 ± 1.3 yr was measured by taking a “generalised” Lomb-Scargle periodogram (Lomb 1976; Scargle 1982; Zechmeister & Kürster 2009) on all available S-index data. This cycle period agrees with the activity cycle period of 7.3 ± 0.1 yr, measured using only the Mount Wilson data (Baliunas et al. 1995).

4.2.2. H α -index

We measured the variability of the chromospheric activity of 61 Cyg A using H α as an activity tracer for each epoch of this analysis. We employed the technique used by Gizis et al. (2002), where a rectangular band pass of 0.36 nm width was used at the H α line at 656.285 nm. Two rectangular band passes H_{blue} and H_{red} of 0.22 nm width were used to measure the nearby continuum,

$$\text{H}\alpha\text{-index} = \frac{H\alpha}{H_{\text{blue}} + H_{\text{red}}} \quad (3)$$

where $H\alpha$, H_{blue} , and H_{red} refers to different band passes used.

4.2.3. CaIRT-index

The CaIRT-index of 61 Cyg A was measured by taking 0.2 nm wide rectangular band passes centred at the Ca II triplet lines at 849.8023 nm, 854.2091 nm, and 866.241 nm respectively. The nearby continuum is measured by taking two 0.5 nm band passes IR $_{\text{red}}$, and IR $_{\text{blue}}$ at 870.49 nm, and 847.58 nm respectively (Marsden et al. 2014), as shown in Eq. (4),

$$\text{CaIRT-index} = \frac{\text{IR1} + \text{IR2} + \text{IR3}}{\text{IR}_{\text{red}} + \text{IR}_{\text{blue}}} \quad (4)$$

where IR1, IR2, and IR3 represent flux in the Ca II infra-red triplet lines. The continuum flux is represented by IR $_{\text{red}}$ and IR $_{\text{blue}}$ respectively.

The H α and Ca IRT lines are in absorption throughout our observations. The H α -index variability as a function of the rotational phase is shown in Fig. 3. The rotational variability of the CaIRT-index is also shown in Fig. 3. Surprisingly the H α and Ca II IRT indices do not show the strong variations shown by the S-index. This might indicate that the magnetic field is not strong enough for these lines to go into emission explaining such low variations as compared to the S-index. The uncertainties are obtained by carrying out error propagation and are shown in Table A.2.

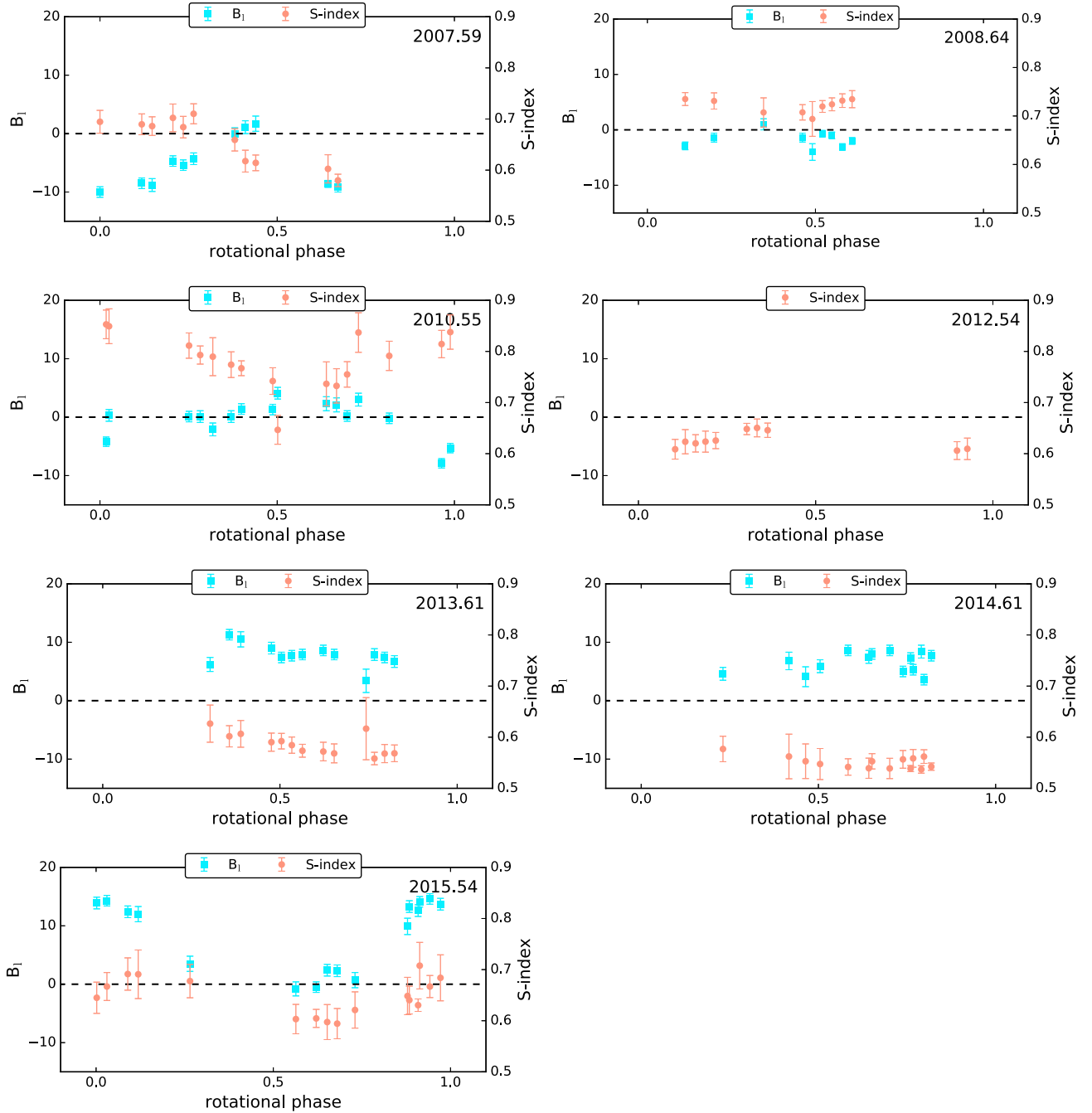


Fig. 1. Variability of the longitudinal field (B_l) (left y -axis) and S-index (right y -axis) as a function of rotational phase for each epoch (2007.59, 2008.64, 2010.55, 2012.54, 2013.61, 2014.61 and 2015.54). For epoch 2012.54 only the S-index is plotted, as the Stokes V observations were not available (see Sect. 3). The subplots are all on the same scale for easier comparison between epochs, and the dashed line represents $B_l = 0$. The scale was chosen based on the minimum and maximum value.

4.2.4. Coronal activity

The X-ray luminosity variation of 61 Cyg A is shown in Fig. 5, where the data up to end of 2011 were previously published by Robrade et al. (2012) and the results from 2012 onward are derived for the first time here. While comparison to the results of Robrade et al. (2012) yields minor discrepancies caused by the different data fitting approaches, the overall picture is consistent. A clear cyclic variation in the X-ray luminosity is recovered with a cycle period of 6.6 ± 0.5 yr. This is in good agreement with the chromospheric cycle of 7.2 ± 1.3 yr.

4.2.5. Correlation between activity indicators and longitudinal field

As shown in Fig. 4, a possible weak anti-correlation between the absolute value of the longitudinal magnetic field $|B_l|$ and the S-index is detected. The field measurements are the weakest in epoch 2010.55, when the chromospheric activity is at a maximum in our observations, and are the strongest close to minimum chromospheric activity. The average of the absolute longitudinal magnetic field ($|B_l|$) per epoch is strongest in epoch 2015.54 with 9.3 ± 5.9 G, and is weakest in epoch 2010.55 with 0.6 ± 3.1 G. On the contrary strong correlation is detected between the three

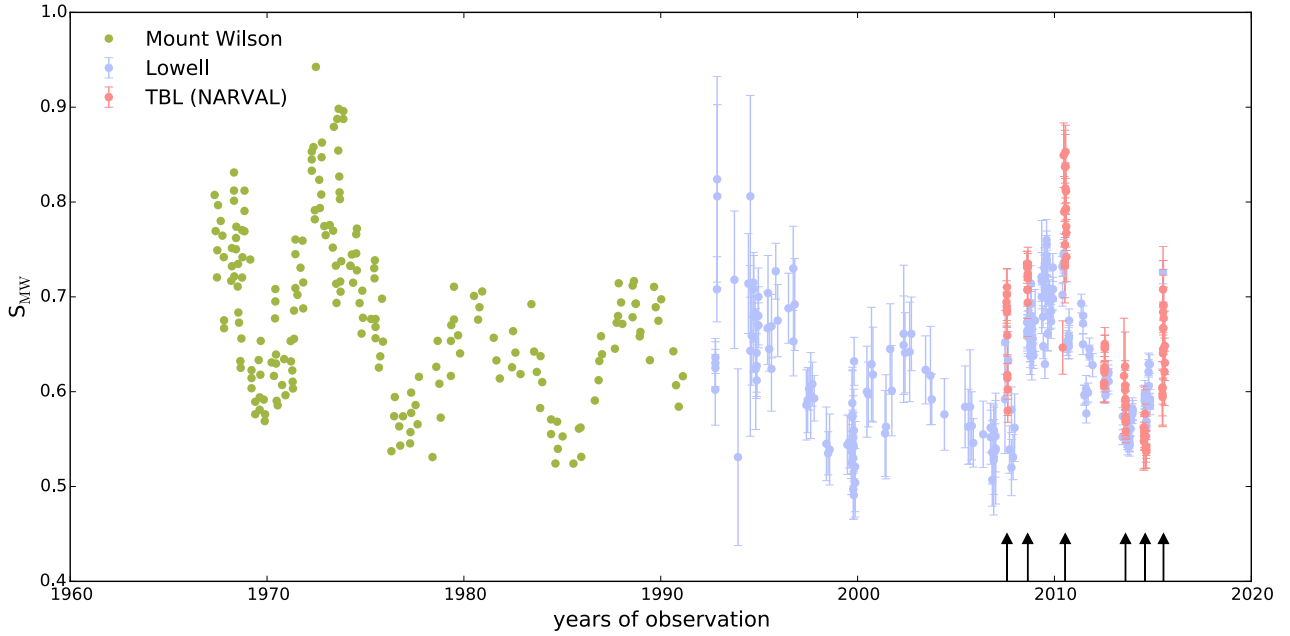


Fig. 2. Long-term chromospheric activity of 61 Cyg A shown as S-index, which is calibrated to the Mount Wilson S-index. The arrow marks represent epochs for which the large-scale field map is reconstructed using ZDI.

chromospheric activity indicators as shown in Fig. 4. The CaIRT-index exhibits a slightly stronger correlation with the S-index ($r = 0.7$) compared to the $H\alpha$ -index ($r = 0.6$), which is not surprising as those two indicators are sensitive to similar chromospheric features. A strong correlation between the S-index and the CaIRT-index was also detected in the G dwarf ξ Bootis A by Morgenthaler et al. (2012). Correlation between the $\log R'_{\text{HK}}$ and CaIRT-index was also studied for 170 solar-type stars as part of the BCoolest snapshot survey (Marsden et al. 2014). The $\log R'_{\text{HK}}$ is also an indicator of stellar activity where the photospheric contribution is corrected for the S-index (Wright et al. 2004).

5. Large-scale magnetic field geometry

The large-scale surface magnetic geometry of 61 Cyg A is reconstructed using the ZDI technique. ZDI is an inverse tomographic technique which reconstructs the stellar magnetic geometry from a time series of Stokes V LSD profiles using maximum entropy as a regularisation technique (Skilling & Bryan 1984). The ZDI version used in this paper reconstructs the magnetic geometry into its poloidal (potential field) and toroidal (non-potential field) components expressed as spherical harmonic decomposition (see Donati et al. 2006).

If the total surface magnetic field is a combination of poloidal and toroidal field, it can be formalised into three field components B_r , B_θ , and B_ϕ . B_r is the radial component of the field and is entirely poloidal, described by the spherical harmonics coefficient $\alpha_{l,m}$, where l and m are the order and degree respectively. B_θ is the meridional component of the total field and B_ϕ is the azimuthal component of the total field. B_θ and B_ϕ can be described by a combination of the spherical harmonics coefficients $\beta_{l,m}$ and $\gamma_{l,m}$ (Donati et al. 2006, see Sect. 5.1). The order of the spherical harmonics coefficients determines if the field is simple dipolar or complex, and the degree determines if the field is axisymmetric or non-axisymmetric. For example, $l = 1, m = 0$ denotes the axisymmetric dipolar field. The quadrupolar and octopolar modes are denoted by higher order

$l = 2$ and $l = 3$ respectively. The surface magnetic map and the associated Stokes V spectra of a star can be derived using the spherical harmonics coefficients. ZDI tackles the inverse problem and reconstructs the large-scale surface magnetic field geometry by comparing model Stokes V spectra to the observed Stokes V spectra.

For 61 Cyg A we use a stellar model of 5000 grid points, where each grid point is associated with both a local Stokes V and Stokes I profile. A solar-type limb darkening law is used in the stellar model and synthetic Stokes V profiles are computed for all rotational phases. Each of these local Stokes V profiles are calculated under the weak field assumption and a uniform Stokes I profile is assumed over the entire photosphere. Using maximum entropy as a regularisation scheme (Skilling & Bryan 1984), ZDI reconstructs $\alpha_{l,m}$, $\beta_{l,m}$, and $\gamma_{l,m}$ by iteratively comparing the synthetic Stokes V profiles to the observed Stokes V profiles.

5.1. Vector magnetic field

61 Cyg A exhibits chromospheric variability on rotational timescales, indicating a high stellar inclination. We choose an inclination angle of 70° , which together with a rotation period of 34.2 days and a stellar radius of $0.665 R_\odot$, results in a $v \sin i$ of 0.92 km s^{-1} . This $v \sin i$ is consistent with the v_e of 0.94 km s^{-1} mentioned in Sect. 2. We therefore used the $v \sin i$ of 0.92 km s^{-1} and an inclination angle of 70° in our magnetic field reconstructions. We tested a range of spherical harmonics orders $l_{\text{max}} = 5, 7, 9, 11$ to reconstruct the magnetic maps. For all epochs of our analysis except 2010.55 and 2015.54, the target χ^2 was achieved for $l_{\text{max}} = 7$. Using $l_{\text{max}} \geq 7$ does not improve the χ^2 or the fit to the spectra at epochs 2007.59, 2008.64, 2013.61, and 2014.61. However, for epochs 2010.55 and 2015.54, the best χ^2 fit was obtained for $l_{\text{max}} \geq 11$. To make comparisons between the epochs we kept the l_{max} to 11 throughout our analysis. We also include differential rotation in the large-scale field reconstructions (see Sect. 5.3 for more details).

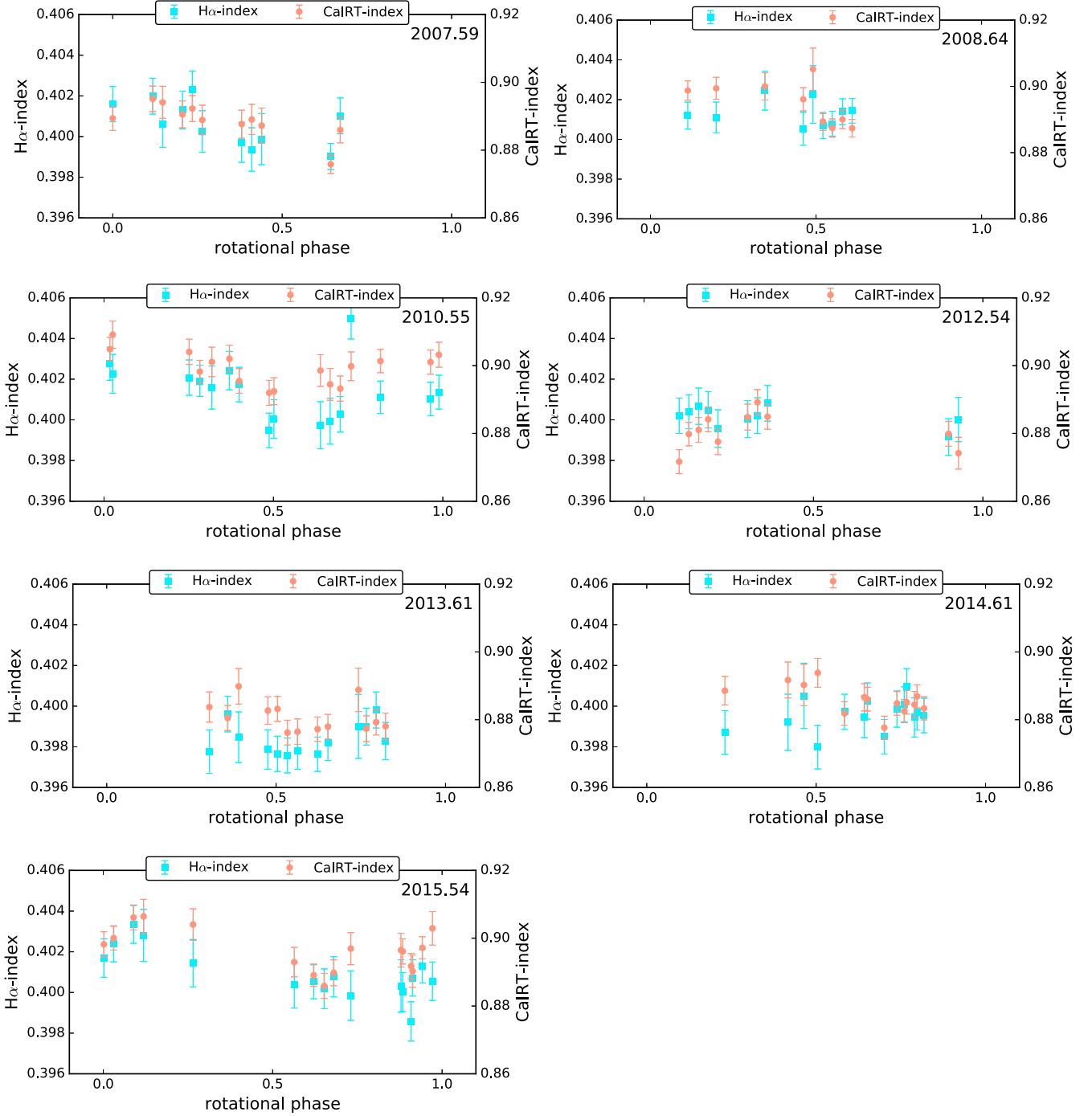


Fig. 3. Same as Fig. 1 for $H\alpha$ -index (left y -axis) and CaIRT-index (right y -axis). The scales are chosen based on the minimum and maximum value.

Figure 6 shows the Stokes V LSD time series of 61 Cyg A for our observational time span. The fit between the modelled Stokes V profiles and observed Stokes V profiles was obtained with a reduced χ^2 of 1.0 for epoch 2007.59. A reduced χ^2 of 1.1 was achieved for epochs 2008.64, 2013.61, and 2014.61. For epochs 2010.55 and 2015.54 the observed fit was obtained with a reduced χ^2 of 1.2 and 1.5 respectively. It is evident from Fig. 6 that during epoch 2007.59, which is around activity minimum, the time series of Stokes V profiles is dominated by a single sign. The Stokes V profiles in epoch 2010.55 are not dominated by a single orientation. The high variability shown by the Stokes V profile in 2010.55 is caused by a complex magnetic

field geometry. This complexity of the field is likely to make the impact of differential rotation more apparent in this epoch. In the following three epochs (2013.61, 2014.61, and 2015.54) the Stokes V profiles flip sign and are dominated again by a single orientation. This evolution of the Stokes V time series indicates that the surface magnetic field geometry is strongly evolving, with a single polarity field during low activity and a more complex field during strong activity.

Since the maximum entropy reconstruction does not allow error calculations of our magnetic maps, we employ the technique described by Petit et al. (2008), where a range of magnetic maps were calculated for the wide range of input stellar

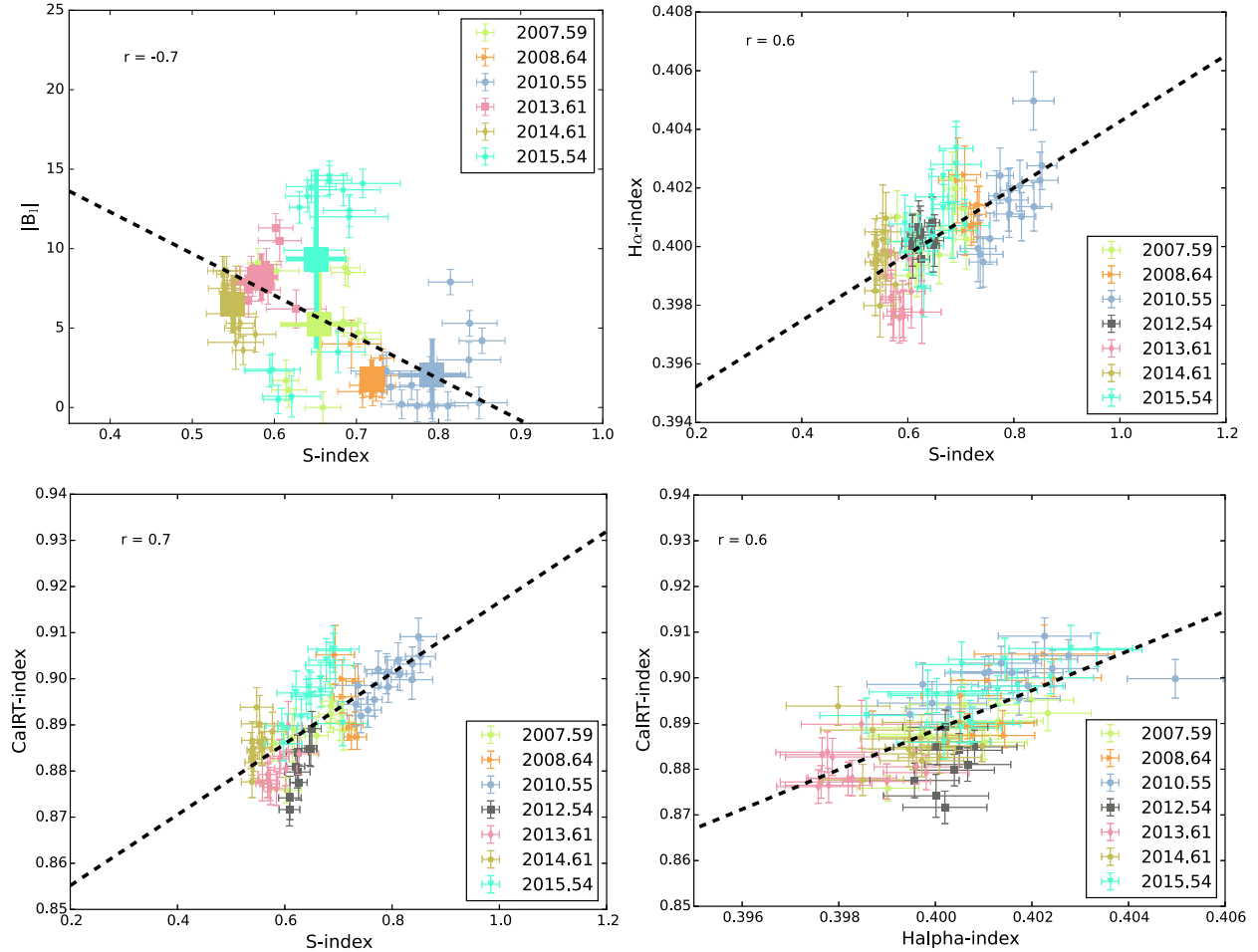


Fig. 4. Correlation between the different activity indicators and the longitudinal field $|B_l|$, where the dashed line represents least square fit to the data and the Pearson's r is shown in the plot. *Top left:* anti-correlation between S-index and $|B_l|$ over six epochs of observations. The bold crosses represent the mean and standard deviation as dispersion. Only the mean values are included in the fit. *Top right:* plot showing the correlation between S-index and H α -index over seven epochs of observations. *Bottom left:* correlation between S-index and CaIRT-index over seven epochs of data. *Bottom right:* correlation between H α -index and CaIRT-index over seven epochs of data.

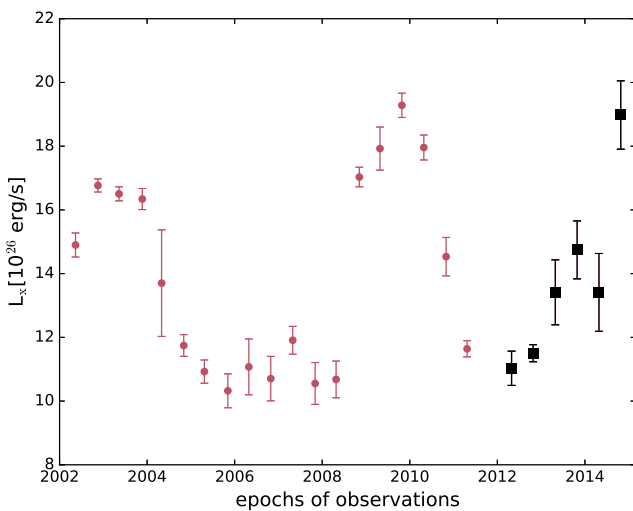


Fig. 5. Variation of X-ray luminosity of 61 Cyg A determined through *XMM-Newton* observations. The black squares show new previously unpublished results.

parameters used in our ZDI code. These input parameters are varied within their error bars and the resulting dispersion is taken as error bars, as shown in Table 2.

5.1.1. Epoch 2007.59

The large-scale surface magnetic field of 61 Cyg A at epoch 2007.59 exhibits a simple dipolar magnetic geometry. Figure 7 shows this large-scale magnetic field geometry reconstructed into three components: radial, azimuthal and meridional. The radial component of the magnetic field exhibits strong negative-polarity magnetic regions at higher latitudes and positive-polarity magnetic field at lower latitudes. On the other hand, the azimuthal field component is dominated by two positive polarity magnetic field regions between the pole and lower latitudes. The B_{mean} , which is the average of the surface magnetic field from the reconstructed maps, is 12 ± 3 G. The poloidal magnetic field dominates, accounting for $93 \pm 5\%$ of the magnetic energy seen in Stokes V. The fraction of the magnetic energy in other field components is shown in Table 2. As shown in Fig. 8, almost 100% of the poloidal magnetic energy is stored in the lower order spherical harmonics modes, $l \leq 3$.

5.1.2. Epoch 2008.64

In epoch 2008.64 the large-scale magnetic field in the radial component exhibits a similar magnetic geometry as reconstructed for epoch 2007.59, shown in Fig. 7. However, the mean magnetic field (B_{mean}) is weaker than in the previous epoch,

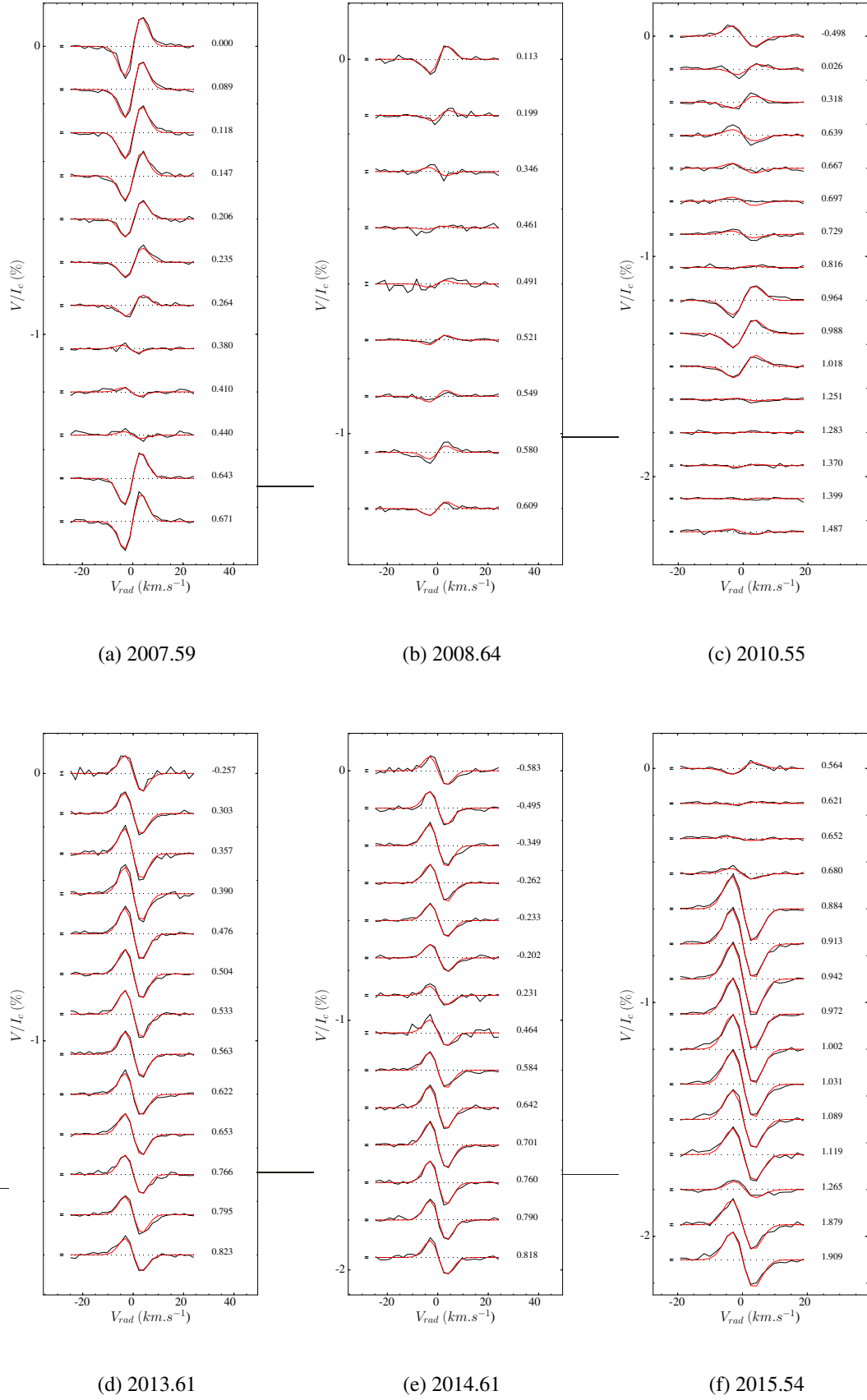
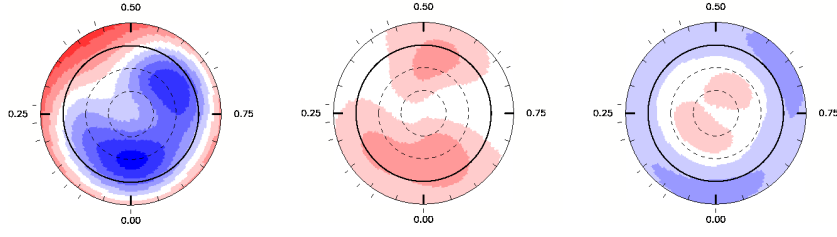
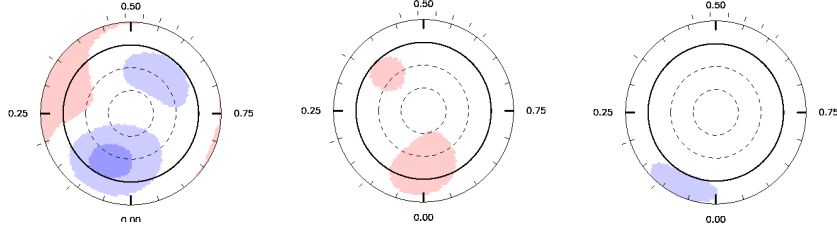


Fig. 6. Time series of LSD Stokes V spectra of 61 Cyg A, where *top row* (from left to right) represents epochs 2007.59, 2008.64, and 2010.55, and *the bottom row* shows (left to right) epochs 2013.61, 2014.61, and 2015.54. The observed spectra is shown in black and the corresponding fit is shown in red. For each observed spectra, the phase of the rotational cycle is shown on the right, and on the left side of each plot 1σ error is shown. For the purpose of clarity successive Stokes V profiles are shifted vertically and the scale is expanded by a factor of two.

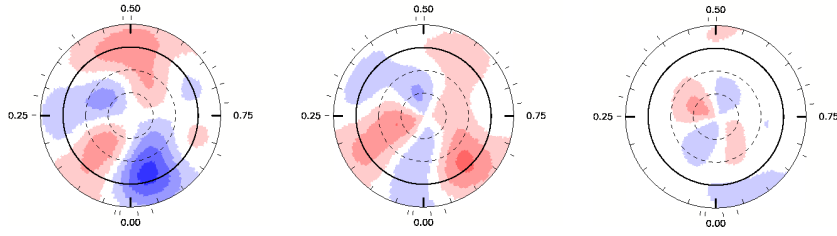
2007.59



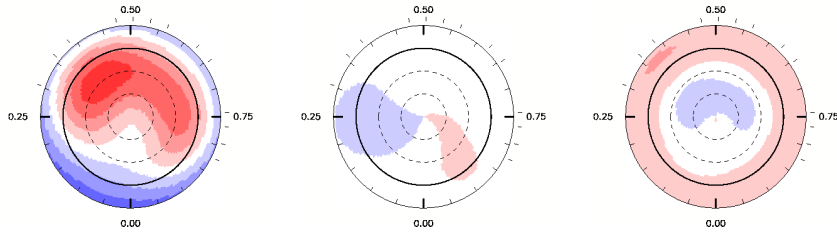
2008.64



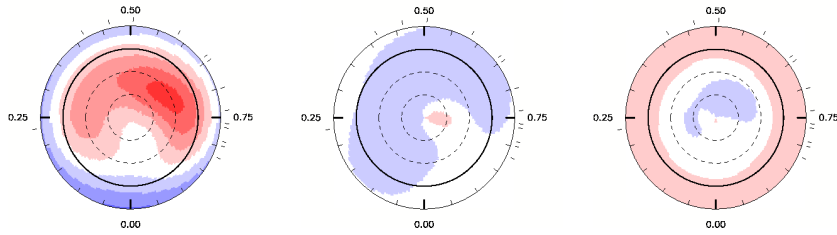
2010.55



2013.61



2014.61



2015.54

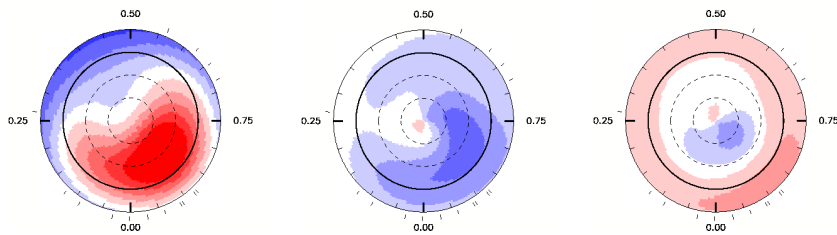
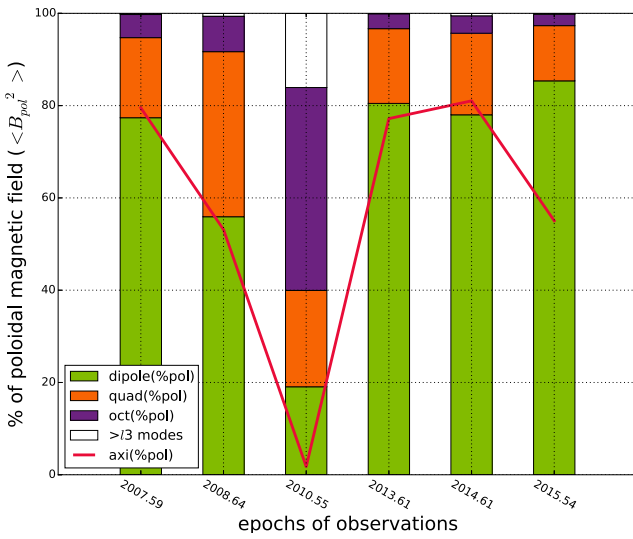


Fig. 7. Surface magnetic geometry of 61 Cyg A over six observational epochs (2007.59, 2008.64, 2010.55, 2013.61, 2014.61, and 2015.54). The three columns represent the radial, azimuthal, and the meridional component of the large-scale field. The star is shown in flattened polar projection down to latitudes of -30° , where the dark line depicts the stellar equator.

Table 2. Magnetic properties of 61 Cyg A extracted from the ZDI maps.

Dates	No of obs	v_r (km s ⁻¹)	B_{mean} (G)	Poloidal (%tot)	Dipole (%pol)	Quad (%pol)	Oct (%pol)	Axi (%tot)	Axi (%pol)	Ω_{eq} (rad d ⁻¹)	d Ω (rad d ⁻¹)
2007.59	12	-65.47 ± 0.05	12 ± 3	93 ± 5	77 ± 9	17 ± 7	5 ± 3	80 ± 13	80 ± 17		
2008.64	9	-65.48 ± 0.05	3 ± 1	92 ± 2	56 ± 5	36 ± 5	8 ± 2	56 ± 10	53 ± 11		
2010.55	16	-65.48 ± 0.05	5 ± 2	87 ± 3	19 ± 6	21 ± 5	44 ± 13	7 ± 2	2 ± 4	0.18 ± 0.03	0.04 ± 0.02
2013.61	13	-65.49 ± 0.05	9 ± 6	99^{+1}_{-3}	80 ± 9	16 ± 7	3 ± 1	77 ± 6	77 ± 6		
2014.61	14	-65.50 ± 0.05	8 ± 4	93 ± 4	78 ± 8	18 ± 6	4 ± 3	81 ± 7	81 ± 9		
2015.54	15	-65.52 ± 0.05	12 ± 5	87 ± 5	85 ± 9	12 ± 6	2 ± 1	59 ± 12	55 ± 20		

Notes. The columns represent fractional dates, number of observations, radial velocity (v_r), mean magnetic field strength (B_{mean}), fraction of magnetic energy reconstructed as the poloidal component, and fraction of poloidal magnetic field stored as dipole, quadrupole, and octopole modes. Finally it also shows the fraction of the total magnetic energy in the axisymmetric component of the magnetic field, and the poloidal axisymmetric fraction (for $m \leq \frac{l}{2}$). The differential rotation parameters Ω_{eq} and d Ω are also shown in the last two columns.

**Fig. 8.** Poloidal magnetic field distribution of 61 Cyg A for six epoch of observations.

as shown in Table 2, where B_{mean} is 3 ± 1 G. In the azimuthal field component the positive polarity magnetic regions seen in the previous epoch are almost non-existent, as shown in Fig. 7. The magnetic field is predominantly poloidal, constituting $92 \pm 2\%$ of the magnetic energy. The dipolar component of the poloidal field accounts for $56 \pm 5\%$ of the poloidal magnetic energy. As in the previous epoch, the lower order modes ($l \leq 3$) account for almost 100% of the poloidal magnetic energy budget as shown in Fig. 8.

5.1.3. Epoch 2010.55

The large-scale magnetic field of 61 Cyg A changes dramatically in epoch 2010.55. The high-latitude negative polarity field reconstructed in epoch 2007.59 is absent from this epoch and the radial field component is dominated by a more complex magnetic geometry between the equator and the poles, as shown in Fig. 7. The azimuthal field component also exhibits a more complex geometry compared to previous epochs. The averaged B_{mean} for epoch 2010.55 is 5 ± 2 G, which is still weaker than epoch 2007.59, as shown in Table 2. The poloidal field dominates and constitutes $87 \pm 3\%$ of the magnetic energy. Only $19 \pm 6\%$ of the poloidal field is dipolar, $21 \pm 5\%$ is quadrupolar

and $44 \pm 13\%$ of the poloidal field is reconstructed into its octopolar component, as shown in Table 2. The field is also the least axisymmetric in this epoch. Unlike previous epochs, a few % of the poloidal magnetic field is also distributed in the higher order spherical harmonics modes, as shown in Fig. 8. The percentage of the magnetic field reconstructed into its toroidal component is higher at this epoch when compared to the previous epochs.

5.1.4. Epoch 2013.61

The large-scale magnetic field geometry at epoch 2013.61 is shown in Fig. 7, where higher latitude magnetic regions are reconstructed in the radial field component with opposite polarity, as seen in epochs 2007.59 and 2008.64. The azimuthal field component exhibits mixed polarity but it is weak and almost non-existent, as shown in Fig. 7. The mean magnetic field derived from the magnetic maps ($B_{\text{mean}} = 9 \pm 6$ G) is stronger than the previous two epochs. The poloidal component comprises $99^{+1}_{-3}\%$ of the magnetic energy, as shown in Table 2. The poloidal magnetic field is $80 \pm 9\%$ dipolar. The distribution of the magnetic field is primarily concentrated at the lower spherical harmonics modes (dipolar, quadrupolar, and octopolar) as shown in Fig. 8.

5.1.5. Epoch 2014.61

The strong positive field in epoch 2013.61 is also reconstructed in the radial component in epoch 2014.61, as shown in Fig. 7. The azimuthal field, as shown in Fig. 7, is reconstructed as a band of negative-polarity magnetic field at the surface of the star, where the field is of opposite polarity, as seen in epoch 2007.59. Table 2 shows the mean magnetic field B_{mean} of epoch 2014.61, where the field strength is 8 ± 4 G. The fraction of the magnetic field distributed into the different field components is shown in Table 2, where $93 \pm 4\%$ of the magnetic field is reconstructed into its poloidal component. $78 \pm 8\%$ of the poloidal field is in its dipolar component. The magnetic energy budget is spread across the lower spherical harmonics mode $l \leq 3$, as shown in Fig. 8.

5.1.6. Epoch 2015.54

During epoch 2015.54 the radial field component exhibits a positive polarity magnetic region between the poles and the equator, together with a negative polarity magnetic field at the equator, as shown in Fig. 7. A band of negative polarity magnetic field is reconstructed in the azimuthal component of the magnetic field, as shown in Fig. 7. The mean magnetic field B_{mean} of 12 ± 5 G is similar to that measured in epoch 2007.59. $87 \pm 5\%$ of the

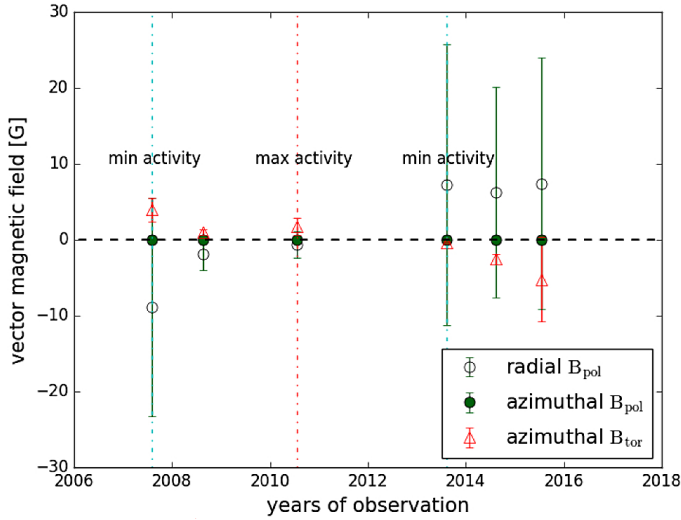


Fig. 9. Vector component of the poloidal and toroidal magnetic field strength averaged over the stellar surface. The field strengths in the radial and azimuthal components are shown. The horizontal dashed line shows where the field strength is 0. The cyan vertical dashed lines show the epochs with minimum activity and the red vertical dashed line shows the epoch with maximum chromospheric activity.

magnetic field is in the poloidal component, as shown in Table 2. The percentage of the poloidal field reconstructed into the dipolar component is $85 \pm 9\%$ and, as shown in Fig. 8 the poloidal field is distributed across the lower spherical harmonics modes.

5.2. Evolution of the different multipolar modes over the magnetic cycle

In order to assess the long-term variability of the magnetic field of 61 Cyg A, we consider several ways of paraphrasing the reconstructed magnetic field. First, we measure the vector component of the magnetic field for each epoch, averaged over the stellar surface. The vector radial field component, as mentioned earlier in this section, is entirely poloidal in nature and the signed value of this component is shown in Fig. 9. The field is averaged over the stellar surface for co-latitudes less than 30° , which is close to the rotational pole. Here co-latitudes are surface co-latitudes relative to the stellar rotational pole in spherical co-ordinates. The vector azimuthal field, on the other hand, is comprised of both poloidal and toroidal field components in this case. Therefore, the signed azimuthal poloidal field and the signed azimuthal toroidal field are treated separately and calculated by averaging the surface for co-latitudes less than 30° . As seen in Fig. 9, the poloidal field strength B_{pol} in the radial component is weakest in epoch 2010.55 and is stronger during epochs of low chromospheric activity. The poloidal component of the azimuthal field is almost negligible for the total azimuthal field. From Fig. 9 it is clear that the poloidal field strength on the stellar surface, for co-latitudes $\leq 30^\circ$, is mostly radial. The toroidal field strength B_{tor} exhibits weak anti correlation to the radial field strength B_{pol} .

Since the axisymmetric modes are of greater interest for existing dynamo models, we investigate the time evolution of both even and odd axisymmetric modes ($m = 0$) for $l = 1, 2, 3, 4$. Figure 10 shows the field strength, at the pole of the axisymmetric modes of the poloidal field, for the first four orders of spherical harmonics. The poloidal field strength in the axisymmetric modes is the signed (radial) field strength at the visible rotational pole of the star. The first odd mode $l = 1$, which is the

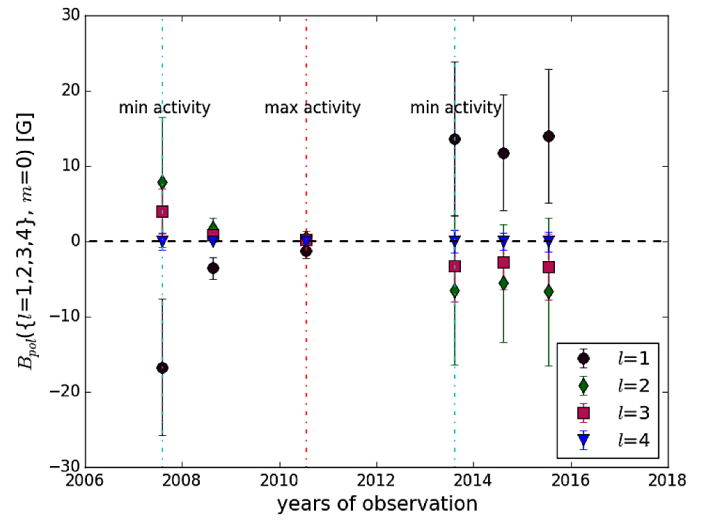


Fig. 10. Magnetic field strength B_r at the magnetic pole for the different axisymmetric modes ($l = \{1, 2, 3\}$). The horizontal dashed line represents $B_r = 0$, and the vertical dashed lines represents the epochs with minimum and maximum activity as shown in Fig. 9.

dipolar mode, exhibits similar behaviour to the full averaged B_{pol} in Fig. 9. The quadrupolar $l = 2$ mode is anti-correlated with the dipolar mode. The octopolar mode ($l = 3$) exhibits correlation with the quadrupolar mode.

Figure 11 shows the field strength of only the lowest order ($l = 1$) axisymmetric ($m = 0$) poloidal and toroidal modes. The poloidal field strength is the same as that shown in Fig. 10, but only the dipolar mode is plotted. The evolution of the dipolar field is consistent with the polarity reversals of the large-scale radial field shown in Fig. 7 and shows strong correlation with the averaged radial poloidal field strength in Fig. 9. The signed toroidal field strength is calculated by measuring the strength of the azimuthal component of the toroidal ($l = 0, m = 0$) field, evaluated at the maximum of the toroidal band. The toroidal field strength is anti-correlated with the poloidal field strength in Fig. 11.

The uncertainties were estimated using the same technique as described in Sect. 5.1.

5.3. Differential rotation

Each epoch spanned over multiple weeks; we therefore investigated the effect of differential rotation on our data. In order to investigate the differential rotation parameters of 61 Cyg A, we incorporated a simple solar-type differential rotation law into our ZDI code, as described by Petit et al. (2002),

$$\Omega(l) = \Omega_{\text{eq}} - d\Omega \sin^2 l \quad (5)$$

where $\Omega(l)$ represents the stellar rotation at latitude l , Ω_{eq} is the equatorial rotation rate, and $d\Omega$ is the difference between the rotation rate at the equator and the poles.

We set the magnetic field strength to a fixed value and carry out ZDI reconstruction for a wide range of differential rotation parameters Ω_{eq} and $d\Omega$ in a 2D parameter space. The set of differential rotation parameters for which the χ^2 minimum is achieved is then selected, as shown in Fig. 12. The uncertainties were calculated by the bootstrap technique, where the input stellar parameters were varied within their error bars and the differential rotation was measured for this wide range of input parameters. The resulting dispersion in the obtained values is considered as the associated uncertainties and is shown in Table 2.

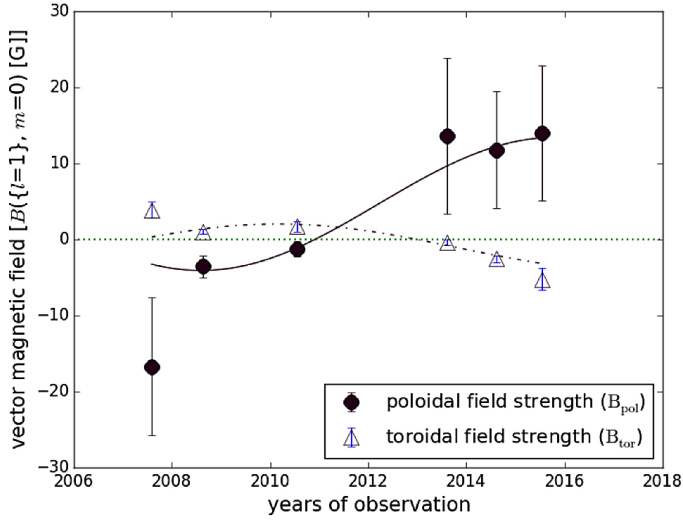


Fig. 11. Magnetic field strength B at the magnetic pole for the axisymmetric $l = 1$ mode for the poloidal component, which is similar to $l = 1$ in Fig. 10. For the toroidal component, B represents the strength of the latitudinal or azimuthal field. The green horizontal dashed line represents $B = 0$. The fits are obtained with a magnetic cycle period of 14.4 yr which is twice the activity cycle period of 7.2 yr.

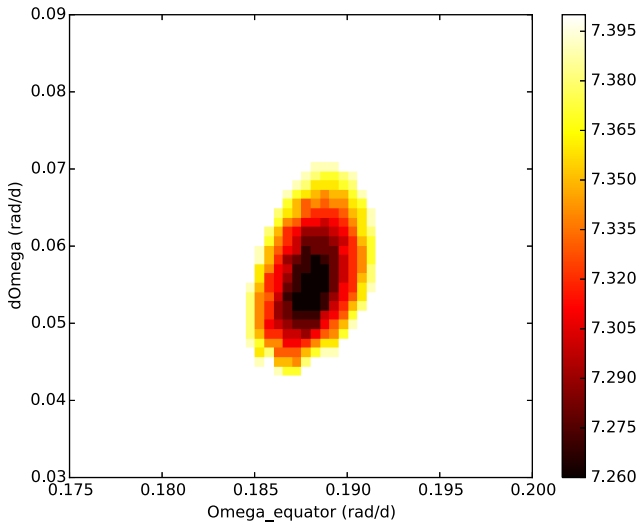


Fig. 12. Best fit B_{mean} map obtained by varying the differential parameters for epoch 2010.55. The Ω_{eq} and $d\Omega$ values obtained from this map are 0.184 rad d^{-1} and 0.04 rad d^{-1} respectively.

To obtain reliable differential rotation parameters, it is necessary to have observations with good phase coverage spread over multiple rotation periods of the star (Petit et al. 2002; Morgenthaler et al. 2012). Consequently, the differential parameters for 61 Cyg A could only be obtained for epoch 2010.55, as shown in Table 2. For epoch 2010.55 we measured a $\Omega_{\text{eq}} = 0.18 \text{ rad d}^{-1}$ and $d\Omega = 0.04 \text{ rad d}^{-1}$. We obtained a rotation period P_{rot} of 34.2 ± 3.7 days. This rotation period is in agreement with the rotation period obtained from chromospheric measurements, as shown in Table 1. Using this $d\Omega$ and Ω value we determine the relative horizontal shear (α) of 0.2 which is equivalent to the $\alpha_{\odot} = 0.2$.

Since the phase coverage of the other epochs was sparse, we could not obtain differential rotation measurements as reliable as those in epoch 2010.55. The magnetic field geometry is also

simple with single polarity dominating the surface in all epochs except 2010.55 making it harder to determine the differential rotation for those epochs. Hence, we adopted the differential rotation parameters of 2010.55 for the rest of the observational epochs.

6. Long-term evolution of the magnetic and activity cycle

The long-term evolution of the chromospheric indicators and the magnetic field strength is shown in Fig. 13. The average value of the different chromospheric activity indicators exhibits good correlation over a long timescale. Figure 13 also shows the mean longitudinal magnetic field of 61 Cyg A averaged over each epoch, where the error bars represent the standard deviation in B_l per observational epoch. Strong temporal variations are detected for the average longitudinal field, with stronger field strength in epochs 2007.59, 2013.61, 2014.61 and 2015.54. The field strength is the weakest in epochs 2008.64 and 2010.55. However, we note that epochs 2007.59 and 2010.55 exhibit stronger dispersion in B_l compared to the other epochs. Long-term observations of 61 Cyg A exhibits a weak anti-correlation between the mean magnetic field strength and the different chromospheric activity indicators.

The mean magnetic field obtained from the reconstructed ZDI maps averaged over each epoch is also shown in Fig. 13. B_{mean} follows a similar trend to B_l ; however, B_{mean} is weakest in epoch 2008.64 and does not show any pronounced anti-correlation with the S-index.

7. Discussion

The large-scale magnetic field geometry of the solar-type K5 dwarf 61 Cyg A was reconstructed using ZDI for five observational epoch spread over nine years, covering a complete chromospheric activity cycle. It exhibits clear polarity reversals of its large-scale magnetic field geometry.

7.1. Large-scale magnetic field

7.1.1. Polarity reversal of the large-scale field

We report polarity reversals of the large-scale magnetic field of 61 Cyg A, where, for the first time, a strong correlation is detected between the magnetic cycle and the chromospheric activity cycle. Polarity reversals have been previously detected in the large-scale magnetic field reconstructions of several cool stars such as τ Boo, HD 190771, and HD 78366. However, the polarity reversals detected in previous targets do not exhibit any correlation with the star's chromospheric activity cycle.

The $v \sin i$ of 61 Cyg A is lower than other ZDI targets. A previous numerical study by Kochukhov & Piskunov (2002) has shown that although reconstructions of the surface geometry from spectropolarimetric data can be carried out for slow rotators, the resolution of the surface reconstruction decreases with lower $v \sin i$. Since 61 Cyg A is a slow rotator, we can only resolve down the surface elements into the lower order spherical harmonics modes. In spite of this shortcoming, the polarity reversal of 61 Cyg A is clearly seen in the Stokes V spectra itself. Figure 6 clearly demonstrates that the Stokes V profiles change with rotational phase and also change with the chromospheric activity cycle. As shown in Fig. 6, the form of the Stokes V profile in epoch 2007.59 flips in epoch 2013.61, indicating that the

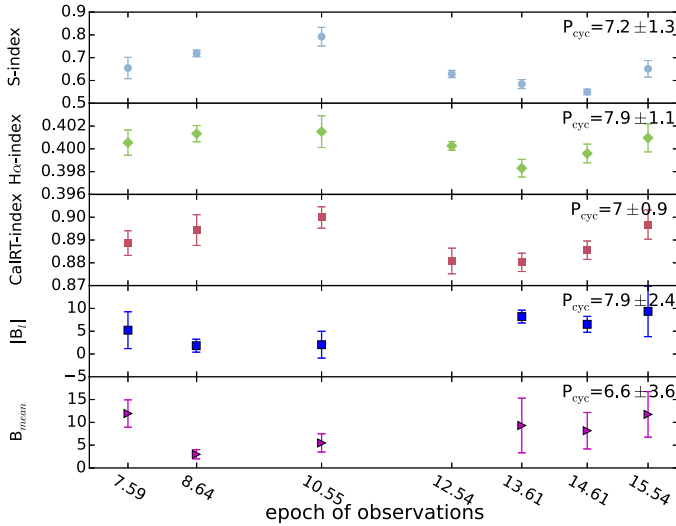


Fig. 13. Long-term evolution of the average values of the activity indices, the average longitudinal field, and the mean magnetic field measured from the ZDI maps.

large-scale field also flips polarity. The lack of any dominant Stokes V shape in epoch 2010.55 also indicates that the large-scale field is more complex during that epoch. The time series of Stokes V profiles reconfirms that the polarity reversals of the large-scale field are in fact independent of any artefact.

7.1.2. Evolution of the multipolar modes of the large-scale field

The magnetic field of 61 Cyg A's large-scale field is strongly poloidal at all epochs of our observations, which we quantify through energy fractions ($\langle B^2 \rangle$). Particularly at epoch 2013.61, which is the epoch around which the field geometry flips polarity, the poloidal magnetic field constitutes 99% of the total magnetic energy budget as seen in Stokes V . During epoch 2010.55 which is coincidentally the chromospheric activity maximum, the percentage of the toroidal magnetic field increases in comparison to other epochs. The axisymmetry of the poloidal field component also reaches its minimum in epoch 2010.55, as shown in Fig. 8. It indicates that the field is complex during activity maximum, in a similar way to the Sun. The poloidal field is dominated by the dipolar component except during activity maximum, when the other higher order modes dominate.

We investigate the lower order spherical harmonics modes of the star and interpret these results in the context of the Sun. Throughout our observational time span, the vector magnetic field, averaged over the surface, is stronger in the poloidal component except during epoch 2010.55. During activity maximum the vector component of the both the poloidal and toroidal field averaged over the surface of the star is close to zero. During epochs of minimum activity, we detected a weak anti-correlation between the poloidal and toroidal field.

The low order axisymmetric modes of the poloidal field indicate that the dipolar component is strongest compared to the quadrupolar and octopolar modes throughout our observational time span. The poloidal dipolar component ($l = 1$) is strongest during epochs of low activity and is almost negligible during activity maximum in epoch 2010.55. The other higher order modes are almost negligible throughout our observations. The large-scale solar magnetic field of the Sun is also more dipolar during epochs of low activity (DeRosa et al. 2012). The quadrupolar

component of the solar large-scale field dominates during activity maximum which is not seen in 61 Cyg A. One explanation for this disparity between the modes of 61 Cyg A might be caused by cross-talk between the odd and even sets of modes. To investigate that we suppressed the even modes and measured the field coefficients again. Although the odd modes are favoured over the even modes, no improvements were detected. An alternative explanation might be that the ZDI reconstruction technique has different sensitivities to the different spherical harmonics.

7.2. Chromospheric activity

The chromospheric activity of 61 Cyg A was determined by using three activity proxies: the Ca II H&K, $H\alpha$, and Ca II IRT lines. From the combined S-index data, we measure a cycle period of 7.2 ± 1.3 yr. An activity cycle period of 7.3 ± 0.1 yr was also previously determined by the long-term Mount Wilson survey (Baliunas et al. 1995).

A correlation is detected between the S-index and $H\alpha$ -index of 61 Cyg A as shown in Fig. 4. Such correlation is also observed in the Sun, where the correlation between the two indices follows the activity cycle (Livingston et al. 2007). Correlations were also observed for other BCoolest targets such as ξ Boo A (Morgenthaler et al. 2012) and HN Peg (Boro Saikia et al. 2015). The correlation between the fluxes in Ca II H&K lines and the $H\alpha$ lines was investigated by Cincunegui et al. (2007) for 109 cool stars. They observed strong dispersion in correlations with both positive and negative correlations. They also suggested that the correlation does not depend on activity and spectral type but on stellar colour.

The correlation between a star's S-index and $H\alpha$ -index has been further studied by Meunier & Delfosse (2009), who suggest that one of the main factors affecting the correlation of these two indices is the filling factor of plages and filaments on the stellar surfaces. Since plages affect both Ca II H&K and $H\alpha$ lines, and filaments are mostly observed as absorption in $H\alpha$, they claim that the correlations depended strongly on the spatial and temporal distribution of filaments on the stellar surface. They also observe an activity dependence on the correlations, where stronger correlation was observed near the cycle maxima. They also suggest that the correlations might also depend on time span of the observations and stellar inclination.

A correlation was also observed between 61 Cyg A's S-index and IRT-index and between $H\alpha$ -index and IRT-index (both shown in Fig. 4). A stronger correlation is observed between the S-index and the IRT-index. This is not surprising as the Ca II IRT lines also arise from plage dominated areas on the stellar surface (Andretta et al. 2005; Busà et al. 2007). In the BCoolest snapshot study Marsden et al. (2014) also reports a good correlation between these two indices.

A weak anti-correlation was observed between the S-index and the longitudinal magnetic field. In previous ZDI targets no correlation was observed between the chromospheric activity and longitudinal magnetic field B_l , for example, in ξ Boo A (Morgenthaler et al. 2012) and HN Peg (Boro Saikia et al. 2015). The lack of correlation is expected due to the cancellation of small scale magnetic features in the mean field measurements. However, for 61 Cyg A the longitudinal magnetic field B_l is weaker during activity maximum and stronger during activity minimum. No anti-correlation was detected, however, between the chromospheric S-index and the mean magnetic field (B_{mean}) measured from the magnetic maps. The chromospheric activity of 61 Cyg A shows good agreement with its coronal activity cycle which is also seen in the Sun.

8. Summary

In this paper we reconstruct the large-scale surface geometry of 61 Cyg A for six epochs over nine years of observations. We report the presence of a possible magnetic cycle which is twice the length of the activity cycle. This is the first detection of a cool star, apart from the Sun, where the magnetic cycle is in phase with its chromospheric cycle. The large-scale surface magnetic field geometry of 61 Cyg A flips its polarity and is highly variable, with the radial field showing strong polar field during activity minimum and a complex field during activity maximum. Throughout our observations, the large-scale field of 61 Cyg A is strongly poloidal except during activity minimum when the poloidal field is relatively weaker. The dipolar component of the poloidal field is strongest close to activity minimum and is weakest during activity maximum. During activity maximum, higher order modes such as quadrupolar and octopolar modes dominate over the dipolar mode. The evolution of the large-scale field of 61 Cyg A over the activity cycle shows close resemblance to the solar large-scale field, which has never been seen before in cool stars.

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Appendix A: Additional tables**Table A.1.** Journal of observations for seven epochs (2007.59, 2008.64, 2010.55, 2012.54, 2013.61, 2014.61, and 2015.54).

Date	Julian date (2 450 000+)	Exposure time (s)	S/N	σ_{LSD} ($10^{-5} I_c$)
2007 July 26	4308.49809	1200	35 765	2.7960
2007 July 30	4312.52989	800	35 807	2.7928
2007 July 31	4313.53121	800	27 096	3.6905
2007 August 02	4315.53608	800	33 873	2.9522
2007 August 03	4316.53315	800	35 102	2.8489
2007 August 04	4317.53462	800	30 572	3.2710
2007 August 08	4321.50738	800	31 550	3.1696
2007 August 09	4322.52760	800	28 744	3.4790
2007 August 10	4323.52773	800	24 362	4.1048
2007 August 17	4330.50089	1600	48 802	2.0491
2007 August 18	4331.46147	800	35 561	2.8121
2008 August 09	4688.54994	1200	46 210	2.1641
2008 August 12	4691.48653	1200	40 079	2.4951
2008 August 17	4696.52126	1200	31 218	3.2033
2008 August 21	4700.47785	1200	36 751	2.7210
2008 August 22	4701.48433	1200	21 527	4.6454
2008 August 23	4702.50793	1200	50 016	1.999
2008 August 24	4703.47966	1200	49 602	2.0161
2008 August 25	4704.53225	1200	48 263	2.0720
2008 August 26	4705.52953	1200	49 904	2.0039
2010 June 03	5351.64786	900	32 082	3.1171
2010 June 21	5369.57210	900	31 021	3.2236
2010 July 01	5379.58460	900	27 037	3.6987
2010 July 12	5390.54209	900	25 343	3.9458
2010 July 13	5391.52388	900	26 541	3.7678
2010 July 14	5392.54747	900	34 955	2.8608
2010 July 15	5393.63958	900	31 139	3.2114
2010 July 18	5396.60809	900	38 824	2.5757
2010 July 23	5401.65920	900	37 562	2.6623
2010 July 24	5402.49800	900	37 471	2.6687
2010 July 25	5403.49549	900	37 945	2.6354
2010 August 02	5411.49161	900	35 432	2.8224
2010 August 03	5412.56953	900	40 189	2.4883
2010 August 06	5415.55556	900	32 469	3.0799
2010 August 07	5416.54320	900	36 301	2.7548
2010 August 10	5419.56875	900	36 220	2.7610
2012 July 08	6117.63676	900	33 129	..
2012 July 09	6118.63476	900	28 811	..
2012 July 15	6124.64113	900	35 051	..
2012 July 16	6125.61827	900	36 914	..
2012 July 17	6126.61222	900	34 327	..
2012 July 18	6127.56778	900	34 221	..
2012 July 19	6128.55829	900	32 787	..
2012 July 22	6131.55361	900	34 150	..
2012 July 23	6132.54263	900	35 066	..
2012 July 24	6133.56659	900	34 955	..
2013 July 14	6488.51321	800	15 411	6.4889
2013 August 02	6507.65889	800	28 131	3.5548
2013 August 04	6509.51420	800	34 901	2.8653
2013 August 05	6510.61746	800	24 113	4.1472
2013 August 08	6513.57114	800	31 163	3.2090
2013 August 09	6514.54087	800	35 398	2.8251
2013 August 10	6515.54442	800	35 718	2.7997
2013 August 11	6516.56402	800	33 745	2.9634
2013 August 13	6518.58190	800	36 075	2.7720

Notes. Column 1 represents the year and date of observations, Col. 2 is the Julian date, col. 3 is the exposure time, Col. 4 is the signal-to-noise ratio of each Stokes V LSD profile, and Col. 5 represents the error bars in the Stokes V LSD profile.

Table A.1. continued.

Date	Julian date (2 450 000+)	Exposure time (s)	S/N	σ_{LSD} ($10^{-5} I_c$)
2013 August 14	6519.62914	800	35 479	2.8186
2013 August 18	6523.51294	800	33 442	2.9903
2013 August 19	6524.49914	800	34 439	2.9036
2013 August 20	6525.44457	800	33 347	2.9988
2014 July 14	6853.55493	800	21 642	4.6207
2014 July 17	6856.55743	800	28 094	3.5596
2014 July 22	6861.57410	800	34 595	2.8906
2014 July 25	6864.55546	800	34 374	2.9092
2014 July 26	6865.53839	800	35 150	2.8449
2014 July 27	6866.58350	800	35 645	2.8055
2014 August 11	6881.38890	800	27 850	3.5907
2014 August 19	6889.36216	800	18 535	5.3954
2014 August 23	6893.46397	800	35 628	2.8068
2014 August 25	6895.46092	800	29 424	3.3987
2014 August 27	6897.47963	800	36 306	2.7544
2014 August 29	6899.49512	800	35 859	2.7888
2014 August 30	6900.52575	800	31 814	3.1433
2014 August 31	6901.48870	800	36 408	2.7467
2015 June 26	7200.57599	900	26 403	3.7874
2015 June 28	7202.53874	900	36 708	2.7242
2015 June 29	7203.59939	900	30 539	3.2745
2015 June 30	7204.56304	900	30 669	3.2607
2015 July 07	7211.52277	900	31 832	3.1415
2015 July 08	7212.52460	900	34 962	2.8603
2015 July 09	7213.50574	900	37 205	2.6878
2015 July 10	7214.52800	900	33 875	2.9521
2015 July 11	7215.56071	900	32 176	3.1080
2015 July 12	7216.54418	900	34 501	2.8985
2015 July 14	7218.55681	900	32 888	3.0407
2015 July 15	7219.56343	900	23 679	4.2233
2015 July 20	7224.56050	900	25 336	3.9469
2015 August 05	7240.48208	900	24 468	4.0871
2015 August 10	7245.54843	900	22 931	4.3610
2015 August 11	7246.57034	900	31 250	3.2000

Table A.2. The chromospheric activity and longitudinal field measurements of 61 Cyg A for six epochs (2007.59, 2008.64, 2010.55, 2012.54, 2013.61, 2014.61, and 2015.54).

Epoch	Julian date (2 454 000+)	Rot. phase	S-index	H α -index	CaIRT-index	B_l (G)	N_l (G)
2007.59	4308.49809	0.00000	0.69 ± 0.02	0.402 ± 0.001	0.889 ± 0.003	-10.0 ± 0.9	-0.8 ± 0.9
	4312.52989	0.11789	0.69 ± 0.02	0.402 ± 0.001	0.895 ± 0.003	-8.5 ± 0.9	-0.0 ± 0.9
	4313.53121	0.14717	0.69 ± 0.01	0.401 ± 0.001	0.894 ± 0.004	-8.8 ± 1.1	-1.1 ± 1.1
	4315.53608	0.20579	0.70 ± 0.02	0.401 ± 0.001	0.891 ± 0.003	-4.7 ± 0.9	0.1 ± 0.9
	4316.53315	0.23494	0.68 ± 0.02	0.402 ± 0.001	0.892 ± 0.003	-5.4 ± 0.9	0.1 ± 0.9
	4317.53462	0.26423	0.71 ± 0.01	0.400 ± 0.001	0.889 ± 0.004	-4.3 ± 1.0	0.6 ± 1.0
	4321.50738	0.38039	0.66 ± 0.02	0.399 ± 0.001	0.888 ± 0.004	0.0 ± 1.0	-0.2 ± 1.0
	4322.52760	0.41022	0.62 ± 0.02	0.399 ± 0.001	0.889 ± 0.004	1.1 ± 1.1	-0.6 ± 1.1
	4323.52773	0.43946	0.61 ± 0.01	0.400 ± 0.001	0.887 ± 0.005	1.7 ± 1.3	0.6 ± 1.3
	4330.50089	0.64336	0.60 ± 0.02	0.399 ± 0.001	0.876 ± 0.002	-8.6 ± 0.6	-0.4 ± 0.6
	4331.46147	0.67144	0.58 ± 0.01	0.401 ± 0.001	0.886 ± 0.003	-9.1 ± 0.9	0.4 ± 0.9
2008.64	4688.54994	0.11263	0.73 ± 0.01	0.401 ± 0.001	0.898 ± 0.003	-2.9 ± 0.7	0.3 ± 0.7
	4691.48653	0.19849	0.73 ± 0.02	0.401 ± 0.001	0.899 ± 0.003	-1.4 ± 0.8	-0.9 ± 0.8
	4696.52126	0.34571	0.71 ± 0.03	0.402 ± 0.001	0.900 ± 0.004	1.0 ± 1.0	0.4 ± 1.0
	4700.47785	0.46140	0.71 ± 0.02	0.400 ± 0.001	0.896 ± 0.004	-1.4 ± 0.8	0.2 ± 0.8
	4701.48433	0.49083	0.69 ± 0.04	0.402 ± 0.002	0.905 ± 0.006	-4.0 ± 1.5	-0.3 ± 1.5
	4702.50793	0.52076	0.72 ± 0.01	0.400 ± 0.001	0.889 ± 0.003	-0.7 ± 0.6	0.4 ± 0.6
	4703.47966	0.54917	0.72 ± 0.01	0.401 ± 0.001	0.887 ± 0.003	-1.0 ± 0.6	1.1 ± 0.6
	4704.53225	0.57995	0.73 ± 0.01	0.401 ± 0.001	0.890 ± 0.003	-3.1 ± 0.6	-0.1 ± 0.6
	4705.52953	0.60911	0.73 ± 0.02	0.401 ± 0.001	0.887 ± 0.003	-2.0 ± 0.6	0.2 ± 0.6
2010.55	5351.64786	0.50146	0.65 ± 0.03	0.400 ± 0.001	0.892 ± 0.004	4.1 ± 1.0	-0.8 ± 1.0
	5369.57210	0.02556	0.85 ± 0.03	0.402 ± 0.001	0.909 ± 0.004	0.3 ± 1.0	-0.6 ± 1.0
	5379.58460	0.31832	0.79 ± 0.04	0.402 ± 0.001	0.901 ± 0.004	-2.1 ± 1.1	-1.9 ± 1.1
	5390.54209	0.63871	0.74 ± 0.04	0.400 ± 0.001	0.899 ± 0.005	2.3 ± 1.2	1.4 ± 1.2
	5391.52388	0.66742	0.73 ± 0.03	0.399 ± 0.001	0.894 ± 0.005	2.1 ± 1.2	1.6 ± 1.2
	5392.54747	0.69735	0.76 ± 0.02	0.400 ± 0.001	0.893 ± 0.004	0.2 ± 0.9	1.0 ± 0.9
	5393.63958	0.72928	0.84 ± 0.04	0.405 ± 0.001	0.900 ± 0.004	3.0 ± 1.1	-0.8 ± 1.1
	5396.60809	0.81608	0.79 ± 0.03	0.401 ± 0.001	0.901 ± 0.003	-0.2 ± 0.9	-0.7 ± 0.9
	5401.65920	0.96378	0.81 ± 0.03	0.401 ± 0.001	0.901 ± 0.004	-7.9 ± 0.8	0.2 ± 0.8
	5402.49800	0.98830	0.84 ± 0.03	0.401 ± 0.001	0.903 ± 0.004	-5.3 ± 0.8	-0.3 ± 0.8
	5403.49549	0.01747	0.85 ± 0.03	0.403 ± 0.001	0.905 ± 0.004	-4.2 ± 0.8	0.5 ± 0.8
	5411.49161	0.25127	0.81 ± 0.02	0.402 ± 0.001	0.904 ± 0.004	0.1 ± 0.9	-0.3 ± 0.9
	5412.56953	0.28279	0.79 ± 0.02	0.402 ± 0.001	0.898 ± 0.003	0.1 ± 1.0	-0.6 ± 1.0
	5415.55556	0.37010	0.77 ± 0.03	0.402 ± 0.001	0.902 ± 0.004	0.1 ± 1.0	-0.6 ± 1.0
	5416.54320	0.39897	0.77 ± 0.01	0.402 ± 0.001	0.895 ± 0.004	1.4 ± 0.9	-1.5 ± 0.9
	5419.56875	0.48745	0.74 ± 0.03	0.399 ± 0.001	0.892 ± 0.004	1.3 ± 0.9	0.1 ± 0.9
2012.54	6117.63676	0.89879	0.61 ± 0.02	0.399 ± 0.001	0.880 ± 0.004	..	..
	6118.63476	0.92797	0.61 ± 0.02	0.400 ± 0.001	0.874 ± 0.005	..	..
	6124.64113	0.10360	0.61 ± 0.02	0.400 ± 0.001	0.872 ± 0.004	..	..
	6125.61827	0.13217	0.62 ± 0.02	0.400 ± 0.001	0.880 ± 0.003	..	..
	6126.61222	0.16123	0.62 ± 0.02	0.401 ± 0.001	0.881 ± 0.004	..	..
	6127.56778	0.18917	0.62 ± 0.02	0.400 ± 0.001	0.884 ± 0.004	..	..
	6128.55829	0.21813	0.63 ± 0.02	0.400 ± 0.001	0.878 ± 0.004	..	..
	6131.55361	0.30572	0.65 ± 0.01	0.400 ± 0.001	0.885 ± 0.004	..	..
	6132.54263	0.33464	0.65 ± 0.02	0.400 ± 0.001	0.889 ± 0.004	..	..
	6133.56659	0.36458	0.65 ± 0.01	0.401 ± 0.001	0.885 ± 0.004	..	..

Notes. From left to right it represents: Julian date, rotational phase, S-index, H α -index, CaIRT-index, longitudinal magnetic field (B_l), and magnetic field of the Null profile (N_l). The longitudinal magnetic field measurements are not included for 2012.54 epoch due to spectropolarimetric errors at NARVAL. See Sect. 3 for details.

Table A.2. continued.

Epoch	Julian date (2 454 000+)	Rot. phase	S-index	H α -index	CaIRT-index	B_l (G)	N_l (G)
2013.61	6488.51321	0.74313	0.62 ± 0.06	0.399 ± 0.002	0.889 ± 0.006	3.4 ± 2.0	-2.7 ± 2.1
	6507.65889	0.30295	0.63 ± 0.04	0.398 ± 0.001	0.884 ± 0.004	6.2 ± 1.2	0.5 ± 1.2
	6509.51420	0.35720	0.60 ± 0.02	0.400 ± 0.001	0.881 ± 0.004	11.3 ± 0.9	0.3 ± 0.9
	6510.61746	0.38946	0.61 ± 0.03	0.398 ± 0.001	0.890 ± 0.005	10.5 ± 1.3	-0.2 ± 1.3
	6513.57114	0.47582	0.59 ± 0.02	0.398 ± 0.001	0.883 ± 0.004	9.0 ± 1.0	-0.1 ± 1.0
	6514.54087	0.50417	0.59 ± 0.02	0.398 ± 0.001	0.883 ± 0.004	7.4 ± 0.9	-0.0 ± 0.9
	6515.54442	0.53352	0.58 ± 0.02	0.397 ± 0.001	0.876 ± 0.004	7.7 ± 0.9	-0.1 ± 0.9
	6516.56402	0.56333	0.57 ± 0.01	0.398 ± 0.001	0.876 ± 0.004	7.9 ± 0.9	0.7 ± 0.9
	6518.58190	0.62233	0.57 ± 0.02	0.398 ± 0.001	0.877 ± 0.004	8.6 ± 0.9	0.7 ± 0.9
	6519.62914	0.65295	0.57 ± 0.02	0.398 ± 0.001	0.878 ± 0.004	7.9 ± 0.9	-0.2 ± 0.9
	6523.51294	0.76652	0.56 ± 0.01	0.398 ± 0.001	0.877 ± 0.004	7.9 ± 1.0	-0.1 ± 1.0
	6524.49914	0.79535	0.57 ± 0.02	0.399 ± 0.001	0.879 ± 0.004	7.4 ± 0.9	-1.4 ± 0.9
2014.61	6525.44457	0.82300	0.57 ± 0.02	0.398 ± 0.001	0.878 ± 0.004	6.7 ± 1.0	-1.1 ± 1.0
	6853.55493	0.41687	0.56 ± 0.04	0.399 ± 0.001	0.892 ± 0.005	6.8 ± 1.5	-1.6 ± 1.5
	6856.55743	0.50466	0.55 ± 0.03	0.398 ± 0.001	0.894 ± 0.004	5.9 ± 1.1	-1.4 ± 1.1
	6861.57410	0.65135	0.55 ± 0.02	0.400 ± 0.001	0.886 ± 0.004	8.0 ± 0.9	0.4 ± 0.9
	6864.55546	0.73852	0.56 ± 0.02	0.400 ± 0.001	0.885 ± 0.004	5.0 ± 0.9	-0.7 ± 0.9
	6865.53839	0.76726	0.56 ± 0.02	0.401 ± 0.001	0.885 ± 0.003	5.3 ± 0.9	-1.2 ± 0.9
	6866.58350	0.79782	0.56 ± 0.01	0.400 ± 0.001	0.887 ± 0.003	3.6 ± 0.9	1.1 ± 0.9
	6881.38890	0.23073	0.58 ± 0.02	0.399 ± 0.001	0.889 ± 0.004	4.6 ± 1.1	1.7 ± 1.1
	6889.36216	0.46386	0.55 ± 0.03	0.401 ± 0.002	0.890 ± 0.006	4.1 ± 1.7	0.0 ± 1.7
	6893.46397	0.58380	0.54 ± 0.02	0.400 ± 0.001	0.882 ± 0.003	8.6 ± 0.9	1.2 ± 0.9
	6895.46092	0.64219	0.54 ± 0.02	0.399 ± 0.001	0.887 ± 0.004	7.5 ± 1.1	0.9 ± 1.1
	6897.47963	0.70121	0.54 ± 0.02	0.398 ± 0.001	0.878 ± 0.003	8.6 ± 0.9	-0.3 ± 0.9
2015.54	6899.49512	0.76015	0.54 ± 0.01	0.400 ± 0.001	0.883 ± 0.003	7.3 ± 0.9	0.7 ± 0.9
	6900.52575	0.79028	0.54 ± 0.01	0.399 ± 0.001	0.884 ± 0.004	8.4 ± 1.1	-0.9 ± 1.1
	6901.48870	0.81844	0.54 ± 0.01	0.400 ± 0.001	0.883 ± 0.003	7.7 ± 0.9	-1.4 ± 0.9
	7200.57599	0.56368	0.60 ± 0.03	0.400 ± 0.001	0.893 ± 0.004	-0.8 ± 1.2	0.3 ± 1.2
	7202.53874	0.62107	0.60 ± 0.02	0.401 ± 0.001	0.889 ± 0.003	-0.5 ± 0.9	-1.1 ± 0.9
	7203.59939	0.65208	0.59 ± 0.03	0.400 ± 0.001	0.886 ± 0.004	2.4 ± 1.0	1.0 ± 1.0
	7204.56304	0.68026	0.59 ± 0.03	0.401 ± 0.001	0.890 ± 0.004	2.3 ± 1.0	0.4 ± 1.0
	7211.52277	0.88376	0.64 ± 0.03	0.400 ± 0.001	0.896 ± 0.004	13.3 ± 1.0	-1.2 ± 1.0
	7212.52460	0.91306	0.71 ± 0.05	0.401 ± 0.001	0.890 ± 0.005	14.1 ± 0.9	-0.5 ± 0.9
	7213.50574	0.94174	0.67 ± 0.02	0.401 ± 0.001	0.897 ± 0.003	14.6 ± 0.9	0.3 ± 0.9
	7214.52800	0.97163	0.68 ± 0.05	0.401 ± 0.001	0.903 ± 0.005	13.7 ± 1.0	1.0 ± 1.0
	7215.56071	0.00183	0.64 ± 0.03	0.402 ± 0.001	0.898 ± 0.004	13.9 ± 1.0	-0.7 ± 1.0
2015.54	7216.54418	0.03059	0.67 ± 0.03	0.402 ± 0.001	0.900 ± 0.004	14.3 ± 0.9	-0.9 ± 0.9
	7218.55681	0.08943	0.69 ± 0.03	0.403 ± 0.001	0.906 ± 0.004	12.4 ± 1.0	0.5 ± 1.0
	7219.56343	0.11887	0.69 ± 0.05	0.403 ± 0.001	0.906 ± 0.005	12.0 ± 1.3	0.2 ± 1.3
	7224.56050	0.26498	0.68 ± 0.03	0.401 ± 0.001	0.904 ± 0.005	3.5 ± 1.3	-0.2 ± 1.2
	7240.48208	0.73053	0.62 ± 0.04	0.400 ± 0.001	0.897 ± 0.005	0.7 ± 1.3	1.1 ± 1.3
	7245.54843	0.87866	0.65 ± 0.04	0.400 ± 0.001	0.896 ± 0.005	9.9 ± 1.4	0.2 ± 1.4
	7246.57034	0.90855	0.63 ± 0.01	0.399 ± 0.001	0.892 ± 0.004	12.6 ± 1.0	0.3 ± 1.0