

## A STUDY OF OPTICAL OBSERVING TECHNIQUES FOR EXTRA-GALACTIC SUPERNOVA REMNANTS: CASE OF NGC 300

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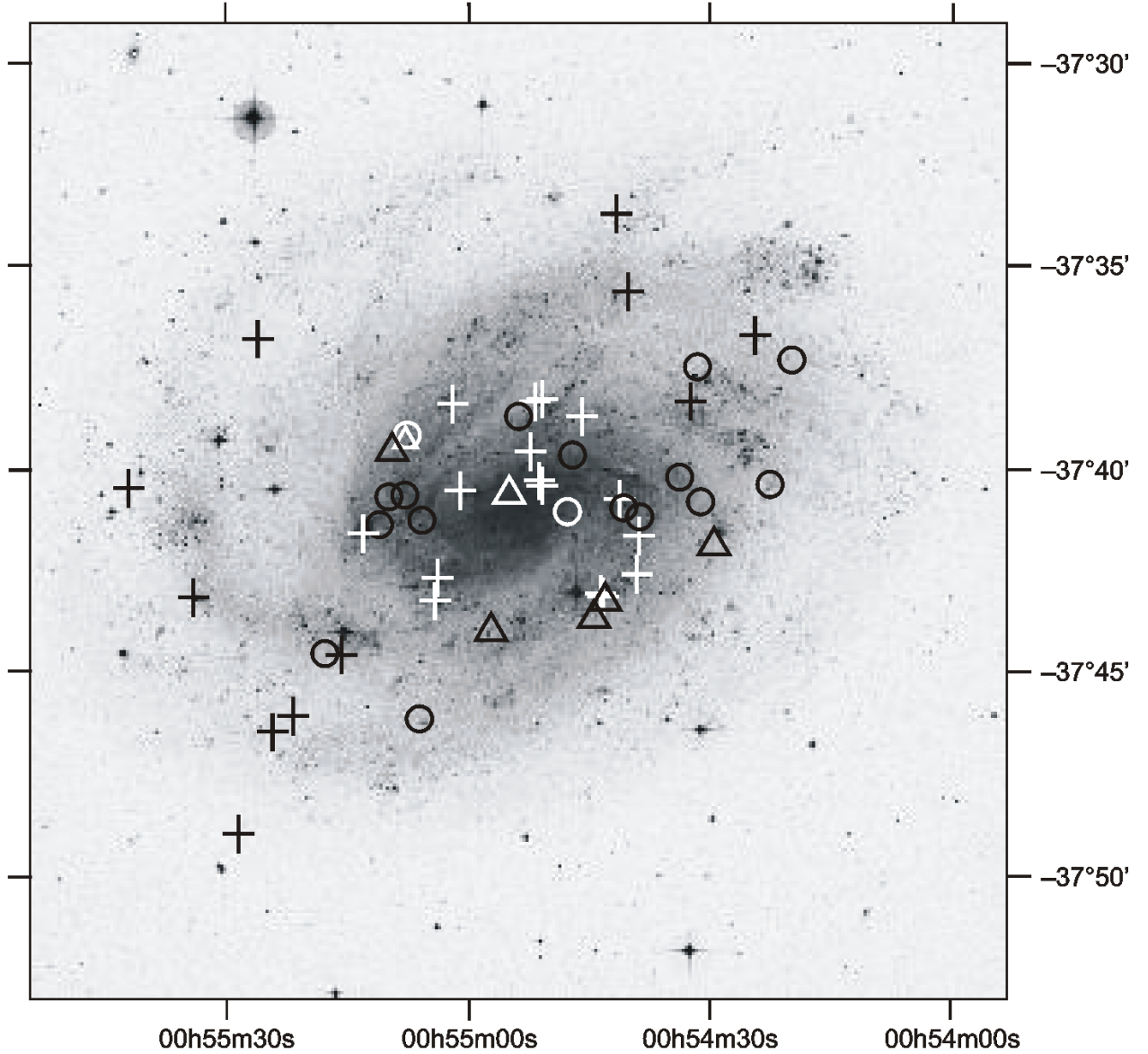
**SUMMARY:** We present the results of a study of observational and identification techniques used for surveys and spectroscopy of candidate supernova remnants (SNRs) in the Sculptor Group galaxy NGC 300. The goal of this study was to investigate the reliability of using  $[\text{S II}]:\text{H}\alpha \geq 0.4$  in optical SNR surveys and spectra as an identifying feature of extra-galactic SNRs (egSNRs), and also to investigate the effectiveness of the observing techniques (which are hampered by seeing conditions and telescope pointing errors) using this criterion in egSNR surveys and spectrographs. This study is based on original observations of these objects and archival data obtained from the Hubble Space Telescope which contained images of some of the candidate SNRs in NGC 300. We found that the reliability of spectral techniques may be questionable and very high-resolution images may be needed to confirm a valid identification of some egSNRs.

**Key words.** ISM: supernova remnants – galaxies: individual: NGC 300 – galaxies: ISM – astrometry – telescopes – techniques: miscellaneous

### 1. INTRODUCTION

The Earth is located in the gas and dust filled disk of the Milky Way galaxy (here after, the Galaxy) making the study of supernova remnants (SNRs) within the Galaxy difficult because the extinction and reddening effects of this interstellar medium (ISM) blocks or severely hampers our ability to see Galactic SNRs at wavelengths other than radio. For observations of SNRs in nearby galaxies this absorption by

the ISM within both the host galaxy and the Galaxy is greatly reduced – particularly for face-on (low inclination angle) spiral galaxies with high Galactic latitude (Matonick et al. 1997, Pannuti et al. 2000). Surveys of SNRs in the Local Group galaxies such as the Large and Small Magellanic Clouds (Filipović et al. 2005, Payne et al. 2008) and galaxies within some nearby clusters (mostly the Sculptor Group) have resulted in observations that are free from these biases. Over 450 SNRs have been found in nearby galaxies



**Fig. 1.** A DSS image of NGC 300 with the positions (in J2000.0 coordinates) indicated of the 51 SNRs and candidate SNRs considered by the present study. Radio sources (SNRs and SNR candidates only) from P04 are shown with crosses. Optical candidates with line ratios measured with long-slit spectra (from BL97) are shown as circles and optical candidates with line ratios measured by interference filters (BL97) are shown with triangles. Symbols are black or white only for increased contrast. (Figure originally published in MWF11. Southern sky DSS image, courtesy of Royal Observatory Edinburgh, Anglo-Australian Observatory, California Institute of Technology.)

and listings are given by Matonick et al. (1997), Urošević et al. (2005), Filipović et al. (2008) and Pannuti et al. (2007).

An image of NGC 300 is shown in Fig. 1. Table 1 provides a brief list of the characteristics of NGC 300 following Millar et al. (2011, hereafter, MWF11); a more complete list is provided by Kim et al. (2004). NGC 300 has a low inclination angle and a high Galactic latitude and observations of

SNRs within NGC 300 suffer very low internal extinction (Butler et al. 2004) and foreground reddening<sup>1</sup>. NGC 300 shows many giant H II regions which are evidence of many star formation episodes (Read et al. 2001). The similarities between NGC 300 and other nearby spiral galaxies (such as M33 in the Local Group, and other members of the Sculptor Group, such as NGC 7793) lead us to conclude that NGC 300 is a typical, normal spiral galaxy (Blair et al. 1997,

<sup>1</sup> $E(B - V) = 0.013$  mag (Bland-Hawthorn et al. 2005).

**Table 1.** Gross Properties of NGC 300 (from MWF11).

Property	Value	Reference
Hubble Type	SA(s)d	Tully et al. (1988) deVaucouleurs et al. (1991)
R.A. (J2000.0)	00 <sup>h</sup> 54 <sup>m</sup> 53.48 <sup>s</sup>	NED
Dec. (J2000.0)	−37°41′03.8″	NED
Galactic Latitude	−77.17°	NED
Radial Velocity	144 km/s (Solar)	Puche et al. (1990) Karachentsev et al. (2003)
Inclination	46° 42.6°	Tully et al. (1988) Puche et al. (1990)
Distance	2.1 Mpc 2.02 Mpc 1.88 Mpc	Freedman et al. (1992) Freedman et al. (2001) Bresolin et al. (2005); Gieren et al. (2005)
Observed Diameter ( $D_{25}$ )	20.2 arcmin	Tully et al. (1988)
Observed Diameter (UV isophotes)	21.9 × 15.5 arcmin	Gil de Paz et al. (2007)
Galaxy Diameter	22.6 kpc, at 2.1 Mpc	Based on Gil de Paz et al. (2007)
Mass ( H I )	$2.4 \times 10^9 M_{\odot}$	Tully et al. (1988)
$N_H$ Column Density	$2.97 \times 10^{20} \text{ cm}^{-2}$	Read et al. (1997)

Note. NED = NASA/IPAC Extragalactic Database (<http://nedwww.ipac.caltech.edu/>).

hereafter, BL97). Most authors have adopted distances of 2.0–2.1 Mpc for NGC 300 (BL97, Payne et al. 2004, hereafter P04; Table 1) hence, we have adopted this distance of 2.1 Mpc to be consistent with previous SNR surveys and spectral observations (BL97, P04). The corresponding linear scale is  $10.2 \text{ pc arcsec}^{-1}$ .

The resolving power (or seeing) of the two-meter class, ground-based telescopes usually used for SNR observations is generally from 0″.5 (Chil ) up to 2″ (Siding Spring, Australia or Sutherland, South Africa). These seeing conditions do not allow the direct imaging of extra-galactic SNRs (hereafter, egSNRs) and therefore, spectral line intensity ratios are used for identification of these objects. The now common technique of using two narrow band optical interference filters, one centered on H  (  6564  ) and the second allowing passage of the [ S II ] (   6717, 6731  ) doublet, to observe (or survey) bright nebulae in nearby galaxies was first described in a series of papers by Mathewson et al. (1972, 1973a,b,c). This technique depends on the strength of the [ S II ] lines in SNRs being about the same strength as the H  line, which is most likely due to shock fronts in the expanding SNR shell as it collides with a dense ISM. In H II regions (where there are few if any shock fronts) this condition would not exist. The total flux of the two [ S II ] lines should be at least an order of magnitude weaker than the H  line in H II regions as compared to SNRs (Mathewson et al. 1972). Because of the filter’s bandwidth, the H  filter was not able to remove the [ N II ] (   6548, 6583  ) doublet lines adjacent to the H  line. In some SNRs these lines (summed) can be as strong as the H  itself. Given this extra flux near the

H  line, a candidate SNR was claimed if the emission region contained a (non-thermal) radio source and the H  + [ N II ] to [ S II ] ratio was less than two (Mathewson et al. 1972).

D’Odorico et al. (1978) present arguments based on observations of SNRs and H II regions within the Galaxy and within the Large Magellanic Cloud to show how SNRs can be identified within M33 when [ S II ]:H  > 0.4. When [ S II ]:H  < 0.2 the nebula is considered an H II region (BL97). When [ S II ]:H  is between these two values the nature of the nebula may be unclear. Fesen et al. (1985) found that [ O I ] (   6300, 6364  ), [ O II ] (  3727  ) and [ O III ] (   4959, 5007  ) are often all simultaneously strong in SNRs and this can also be used to help differentiate SNRs from H II regions in cases where [ S II ]:H  is borderline.

The first observations of SNRs within NGC 300 using this technique were published by D’Odorico et al. (1980). The candidates studied for this paper are the same candidates studied in MWF11 which were selected from those published by BL97 (optical candidates) and P04 (radio candidates). The positions of these SNRs and SNR candidates and the results of previous observations are given in Table 2 and are shown in Fig. 1. For Table 2, column 1 is the optical designation as given in BL97. Column 2 is the radio designation as given in P04. Columns 3 and 4 are the right ascension (RA) and declination (Dec) coordinates. Column 5 is the [ S II ]:H  as reported in (or derived from) BL97. Column 6 is the [ S II ]:H  as reported in MWF11 (including the measurement error discussed in MWF11). Column 7 is the spectral index ( $\alpha$ ) as reported in P04 and column 8 is the measured diameter as reported in MWF11.

**Table 2.** Summary of observational results for the selected objects (from BL97, P04, and MWF11). The cut-in headers are from MWF11.

1	2	3	4	5	6	7	8
Optical	Radio	RA	Dec	[ S II ]:H $\alpha$	[ S II ]:H $\alpha$	$\alpha \pm \Delta\alpha$	Diameter
Object	Object	(h m s)	( $^{\circ}$ ' ")	(BL97)	(MWF11)	(P04)	(pc, MWF11)
SNRs (MWF11)							
N300-S1		00 54 19.21	-37 37 23.96	0.44	$0.46 \pm 0.29$		38
N300-S2		00 54 21.85	-37 40 27.11	0.49	$0.72 \pm 0.39$		69
N300-S4		00 54 30.62	-37 40 53.75	0.71	$0.93 \pm 0.04$		150
N300-S5		00 54 30.99	-37 37 33.96	0.46	$0.56 \pm 0.47$		
N300-S6	J005431-373825	00 54 31.91	-37 38 25.68	0.60	0.69	- <sup>a</sup>	44
N300-S7		00 54 33.17	-37 40 16.90	0.57	$0.57 \pm 0.41$		31
N300-S8		00 54 38.17	-37 41 14.88	0.61	$0.58 \pm 0.06$		49
N300-S9		00 54 40.20	-37 41 02.12	0.44	$0.53 \pm 0.23$		83
N300-S12		00 54 43.86	-37 43 39.08	0.52	$0.73 \pm 0.27$		22
N300-S13		00 54 46.60	-37 39 44.32	0.59	$0.91 \pm 0.05$		35
N300-S14		00 54 47.15	-37 41 07.63	0.63	$1.08 \pm 0.24$		41
N300-S15		00 54 53.32	-37 38 48.24	0.65	$0.57 \pm 0.39$		12
N300-S16		00 54 54.46	-37 40 35.46	0.70	$0.94 \pm 0.06$		57
N300-S17		00 54 56.68	-37 43 57.70	0.69	$0.96 \pm 0.15$		65
N300-S19		00 55 05.41	-37 41 21.04	0.53	$0.70 \pm 0.42$		30
N300-S20		00 55 05.68	-37 46 13.35	0.75	$0.79 \pm 0.11$		48
N300-S22		00 55 07.50	-37 40 43.20	0.27	$0.46 \pm 0.38$		75
N300-S24		00 55 09.48	-37 40 46.21	0.80	$0.64 \pm 0.13$		100
N300-S25		00 55 10.68	-37 41 27.13	0.64	$0.54 \pm 0.40$		80
N300-S26	J005515-374439	00 55 15.46	-37 44 39.11	0.57	$0.86 \pm 0.67$	-	31
N300-S27		00 55 17.54	-37 44 36.65	0.70	$0.64 \pm 0.48$		66
N300-S28	J005533-374314	00 55 33.76	-37 43 13.13	0.61	$0.45 \pm 0.15$	-	63
Other Objects ( H II Regions? MWF11)							
	J005438-374144	00 54 38.16	-37 41 44.2		$0.17 \pm 0.07$	$-0.8 \pm 0.2$	
	J005438-374240	00 54 38.49	-37 42 40.5		$0.25 \pm 0.02$	-	
	J005439-373543	00 54 39.61	-37 35 43.4		0.27	$-0.4 \pm 0.1$	
	J005441-373348	00 54 41.05	-37 33 48.9		0.36	-	
	J005442-374313	00 54 42.70	-37 43 13.3		$0.19 \pm 0.07$	$-0.9 \pm 0.3$	
	J005443-374311	00 54 43.11	-37 43 11.0		$0.18 \pm 0.09$	$-0.6 \pm 0.2$	
	J005445-373847	00 54 45.39	-37 38 47.1		$0.11 \pm 0.03$	$-0.3 \pm 0.1$	
	J005450-374030	00 54 50.28	-37 40 30.0		$0.32 \pm 0.12$	$-0.5 \pm 0.2$	
	J005450-373822	00 54 50.30	-37 38 22.4		$0.22 \pm 0.14$	$-0.2 \pm 0.2$	
	J005450-374022	00 54 50.73	-37 40 22.2		$0.38 \pm 0.31$	$-0.3 \pm 0.1$	130
	J005451-373826	00 54 51.16	-37 38 26.1		$0.26 \pm 0.16$	$-1.2 \pm 0.7$	
	J005451-373939	00 54 51.79	-37 39 39.6		$0.10 \pm 0.18$	$-0.1 \pm 0.2$	
	J005500-374037	00 55 00.58	-37 40 37.4		$0.25 \pm 0.09$	$-0.4 \pm 0.4$	
	J005501-373829	00 55 01.49	-37 38 29.9		$0.35 \pm 0.12$	$-0.9 \pm 0.1$	31
	J005503-374246	00 55 03.50	-37 42 46.0		$0.14 \pm 0.04$	$-0.4 \pm 0.1$	
	J005503-374320	00 55 03.66	-37 43 20.1		$0.15 \pm 0.08$	$-0.7 \pm 0.3$	
	J005512-374140	00 55 12.70	-37 41 40.3		$0.08 \pm 0.02$	$-0.7 \pm 0.1$	
N300-S3		00 54 28.86	-37 41 53.32	0.40	$0.24 \pm 0.31$		26
N300-S10	J005440-374049	00 54 40.87	-37 40 48.73	0.67	$0.35 \pm 0.15$	$-0.5 \pm 0.3$	63
N300-S11		00 54 42.54	-37 43 14.16	0.66	$0.30 \pm 0.12$		150
N300-S18		00 55 01.39	-37 39 18.17	0.53	$0.32 \pm 0.32$		69
N300-S21		00 55 07.15	-37 39 15.17	0.59	$0.37 \pm 0.30$		41
N300-S23		00 55 09.10	-37 39 32.61	0.64	$0.31 \pm 0.08$		43
No Signal (MWF11)							
	J005423-373648	00 54 23.84	-37 36 48.4			$-0.7 \pm 0.1$	
	J005521-374609	00 55 21.35	-37 46 09.6			$-1.0 \pm 0.3$	
	J005523-374632	00 55 23.95	-37 46 32.4			$-0.9 \pm 0.1$	
	J005525-373653	00 55 25.82	-37 36 53.8			$-1.0 \pm 0.1$	
	J005528-374903	00 55 28.25	-37 49 03.3			$-0.6 \pm 0.3$	
	J005541-374033	00 55 41.94	-37 40 33.5			-	

<sup>a</sup>Spectral index not determined.

In MWF11, these candidates were studied with long-slit spectral observations and the accepted criteria ( $[\text{S II}]:\text{H}\alpha \geq 0.4$  and the presence of  $[\text{O I}]$ ) were used giving a result of 22 objects as SNRs, with the remaining left as unknown (no signal) or unclear. For some of those objects which were unclear, the error in the  $[\text{S II}]:\text{H}\alpha$  measurement allows overlap with the  $[\text{S II}]:\text{H}\alpha \geq 0.4$  threshold. The radio sources which do not reach the threshold value to be labeled as optical SNRs may be optically obscured by the emissions from neighboring  $\text{H II}$  regions (P04). This possibility is further supported by our investigation of the radio observations' astrometry here (Section 3.1). There are other well established causes for the poor overlap of observed radio and optical SNR emissions (Duric 2000a,b).

In this paper we investigate the reliability of using  $[\text{S II}]:\text{H}\alpha \geq 0.4$  as a defining criterion for the detection and identification of egSNRs, and the effectiveness of the observing techniques used for egSNR surveys and spectrographs. We present a study of comparisons of the observations of these SNR candidates between BL97, P04, MWF11 and images containing the candidates found in the archives of the Hubble Space Telescope (HST). For reasons given below, this study included an analysis of the astrometry of the original CCD images from BL97, a discussion of the astrometry of the radio observations from P04, and an analysis of the pointing errors inherent to the two-meter class telescopes generally used for egSNR observations. Even with good seeing conditions, this telescope size class may be too small to ensure the reliability of using the  $[\text{S II}]:\text{H}\alpha$  ratio to identify egSNRs with complete confidence.

We found five sources which were comparable between the observation sets – four radio sources and one optical source. We found an apparent systematic error in the radio observation astrometry. Unfortunately, with only four radio sources to work with in the current HST data, testing this possible error any further requires more high-resolution (HST) optical images. Only one source was in both the BL97 and HST observations. However, this one source demonstrated a need for further investigation into the reliability of using the  $[\text{S II}]:\text{H}\alpha$  line ratio for the identification of egSNRs in optical surveys and spectral observations.

## 2. NGC 300 SNRs IN THE HST ARCHIVAL DATA

Table 3 shows the file names associated with the central wavelengths sorted by wavelength along with the original proposal identification and any publications based on the data. In this table, column 1 is the central wavelength of the HST filter, column 2 is the file name (all with .FITS extension) containing the image, column 3 is a list of publications associated

with the observation. The references are included to provide the original objectives of the observations of NGC 300 at these wavelengths and other publications which have used these observations (all taken from the MAST<sup>2</sup> web site).

Table 4 shows the bandwidths for each of the HST filters used for these images. These data were taken from the images' FITS file header<sup>3</sup>. Column 1 is the filter's central wavelength and column 2 is the filter bandwidth. This table also shows emission wavelengths with the associated ions or atoms that are found within the bandwidth of the filters (column 3). These ions or atoms are known to be within (Galactic) SNRs (Fesen et al. 1996).

A computer program was written to read the HST FITS file headers and based on each file's FITS world coordinate system (WCS), the following objects (from Table 2) were found in one or more of the images files: J005438–374144, J005440–374049, J005445–373847, J005450–374030, J005450–373822, J005450–374022, J005451–373826, J005451–373939, J005500–374037, J005501–373829, J005503–374246, J005503–374320, N300-S8, N300-S9, N300-S10, N300-S13, N300-S14, N300-S15, N300-S16, N300-S18.

Table 5 lists the filter central wavelengths and bandwidths containing spectral emissions which are diagnostic to the identification of SNRs (see discussion in Section 1). This table's columns are the same as those in Table 4. The emissions of most interest are:  $\text{H}\alpha$  ( $\lambda$  6565 Å),  $[\text{S II}]$  ( $\lambda$  6717 Å,  $\lambda$  6732 Å),  $[\text{O I}]$  ( $\lambda$  6300 Å), and  $[\text{O III}]$  ( $\lambda$  5007 Å). Only the files possibly containing these wavelengths were examined any further. As can be seen from Table 3, only one image file group contains data with  $\text{H}\alpha$  (u671370#r.drz, where # is one of the digits, 5 to 9). There were no files containing the  $[\text{S II}]$  doublet wavelengths. The images centered on  $\lambda = 5741$  Å and  $\lambda = 6001$  Å contain wavelengths of interest but none of these images contain objects from Table 2. Thus, only the  $\text{H}\alpha$  images were of use for this study. Only 5 of the 51 objects were found within the  $\text{H}\alpha$  images. All of the  $\text{H}\alpha$  images have the same field of view. The final results of the search are shown in Table 6. The columns for this table are the same as those for Table 2.

The HST  $\text{H}\alpha$  images are centered on (00<sup>h</sup>54<sup>m</sup>54<sup>s</sup>.54, −37°40'35".9). They are 400 second exposures taken in May of 2001 with the Wide-Field Planetary Camera 2 (WFPC2) instrument using its (default) PC1 aperture. The filter used in these images (F656N) was centered on  $\lambda = 6563.76$  Å with a RMS bandwidth of  $\Delta\lambda = 53.77$  Å. The bandpass characteristics of this filter are shown in Fig. 2. The data were originally collected as part of the HST proposal 8591 by Douglas Richstone (Richstone 2000). The file u6713709r.drz.fits appears to provide the cleanest image for analysis of the SNR candidates' environment.

<sup>2</sup><http://archive.stsci.edu/>

<sup>3</sup>They are also available on the HST WFPC2 instrument website. <http://www.stsci.edu/hst/wfpc2/documents/IHB-17.html>

**Table 3.** Archival HST files organized by filter center wavelength.

1	2	3	4
Wavelength (Å)	Files	Proposal ID	Associated Publications
2993	u8hhvp02m <sub>drz</sub>	9677	Wadadekar et al. (2006).
4318	j8d702010 <sub>drz</sub>	9492	Gliozzi et al. (2009); Gogarten et al. (2009b);
	j8d702mcq <sub>flt</sub>		Nantais et al. (2010); Gogarten et al. (2010);
	j8d702mj <sub>q</sub> <sub>flt</sub>		Gieren et al. (2004); Bresolin et al. (2005);
	j8d702mr <sub>q</sub> <sub>flt</sub>		Rizzi et al. (2007); Kudritzki et al. (2008);
	j8d703010 <sub>drz</sub>		Bond et al. (2009); Barth et al. (2009);
	j8d703vi <sub>q</sub> <sub>flt</sub>		Dalcanton et al. (2009); Rizzi et al. (2006);
	j8d703vp <sub>q</sub> <sub>flt</sub>		Tikhonov et al. (2005);
	j8d703vx <sub>q</sub> <sub>flt</sub>		Kuntz et al. (2010).
4747	j9rai1010 <sub>drz</sub>	10915	Berger et al. (2009); Gogarten et al. (2009a);
	j9ra1ine <sub>q</sub> <sub>flt</sub>		Holwerda et al. (2009); Lianou et al. (2009);
	j9rai1nf <sub>q</sub> <sub>flt</sub>		Nantais et al. (2010); Williams et al. (2010);
			Melbourne et al. (2010); Gogarten et al. (2010);
			de Mello et al. (2008); Girardi et al. (2008);
			Mould et al. (2008); Bond et al. (2009);
			Williams et al. (2009); Kornei et al. (2009);
			Williams et al. (2009).
5360	j8d702020 <sub>drz</sub>	9492	See $\lambda = 4318$ Å.
	j8d702md <sub>q</sub> <sub>flt</sub>		
	j8d702ml <sub>q</sub> <sub>flt</sub>		
	j8d702mt <sub>q</sub> <sub>flt</sub>		
	j8d703020 <sub>drz</sub>		
	j8d703vj <sub>q</sub> <sub>flt</sub>		
	j8d703vr <sub>q</sub> <sub>flt</sub>		
	j8d703vz <sub>q</sub> <sub>flt</sub>		
5484	u6713701m <sub>drz</sub>	8591	Graham et al. (2001); Schinnerer et al. (2003);
	u6713702r <sub>drz</sub>		Larsen (2004); Lauer et al. (2005);
	u6713703r <sub>drz</sub>		Rosolowsky et al. (2005); Lauer et al. (2007a);
	u6713704r <sub>drz</sub>		Lauer et al. (2007b); González Delgado et al. (2008);
			Dai et al. (2008); Siopis et al. (2009);
			Beifiori et al. (2009).
5741	o6j3bzy <sub>q</sub> <sub>flt</sub>	9285	Shaw et al. (2006).
5921	j9rai1020 <sub>drz</sub>	10915	See $\lambda = 4747$ Å.
	j9rai1nh <sub>q</sub> <sub>flt</sub>		
	j9rai1nj <sub>q</sub> <sub>flt</sub>		
6001	u8hhvp01m <sub>drz</sub>	9677	Wadadekar et al. (2006).
6001	u8ixhd01m <sub>drz</sub>	9676	Tikhonov et al. (2009); Van Dyk et al. (2003);
	u8ixhd02m <sub>drz</sub>		Larsen (2004); Smartt et al. (2004);
			Tikhonov et al. (2005a); Sugerman (2005);
			Milone et al. (2006); Wadadekar et al. (2006);
			Sugerman et al. (2006); Shaw et al. (2007);
			Maund et al. (2009); Guerrero et al. (2008).
6564 (H $\alpha$ )	u6713705r <sub>drz</sub>	8591	See $\lambda = 5484$ Å.
	u6713706r <sub>drz</sub>		
	u6713707r <sub>drz</sub>		
	u6713708r <sub>drz</sub>		
	u6713709r <sub>drz</sub>		
7996	u65w0201r <sub>drz</sub>	8599	Windhorst et al. (2002); Böker et al. (2003a);
	u65w0202r <sub>drz</sub>		Larsen (2004); Butler et al. (2004)
	u65w0203r <sub>drz</sub>		Walcher et al. (2005); de Grijs et al. (2005);
			Tully et al. (2006); Rossa et al. (2006);
			Schinnerer et al. (2006); Peebles et al. (2006);
			Ganda et al. (2006); Cao et al. (2007);
			Seth et al. (2008); Böker et al. (2003b);
			González Delgado et al. (2008); Ghosh et al. (2009);
			Böker et al. (2004); Andersen et al. (2008);
			Böker et al. (2002).
8057	j9rai1030 <sub>drz</sub>	10915	See $\lambda = 4747$ Å.
	j9rai1nl <sub>q</sub> <sub>flt</sub>		
	j9rai1nn <sub>q</sub> <sub>flt</sub>		
8060	j8d702030 <sub>drz</sub>	9492	See $\lambda = 4318$ Å.
	j8d702mf <sub>q</sub> <sub>flt</sub>		
	j8d702mo <sub>q</sub> <sub>flt</sub>		
	j8d702mv <sub>q</sub> <sub>flt</sub>		
	j8d702n0 <sub>q</sub> <sub>drz</sub>		
	j8d702n0 <sub>q</sub> <sub>flt</sub>		
	j8d703030 <sub>drz</sub>		
	j8d703vl <sub>q</sub> <sub>flt</sub>		
	j8d703vu <sub>q</sub> <sub>flt</sub>		
	j8d703w1 <sub>q</sub> <sub>flt</sub>		
	j8d703w6 <sub>q</sub> <sub>drz</sub>		
	j8d703w6 <sub>q</sub> <sub>flt</sub>		

**Table 4.** Central Wavelengths of HST Image File Sequences.

1	2	3
Wavelength (Å)	RMS Bandwidth (Å)	Known SNR Ions/Atoms with Wavelengths Å, Fesen et al. 1996
2993	325	Mg I : 2852.13
4318	293	H I : 4340.49 (H $\gamma$ )
		C II : 4267.00, 4267.26
		[ O III ]: 4363.21
		[ Fe II ]: 4177.21, 4243.98,
		4244.81, 4287.40, 4358.10, 4358.37, 4359.34,
		4413.78, 4414.45, 4416.27, 4452.11, 4457.95
		[ Fe V ]: 4227.20
4747	420	H I : 4861.36 (H $\beta$ )
		He I : 4713.14, 4713.37, 4921.93
		He II : 4685.68
		[ Ne IV ]: 4724.17, 4725.60
		Mg I : 4571.10
		[ Mg I ]: 4562.48
		[ Ar IV ]: 4711.33, 4740.20
		[ Fe II ]: 4632.27, 4774.74, 4814.55,
		4889.63, 4889.70, 4905.35
		[ Fe III ]: 4658.10, 4701.62, 4733.93,
		4754.83, 4769.60, 4777.88, 4813.90, 4881.11,
		4924.50, 4930.50
		[ Fe VII ]: 4893.40, 4942.50
5360	360	He II : 5411.52
		[ N I ]: 5197.90, 5200.26
		[ Cl III ]: 5517.71
		[ Ar III ]: 5191.82
		[ Ca V ]: 5309.18
		[ Fe II ]: 5184.80, 5220.06, 5261.61,
		5268.88, 5273.38, 5296.84, 5333.65, 5376.47,
		5412.64, 5413.34, 5527.33
		[ Fe III ]: 5270.30
		[ Fe VI ]: 5277.80, 5335.20, 5424.20, 5426.60, 5484.80
		[ Fe XIV ]: 5302.86
5484	206	He II : 5411.52
		[ O I ]: 5577.34
		[ Cl III ]: 5517.71
		[ Fe II ]: 5412.64, 5413.34, 5527.33
		[ Fe VI ]: 5424.20, 5426.60, 5484.80
5741	1836	H I : 4861.36 (H $\beta$ ), 6562.85 (H $\alpha$ )
		He I : 4921.93, 5015.68, (5876)
		He II : 5411.52
		[ N I ]: 5197.90, 5200.26
		[ N II ]: 5754.59, 6548.05, 6583.45
		[ O I ]: 5577.34, 6300.30, 6363.78
		[ O III ]: 4958.91, 5006.84
		Na I : 5889.95, 5895.92
		[ S III ]: 6312.06
		[ Cl III ]: 5517.71
		[ Ar III ]: 5191.82
		[ Ar V ]: 6435.10
		[ Ca V ]: 5309.18
		[ Fe II ]: 4889.63, 4889.70, 4905.35,
		4973.39, 5039.10, 5043.53, 5072.40, 5107.95,
		5111.63, 5158.00, 5158.81, 5184.80, 5220.06,
		5261.61, 5268.88, 5273.38, 5296.84, 5333.65,
		5376.47, 5412.64, 5413.34, 5527.33
		[ Fe III ]: 4881.11, 4924.50, 4930.50,
		4985.90, 4987.20, 5270.30
		[ Fe VI ]: 4972.50, 5145.80, 5176.00,
		5277.80, 5335.20, 5424.20, 5426.60, 5484.80,
		5631.10, 5677.00
		[ Fe VII ]: 4893.40, 4942.50, 5158.90, 5720.70, 6087.00
		[ Fe X ]: 6374.51
		[ Fe XIV ]: 5302.86

<sup>a</sup> Unable to find these wavelengths for H I in NIST database.

Continued

**Table 4.** Central Wavelengths of HST Image File Sequences (continued).

1	2	3
Wavelength (Å)	RMS Bandwidth (Å)	Known SNR Ions/Atoms with Wavelengths Å, Fesen et al. 1996
5921	672	He I : (5876) [ N II ]: 5754.59 Na I : 5889.95, 5895.92 [ Fe VI ]: 5631.10, 5677.00 [ Fe VII ]: 5720.70, 6087.00
6001	638	He I : (5876) [ N II ]: 5754.59 [ O I ]: 6300.30 Na I : 5889.95, 5895.92 [ S III ]: 6312.06 [ Fe VII ]: 5720.70, 6087.00
6564	54	H I : 6562.85 (H $\alpha$ ) [ N II ]: 6548.05, 6583.45
7996	646	He II : 8236.77 O I : (7774) [ Ar III ]: 7751.06 [ Cr II ]: 7999.85, 8125.22, 8229.55, 8308.39 [ Fe II ]: 7686.19, 7686.90 [ Fe XI ]: 7891.80 [ Ni II ]: 8300.99
8057	652	H I : 8345.55, 8359.00, 8374.48 <sup>a</sup> He II : 8236.77 O I : (7774) [ Ar III ]: 7751.06 [ Cr II ]: 7999.85, 8125.22, 8229.55, 8308.39, 8357.51 [ Fe XI ]: 7891.80 [ Ni II ]: 8300.99
8060	653	H I : 8345.55, 8359.00, 8374.48 <sup>a</sup> He II : 8236.77 O I : (7774) [ Ar III ]: 7751.06 [ Cr II ]: 7999.85, 8125.22, 8229.55, 8308.39, 8357.51 [ Fe XI ]: 7891.80 [ Ni II ]: 8300.99

<sup>a</sup>Unable to find these wavelengths for H I in NIST database.**Table 5.** HST filters containing ion/atom species important for SNR identification.

1	2	3
Wavelength (Å)	RMS Bandwidth (Å)	SNR Diagnostic Ions/Atoms with Wavelengths (Å)
5741	1836	H I : 4861.36 (H $\beta$ ), 6562.85 (H $\alpha$ ) [ N II ]: 5754.59, 6548.05, 6583.45 [ O I ]: 5577.34, 6300.30, 6363.78 [ O III ]: 4958.91, 5006.84
6001	638	[ O I ]: 6300.30
6564	54	H I : 6562.85 (H $\alpha$ ) [ N II ]: 6548.05, 6583.45

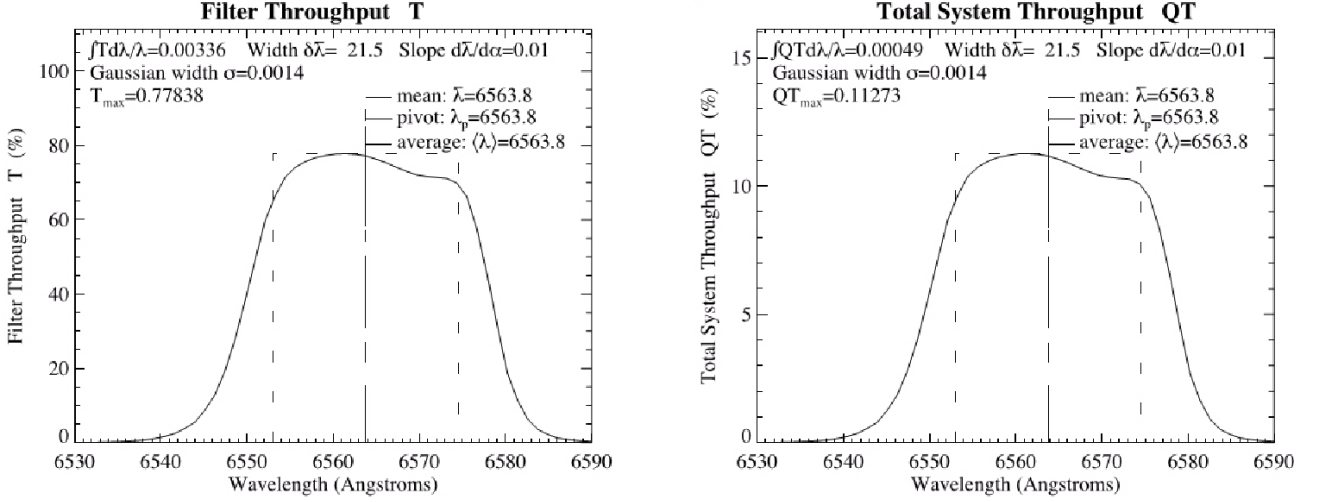
**Table 6.** Five candidates found in the H $\alpha$  HST image `u6713709r_drz.fits`. The positions are the J2000 coordinates reported by P04 (radio sources) and BL97 (optical candidates).

1	2	3	4	5	6	7	8
Optical Object	Radio Object	RA (h m s)	Dec ( $^{\circ}$ / ' / ")	[ S II ]:H $\alpha$ (BL97)	[ S II ]:H $\alpha$ (MWF11)	$\alpha \pm \Delta\alpha$ (P04)	Diameter (pc, MWF11)
	J005450–374030	00 54 50.28	–37 40 30.0		$0.32 \pm 0.12$	$-0.5 \pm 0.2$	
	J005450–374022	00 54 50.73	–37 40 22.2		$0.38 \pm 0.31$	$-0.3 \pm 0.1$	130
	J005451–373939	00 54 51.79	–37 39 39.6		$0.10 \pm 0.18$	$-0.1 \pm 0.2$	
	J005500–374037	00 55 00.58	–37 40 37.4		$0.25 \pm 0.09$	$-0.4 \pm 0.4$	
N300-S16		00 54 54.46	–37 40 35.46	0.70	$0.94 \pm 0.06$		57

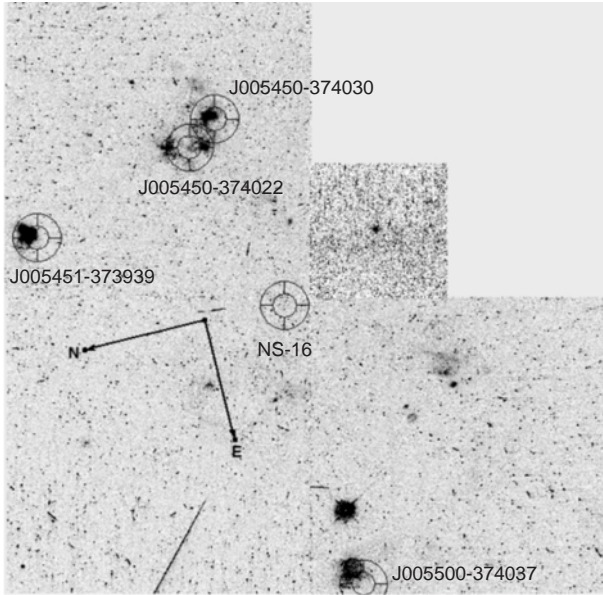
<sup>a</sup>Spectral index not determined.



## F656N (Wheel: 7 Pos: 2)



**Fig. 2.** The pass band characteristics of the N565 filter used on the WFPC2 of the HST. (STScI Institute, WFPC2 Observer's Handbook.)



**Fig. 3.** This image is contained in the file `u6713709r_drz.fits` and is here modified using the SAO's DS9 software. The image is a 400 second exposure centered on  $\alpha = 00\ 54\ 54.54$ ,  $\delta = -37\ 40\ 35.9$  and rotated  $\sim 104^\circ$  east of north. The high-resolution CCD (first quadrant) is centered on the nucleus of NGC 300. The five SNR candidates in this image are described in Table 6. The pandas are centered on each candidate with an inner circle diameter of  $6''$  ( $61.2\text{ pc}$ ).

Fig. 3 shows the image contained in this file in negative gray-scale. The five SNR candidates (Table 6) in the field of view are labeled and positioned by DS9 pandas. For comparison, Fig. 4 shows an image with the same position, size and rotation using DSS2-Red data from SkyView<sup>4</sup>. There is a slight difference in plate scale between the two images. The green line in Fig. 4 is the outline of the HST image's (Fig. 3) first quadrant, which contains the high-resolution CCD chip and the center of NGC 300. A bright star in the second quadrant which appears in the DSS2-Red image does not appear in the HST image (the DSS2-Red survey includes  $H\alpha$  wavelengths). There are some differences in the apparent distribution of gas in the DSS2-Red image compared to the HST image, probably due to the wider pass band of the DSS2-Red image. There is a triangular shaped optical artifact located to the north of the very bright star near the bottom of the DSS2-Red image. The four radio SNR candidates from P04 were not optically confirmed, the optical SNR candidate from BL97 was confirmed (MWF11, Table 2).

### 3. ANALYSIS OF THE RADIO SOURCES

When the HST image was displayed in false color<sup>5</sup> to show emission intensity (Figs. 5 through 8), a number of small sized, high intensity objects appeared in the image. Most likely, these are stars but that is difficult to determine from the available data. Hereafter these objects are referred to simply as "hot spots."

<sup>4</sup><http://skyview.gsfc.nasa.gov/>

<sup>5</sup>Using the Smithsonian Astrophysical Observatory's DS9 software package <http://hea-www.harvard.edu/RD/ds9/>.

Multi-frequency observations of egSNRs are limited by current technologies to radio, optical and X-ray emissions. None of the four radio sources found in this HST image are known to have emissions in other wavelengths. NS16 is observed only in optical.

### 3.1. Radio Observations Astrometry

The majority of the SNRs found in the Galaxy and the Magellanic Clouds are 50 pc or less in diameter (Strom 1996). If the same is assumed for NGC 300, then the SNRs in NGC 300 should be less than 5'' in diameter (at the assumed distance). The radio observations published in P04 used a circular beam width of 6''. Thus the typical SNR would not fill the beam of the radio telescope. This may have led to systematic astrometric errors.

In Figs. 3 through 8 the images (particularly the HST H $\alpha$  images) appear to show that there is such a systematic error in the astrometry of the radio candidates. The candidates show a consistent displacement to the southwest from the neighboring H II region. To account for radio emissions typical of an SNR without corresponding optical emissions (particularly when the candidate is  $\sim 50$  pc in diameter) the radio source should be located within or adjacent to the H cloud region so the candidate SNR's optical emissions could be masked. Though this is not the only possible explanation for such candidate SNR radio/optical observations (Duric 2000a,b, Panuti et al. 2000), it is likely (P04). No ground-based telescope of less than four meters should be able to separate this candidate from the H II region.

Radio observations with a higher resolving power or very high-resolution optical observations (HST) of the other radio sources are needed to resolve this possible apparent error.

### 3.2. SNR Candidate J005450–374030

This radio SNR candidate has a spectral index of  $-0.5 \pm 0.2$ . It has a [S II]:H $\alpha$  of  $0.32 \pm 0.12$  (MWF11, Table 2). Because the threshold value of the [S II]:H $\alpha$  is 0.4 this candidate could not be classified as an optical SNR (MWF11). Note however that the threshold value is within the error of the line ratio measurement. Fig. 5 shows the detail in the HST image at the location of this radio source. The candidate is located very near the edge of an H II region (H II C29, Soffner et al. 1996, 8''33, 85 pc away from center) showing no special structural properties.

A single image in H $\alpha$  does not allow us to determine the nature of the emissions from the region. An HST [S II] image would be very helpful. If there is a systematic error in the astrometry of the radio source, it is possible that the radio source is actually inside the H II region and the H II region would then interfere with the optical line ratio measurement. There are many known SNRs without optical emission so some of the NGC 300 SNRs would be expected to be in that group.

### 3.3. SNR Candidate J005450–374022

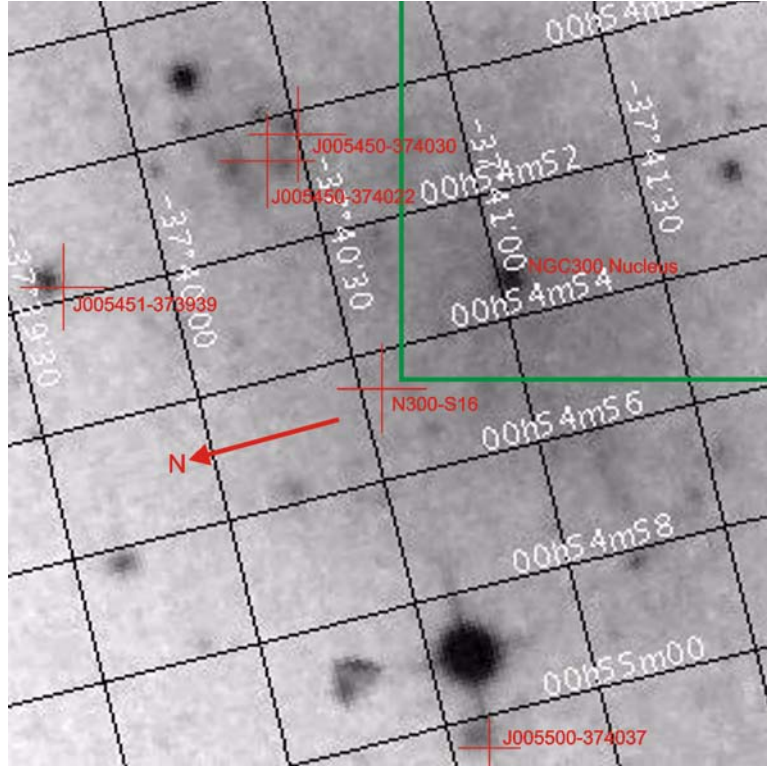
Fig. 6 shows detail in the HST H $\alpha$  emissions surrounding the radio source J005450–374022 which was classified as being either an SNR or an H II region in P04. The radio source is located on the edge of a large bubble in the H II region (76C, Deharveng et al. 1988, 2''36, 24 pc away from center). The small magenta (1'') circle pointed out by the red arrow shows the position of the candidate. The larger magenta ellipse (at the edge of which the radio candidate is located) outlines one possible border to the large bubble. A smaller region within the ellipse, which can be seen at the green emission level, may also be an SNR shock front within this H II region. Without proper [S II] data these remain simply possibilities of shock fronts.

The magenta ellipse is  $2''.54 \times 2''.38$  or  $25.9 \times 24.8$  pc with its semi-major axis aligned in declination. This is approximately the size of an SNR of about 10000 to 20000 years in age (Cioffi 1990). SNRs of this age are typically in the radiative, pressure-driven snowplow (PDS) stage which likely presents a bubble shaped appearance such as the one seen in this image. The [S II]:H $\alpha$  for this candidate is  $0.38 \pm 0.31$  (MWF11, Table 2). While the ratio is just below the critical value of 0.40, the large noise error in this measurement combined with the apparent structure next to the radio candidate raises serious concern about denying its classification as an optical SNR. This evidence is only available because of this high resolution image from the HST.

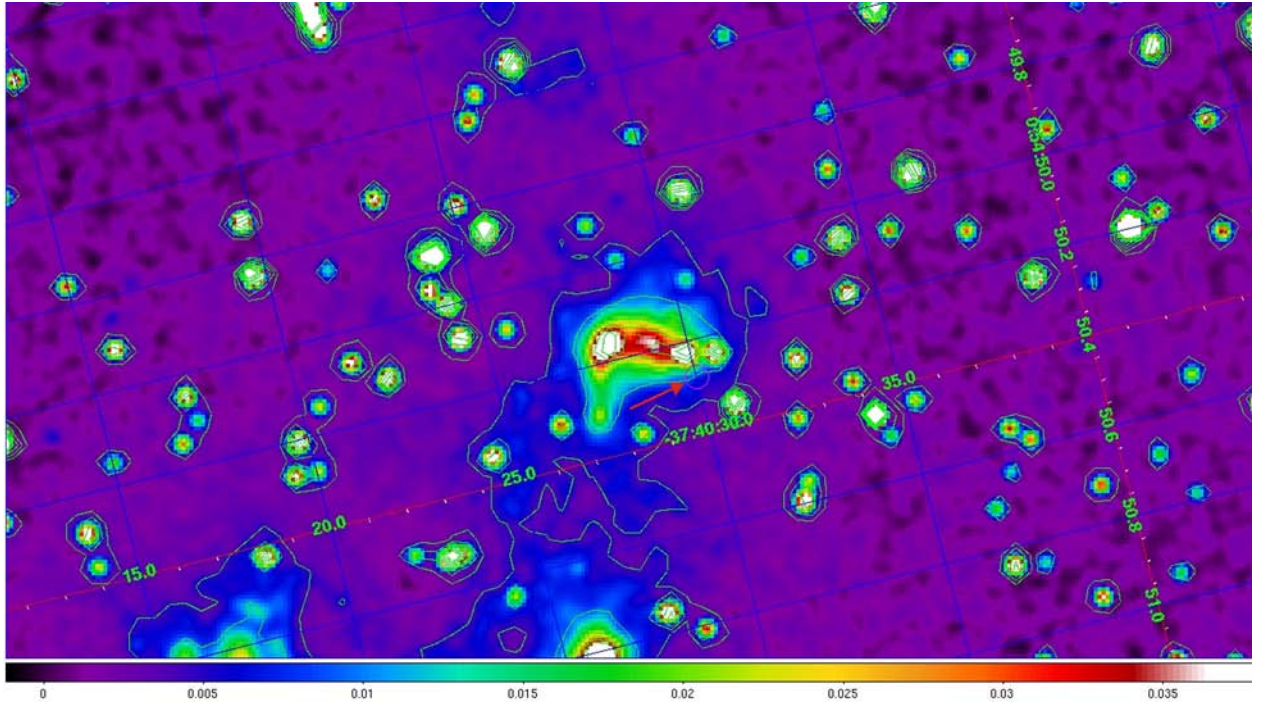
Also shown in Fig. 6 is a black line corresponding to the diameter measurement of this candidate (130 pc) as published in MWF11. This allows comparison of the poorer seeing conditions during those observations with the resolution of the HST image. This 130 pc ( $\sim 12''.7$ ) line is drawn centered on the published position of the radio source. Given this diameter comparison, the [S II]:H $\alpha$  may be that of the H II region with a large error caused by seeing conditions. No firm conclusion about the optical emissions of this radio source can be drawn from these data with the seeing conditions degrading the line ratio technique by increasing the measurement error.

### 3.4. SNR Candidate J005451–373939

Fig. 7 shows the location of J005451–373939, next to an H II region (#20, Bresolin et al. 2009, 3''80, 39 pc away from center). This H II region has a bright emission source off to one side at the same location as the radio source. It is not likely that this "hot spot" could be the SNR but it could be the known Wolf-Reyet (WR) star from (#18, Schild et al. 2003). There is also a known star association in this region (AS 57, Pietrzyński et al. 2001). Emissions from these stars (the association or the WR) can not be separated from the SNR candidate in the spectrographic slit and the WR star may also contribute to the radio emissions (we note the discussion of excessive radio emissions and the spectral index of Galactic WR stars in Montes et al. (2009)).



**Fig. 4.** This is the approximate equivalent to the HST image of Fig. 3 created with SkyView using DSS2-Red data. (Southern sky DSS image, Royal Observatory Edinburgh, Anglo-Australian Observatory, California Institute of Technology.)

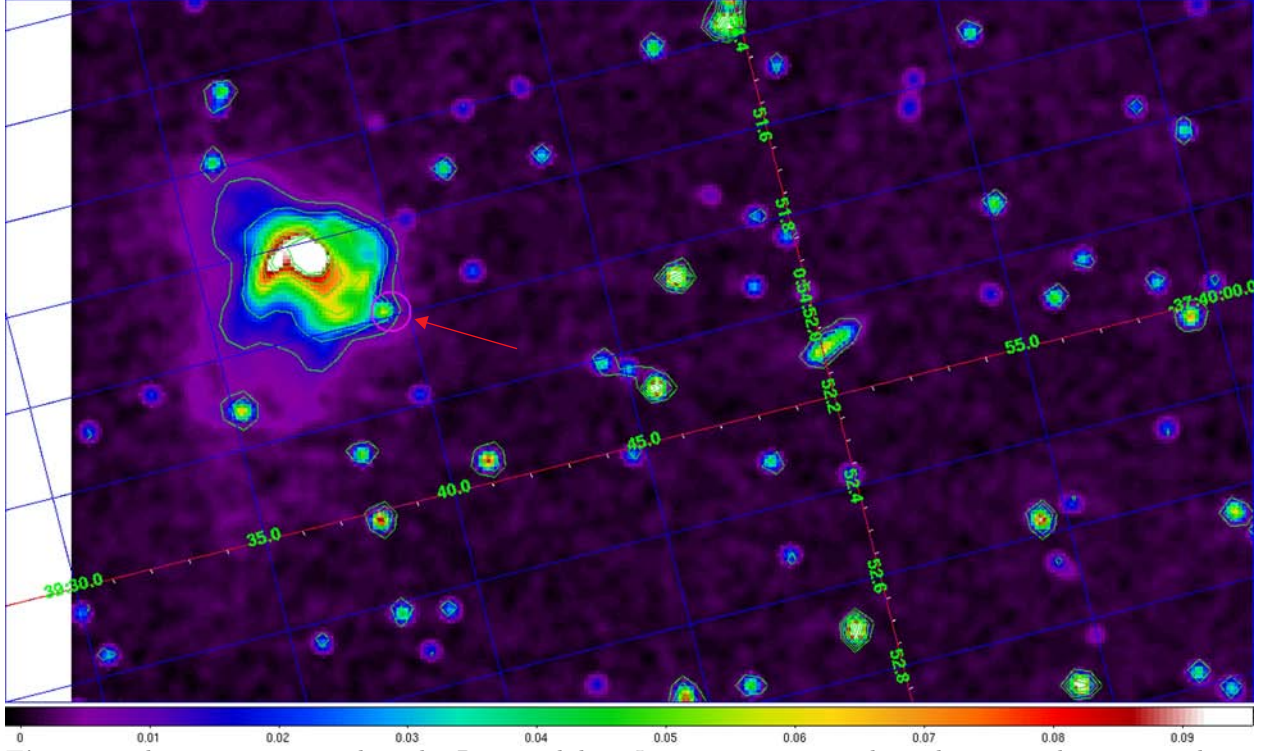


**Fig. 5.** The region surrounding the P04 candidate J005450-374030. The candidate location is shown by the magenta circle (red arrow pointing to it) just to the lower right of center of the intense H II emission. The circle is 1'' (10.2 pc) in diameter – about twice the  $1\sigma$  positional error. This false color image and contour plot was created with DS9. Zoom level: 4; Scale: Linear, 98%; Color: SLS; WCS: Equatorial, J2000; Analysis: Contours: 8, Smooth Parameters: 3, Gaussian.

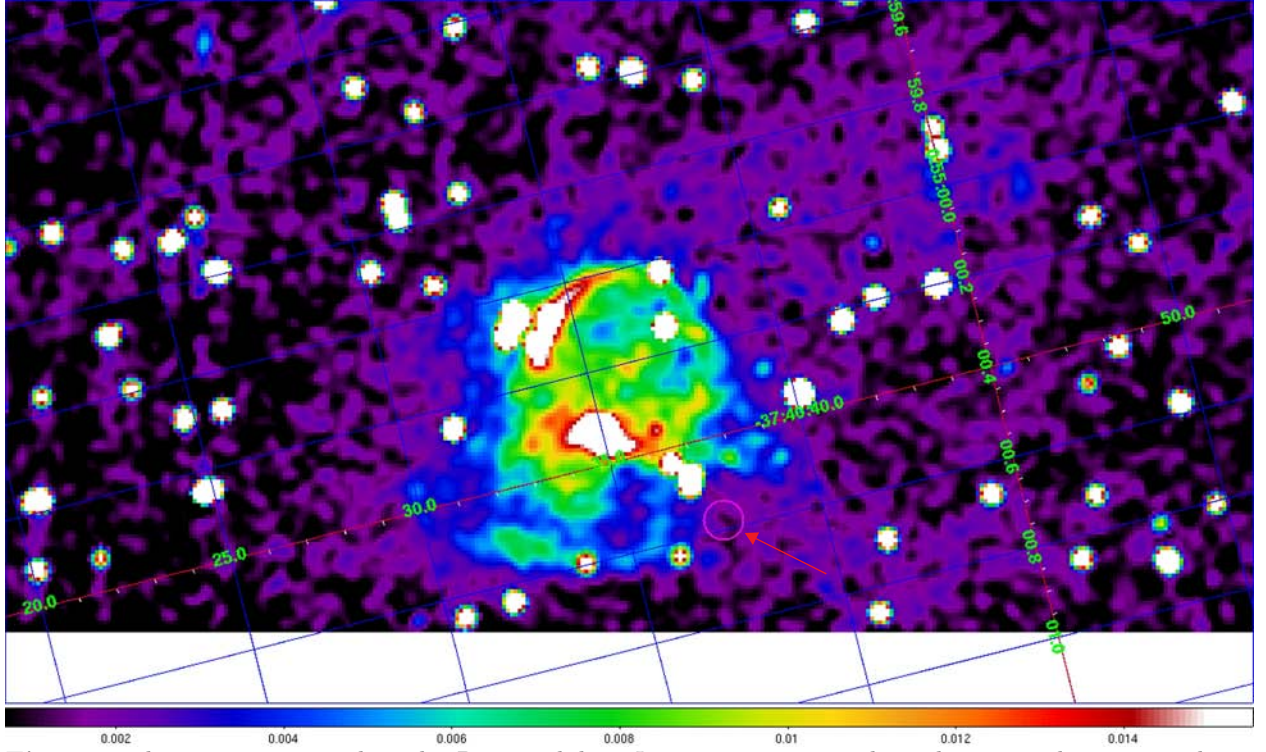




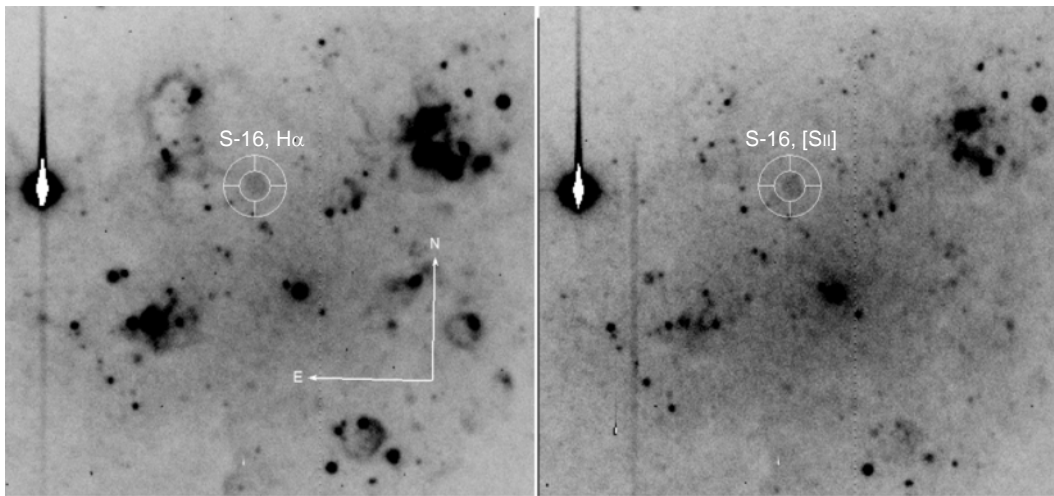




**Fig. 7.** The region surrounding the P04 candidate J005451–373939. The radio source location is shown by the magenta circle (red arrow pointing to it) 1'' (10.2 pc) in diameter ( $\sim 2\sigma$  positional error). This false color image and contour plot was created with DS9. Zoom level: 4; Scale: Linear, 99.5%; Color: SLS; WCS: Equatorial, J2000; Analysis: Contours: 10, Smooth Parameters: 3, Gaussian.



**Fig. 8.** The region surrounding the P04 candidate J005500–374037. The radio source location is shown by the magenta circle (red arrow pointing to it) 1'' in diameter ( $\sim 2\sigma$  positional error). This false color image was created with DS9. Zoom level: 4; Scale: Linear, 95%; Color: SLS; WCS: Equatorial, J2000. The contours were not used on this image.



**Fig. 9.** The left image is the  $H\alpha$  filter image from BL97. The right image is the  $[SII]$  filter image of BL97. An object which could be N300-S16 is visible as a faint smudge in these images. The inner circle of the panda is  $8''$  (81.6 pc) in diameter.

created by the diffuse ionized gas (DIG) of NGC 300. Because this candidate is not near any known HII region this “background ratio” could be subtracted from the MWF11 ratio and the result still exceeds the 0.4 critical  $[SII]:H\alpha$  ratio for indication of an SNR. However, the critical value is then within the small noise error.

Fig. 9 shows the BL97 images (set G) of N300-S16 zoomed and cropped. On the left is the  $H\alpha$  filter image and on the right is the  $[SII]$  filter image. In both images only a faint increase from the background is seen at N300-S16’s given position. The inner panda circle is centered on the N300-S16 coordinates with an apparent  $8''$  (81.6 pc) diameter – somewhat larger than the measured diameter. The outer panda circle is twice as big.

Fig. 10 zooms in to the HST  $H\alpha$  image at N300-S16. In this figure there is no evidence of any structure in the  $H\alpha$  emissions. It is possible that the  $[SII]:H\alpha$  test for SNR character has failed. No satisfactory explanation of the apparent contradiction between the two observations has been found.

## 5. THE ASTROMETRY OF BL97

There is contradictory evidence for the existence of an SNR at the location labeled as N300-S16. BL97 and MWF11 have high  $[SII]:H\alpha$  values but the HST  $H\alpha$  image shows no indication of any structure in the vicinity. In an attempt to find a resolution we checked the astrometry of the BL97 CCD images. The CCD images used for the BL97 optical survey were provided by William Blair. These images were sets and each set was labeled with the letters ‘C’ through ‘J.’ Each lettered image set contained a continuum image, an  $H\alpha$  filter image and an  $[SII]$  filter image with the same center, plate scale and rotation (see BL97). The files did not include the ‘I’ image (which was used in BL97) but did include a ‘K’ image which was not used in BL97. The ‘I’ image contained SNR candidates N300-S12 and N300-S17. The ‘C’ and ‘K’ images contained no SNR candidates

and thus were not considered in this analysis.

The images are in the flexible image transport system (FITS) format, 800 pixels on a side, with a field of view of approximately  $0^\circ 087$ . The central coordinates of each image are given in Table 2 of BL97. The image files did not have a FITS WCS and the  $x$ -axis needed to be inverted to obtain the correct orientation (north up and east to the left). A WCS was thus created for each of the FITS files.

### 5.1. Finding Known Objects Within the Images

A list of 2MASS point sources from the on-line catalog was created from a search for all sources within a radius of one arc-degree of the center of NGC 300. The program used a dynamic array of a record structure to store the coordinates of each 2MASS point source along with a “distance” data field, which is discussed below. NASA’s SkyView was used to create DSS and DSS2-Red FITS file images which appeared equivalent to the BL97 images. The SkyView 2MASS images were used as a visual aid in associating bright, easy to locate 2MASS point sources with optical counterparts in each BL97 image. To reduce errors in calculating the image scale the chosen sources were always closer to (but not at) the edge of the image.

The BL97  $(x, y)$  coordinates and 2MASS  $(\alpha, \delta)$  coordinates of each of the three 2MASS sources from each image were entered into a computer program. The estimated coordinates were compared to the list of catalog coordinates. The comparison created a distance measurement between the estimated and the catalog coordinates which was stored in the “distance” data field of each point source data record structure. The source list was then sorted on the distance field and the source with the smallest distance was put at the top of the sorted list. The catalog coordinates at the top of the list were the coordinates actually used for calculating the BL97 image center, plate scale and rotation.



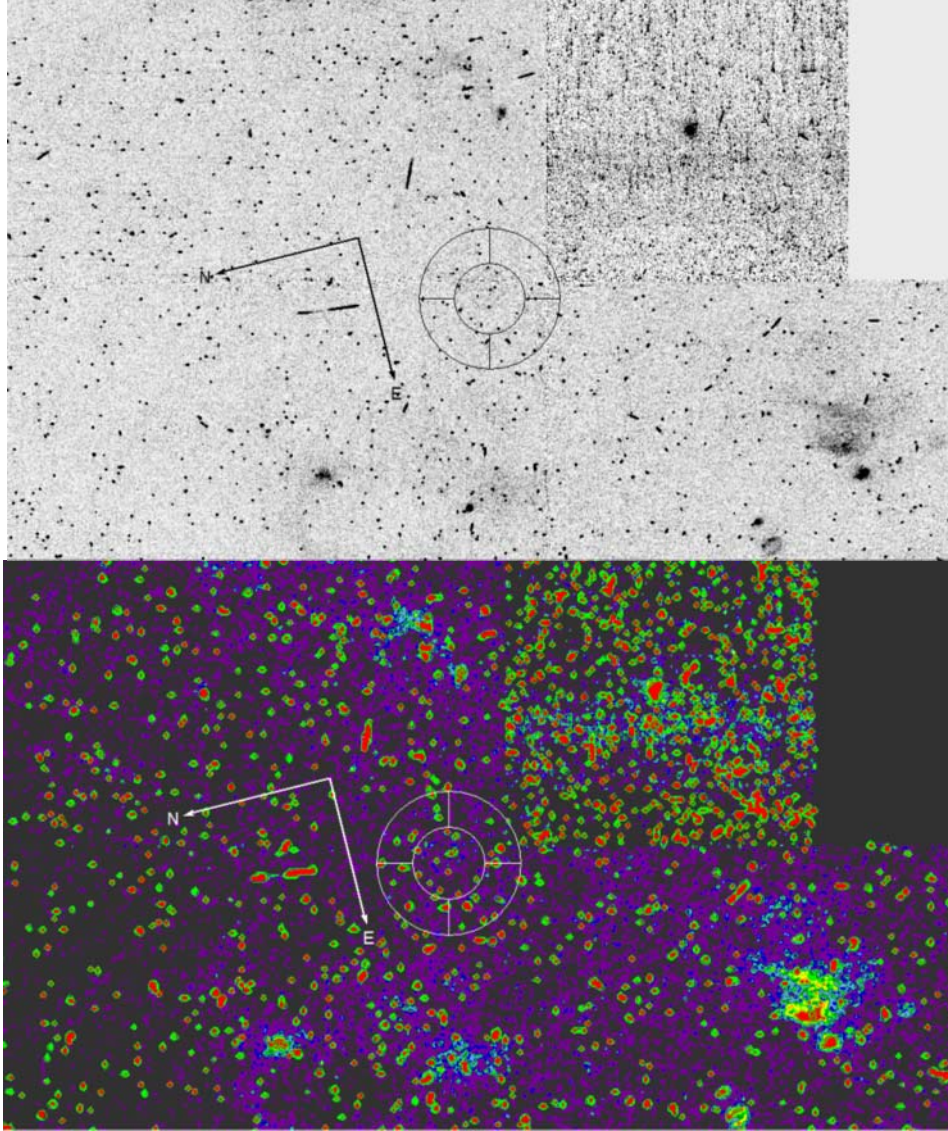
## 5.2. Program Results and Astrometry of the SNRs

Table 7 shows the program results for the calculation of the image centers. Column 1 is the image identification letter. Column 2 is the image center as given in BL97's Table 2. Column 3 is the calculated center from the program and column 4 is the distance (or offset) between them in arcseconds.

After studying the continuum,  $H\alpha$  filter and  $[S II]$  filter images for each lettered image set, any difference in the center coordinates between the image filter types appeared to be negligible. Because the

filters used to make the  $H\alpha$  and  $[S II]$  images were narrow band, it was impossible to locate 2MASS point sources within those images. Thus, center coordinates and plate scale values were calculated only for the continuous images and then applied to all image filter types for the image field (lettered set). Based on the calculated WCS for each image, Table 8 shows the reported equatorial and calculated pixel coordinates for each of the BL97 SNR candidates.<sup>7</sup>

N300-S16 is the only BL97 SNR candidate to appear in an HST image. Fig. 11 shows the BL97  $H\alpha$  filter and  $[S II]$  filter images of N300-S1 (image set D).<sup>8</sup> The pandas mark the coordinates given for



**Fig. 10.** The HST  $H\alpha$  image of N300-S16 shown in inverted gray-scale on top and in AIPS0 with contours at the bottom. In either case no discernible structure is seen in the  $H$  gas at the reported location of this candidate SNR. The panda inner circle is  $8''$  ( $81.6$  pc) in diameter.

<sup>7</sup>Tables 8-10 are available online at <http://saj.matf.bg.ac.rs/184/pdf/Tables8-10.pdf>.

<sup>8</sup>Figures 11-27 are available online at <http://saj.matf.bg.ac.rs/184/pdf/Figs11-27.pdf>.

**Table 7.** Image center from BL97 and 2MASS calibrated.

1	2	3	4
Image	BL97 Center (h m s ° ' ")	Calibrated Center (h m s ° ' ")	Difference (")
D	00 54 23.030 -37 35 50.600	00 54 22.966 -37 35 51.374	1.081
E	00 55 38.770 -37 40 47.900	00 55 38.793 -37 40 48.215	0.420
F	00 55 13.460 -37 40 51.900	00 55 13.550 -37 40 52.086	1.082
G	00 54 47.360 -37 40 50.400	00 54 47.359 -37 40 51.305	0.905
H	00 54 22.840 -37 40 50.700	00 54 22.856 -37 40 52.042	1.356
J	00 55 14.000 -37 45 50.300	00 55 13.987 -37 45 50.463	0.227

N300-S1 in BL97 as calibrated by 2MASS point sources. The BL97 astrometry is good – certainly within the seeing conditions reported in BL97 ( $\sim 1''$ ). Fig. 12 shows a comparison of the reported position (BL97, on the left) of N300-S16 with the apparent center of the assumed image of the candidate (HST, on the right). The offset in these positions ( $0''.7$ ) is again within the reported seeing conditions. The logical next step would be to calibrate the HST images to 2MASS and study the HST H $\alpha$  image of N300-S16 in comparison to the BL97 images.

## 6. HST IMAGE ASTROMETRY

The HST FITS file headers indicated that the HST Guide Star Catalog (GSC) version 1.x was used for the image astrometry. According to the Astrometry Department of the United States Naval Observatory<sup>9</sup> the accuracy of this catalog is approximately 500 mas. The biggest problem with the GSC is that it does not provide for the proper motion of the stars.

The attempt to calibrate the HST images to the 2MASS catalog using the procedure and software from the BL97 analysis proved to be impractical because of the high-resolution of the HST images. The difference in resolution between the SkyView 2MASS images and HST images made it impossible to correctly identify 2MASS point sources in the HST images with reasonable confidence. Therefore we used an alternative method to check the astrometry.

### 6.1. Alternate Analysis

An analysis of the HST and DSS2-Red files' WCS accuracy against the BL97 2MASS-calibrated images was made by comparing the position of a bright star which could be identified in all image sets. The first comparison, between BL97 and the HST images, is shown in Fig. 13. On the left is the BL97 G H $\alpha$  image containing N300-S16 and on the right is the HST H $\alpha$  image containing N300-S16. The position of the selected star within these images is listed in Table 9. In this table column 1 is the image used for the observation (BL97, HST, DSS2-Red), columns 2 and 4 are the J2000 RA and Dec coordinates of the center of the image of the star as found using DS9. Table 9 columns 3 and 5 are the difference between the positions (with BL97 as the

reference) in arcseconds. Column 6 is then the difference in position (offset) between the two images, measured in arcseconds. Between the BL97 and HST images there is an offset of  $0''.15$ .

The second comparison was made between BL97 and the DSS2-Red images used for the finding maps for the observations in MWF11 and is shown in Fig. 14. Between the BL97 and DSS2-Red images there was an offset of  $0''.13$ . Mickaelian (2004) measured the astrometry of the DSS2-Red images based on AGN positions and found an accuracy of  $0''.33$ . These comparisons demonstrate that any error in the astrometry of the images used in BL97 was within the reported seeing conditions.

### 6.2. Estimating Positional Error

Because of the difficulty in recalibrating the HST image WCS, the alternative procedure of convolving the position errors was used. In Section 5.2, it was found that the position error in the BL97 images was  $\sim 1''$  (mainly due to seeing). This was convolved with the approximate positional error of the HST images from Table 9 to find a positioning error of the BL97 SNRs on the HST images: The result is  $1''.0$  – essentially the seeing conditions of BL97. The same was done with the radio SNRs from P04. According to Table 2 of P04 these candidates were observed with a circular beam width of  $6''$ . Convolution 10% of that gives a  $1\sigma$  positioning error for the candidates as  $0''.62$  or  $0''.6$ .

These values were used as the  $1\sigma$  positioning error for finding the SNR candidates within the HST images. We see that positioning (astrometry) errors are mainly due to seeing conditions and telescope pointing.

## 7. SEEING AND TELESCOPE POINTING

Some of the candidates from BL97 which were observed in MWF11 returned no signal. In an attempt to determine the cause, the telescope pointing accuracy was analyzed by superimposing the slit camera images on top of the observation finding maps. Figs. 15 and 16 show examples of the results. Fig. 15 shows a very good alignment with N300-S2. In Fig. 16 there is a confusion of sources surrounding N300-S11. Both N300-S11 and J005442-374313 were clearly within the slit and the seeing conditions were

<sup>9</sup>[http://ad.usno.navy.mil/star/star-cats\\_rec.shtml](http://ad.usno.navy.mil/star/star-cats_rec.shtml)



undoubtedly allowing J005443–374311 to also flow through. Exactly what was being measured here is uncertain but it was labeled as N300-S11 (MWF11) because it was the principal candidate for that observation. This observation also included emission from the neighboring H II region.

For N300-S11, BL97 reports  $[\text{S II}]:\text{H}\alpha = 0.66$  (based on CCD images, not spectra) and MWF11 reports  $[\text{S II}]:\text{H}\alpha = 0.30 \pm 0.12$ . Because of source confusion, neither of these measurements can be trusted. BL97 claimed seeing of  $1''$ , but these sources would be confused in any spectrometer's slit. Due to seeing conditions creating telescope positioning error and the astrometry error in the radio observations it is not possible to determine if J005442–374313 and N300-S11 are actually the same object. Measurements of the diameter of N300-S11 (BL97, MWF11) are greater than 100 pc. This large diameter may be caused by source confusion with the H II region or it may be a multiple supernovae site. Unfortunately multiple SNR sites are usually linked to OB associations and there are no known OB associations in this region. The only way to resolve this problem is with a higher-resolution telescope with no seeing problems – the HST. Unfortunately no archival HST file contained an image of these candidates.

Fig. 17 shows the slit camera image on the finding map of N300-S16 used for the observations for MWF11. For this candidate the telescope pointing was erroneous (by  $1''$  to  $2''$  or 10 to 20 pc) and yet a  $[\text{S II}]:\text{H}\alpha = 0.94 \pm 0.06$  was measured. The cause of the poor alignment is difficult to determine – equipment or seeing conditions. It is most likely due to seeing conditions which were on the order of two to three arcseconds for most nights of the observing run. The seeing conditions may have allowed most of the flux from N300-S16 through the slit but it is then difficult to account for the higher ratio compared to the BL97 results (0.70).

However, we still have the real problem with this observation – the HST image containing N300-S16 shows no evidence of the candidate's existence (Fig. 10).

## 8. ANALYSIS OF THE BL97 IMAGES

BL97 obtained long slit spectra for  $\sim \frac{2}{3}$  of the candidates found in the optical filter survey for that paper. With 100% confirmation ( $[\text{S II}]:\text{H}\alpha > 0.4$ ) of this  $\frac{2}{3}$  sample, there was confidence in all 28 candidates being confirmed as SNRs (Long 1996). N300-S16 was not a member of that  $\frac{2}{3}$  subset.

### 8.1. Candidate Image Profiles

Optical emissions are expected to be greatest during the PDS stage (Cioffi 1990). This stage is typically 50 pc at maximum diameter. At the distance to NGC 300, 50 pc is equivalent to  $4''.9$ . The images for BL97 are  $800 \times 800$  pixels at 5.3 arcminutes square (BL97, Long 1996) and the pixel size is then  $0''.4$ . At

the distance to NGC 300, this is about 4 pc. The maximum theoretical size of a PDS stage SNR on a BL97 image would be about 12 pixels in diameter or an area of about 110 pixels squared.

The  $\text{H}\alpha$  and  $[\text{S II}]$  images used for BL97 were analyzed for candidate image profiles with the intention of investigating the actual number of CCD pixels used to determine the  $[\text{S II}]:\text{H}\alpha$  ratio of the candidate SNRs. Using the “Co-Add” command of SBIG's *CCDOps* program<sup>10</sup> the image of each SNR candidate was stacked to form a composite image. This was done with both the  $\text{H}\alpha$  and  $[\text{S II}]$  images. The  $x$  and  $y$ -axis profiles of these composite images are shown in Figs. 18 and 19.

Fig. 18 shows two renditions of the stacked  $\text{H}\alpha$  candidate images. These two renditions differ by the display intensity scale and the number of plotted contour lines. The images were stacked so that all SNR candidates were located (the pixel corresponding to their equatorial coordinates) at the same resultant pixel. The rendition on the right has the crosshairs and profiles centered on the emissions peak which happens to be the same as the coordinate pixel. As an example of the data, Table 10 shows a  $9 \times 9$  grid of the CCD pixel values from Fig. 18. Fig. 20 shows a surface plot of these values which clearly shows a high signal level against the background. Fig. 21 shows a line plot of the values along the  $x$ -axis (E-W) at  $y$ -axis (N-S) row number 165. The full-width half maximum is about  $10 \times 12$  pixels ( $\sim 41 \times 49$  pc) which is about the same as the expected 12 pixel diameter area and the most expected diameter for candidates at this stage of their evolution.

In Fig. 19 the  $[\text{S II}]$  frames were stacked but did not include the H and J images as these were too noisy for this procedure. Various techniques were used to remove the noise<sup>11</sup>. The problem was that the  $[\text{S II}]$  signal on these images was too close to the noise floor. The H and J images contained candidates: N300-S2, S3, S4, S7 (H image); S20, S26, S27 (J image). Fig. 19 shows two renditions of the stacked  $[\text{S II}]$  images. On the left, the crosshairs (and the profiles) are centered on the coordinate pixel (508, 164). On the right, the crosshairs (and the profiles) are centered on the pixel with maximum value (509, 162). This difference in pixel position corresponds to about 4 (RA) and 8 pc (Dec) in physical distance, respectively. The SNR positions were determined from the  $\text{H}\alpha$  images based on H II regions but this also implies that the maximum  $\text{H}\alpha$  and the maximum  $[\text{S II}]$  emissions are on the opposite side of the SNR. The surface plot of the  $[\text{S II}]$  emissions shown in Fig. 22 shows multiple peaks distributed around the centers of the SNR candidates. The FWHM (centered on the highest peak) is about  $9 \times 9$  pixels ( $\sim 37 \times 37$  pc).

The estimated FWHM of  $\text{H}\alpha$  and  $[\text{S II}]$  profiles were convolved with the FWHM of a small star from the  $[\text{S II}]$  J image. The profile of the star is shown in Fig. 24. When the FWHM of the stacked

<sup>10</sup> Available from Santa Barbara Instruments Group, <http://www.sbig.com/sbwhtmls/ccdopsv5.html>

<sup>11</sup> *CCDOps* has some noise treatment routines such as, “Kill Hot Pixels,” “Kill Cold Pixels,” “Smooth Pixels.”

[ S II ] image profile (Fig. 23) was convolved with the small star image profile (Fig. 24) the resulting apparent size of the stacked [ S II ] regions was,

$$\sqrt{(28 \text{ pc})^2 - (15 \text{ pc})^2} = 23 \text{ pc} = 6 \text{ pixels} . \quad (1)$$

This was taken as the average SNR candidate diameter, so the extent of the image area of the [ S II ] emissions was approximately 28 pixels squared. Because the images could not be flux calibrated, the number of pixels actually contributing to the [ S II ]:H $\alpha$  could not be estimated with reasonable confidence. The theoretically expected maximum number of pixels was about 110. With only 28 pixels (in a stacked image of all sources) contributing to the flux measurement of spectra, flux measurements of BL97 may well have large errors. Weak signals may lead to large flux measurement error (as an overestimate), an effect noted in other egSNR observations by Blair *et al.* (1981).

## 8.2. Images of N300-S16

Fig. 11 shows the BL97 [ S II ] and H $\alpha$  images zoomed in on N300-S16. As measured in MWF11 the H $\alpha$  and [ S II ], flux levels for N300-S16 were about  $2 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$  which was only about 100 CCD pixel counts above the background. The spectrum shows a high-level background along the slit. Fig. 17 shows the slit was at least an arcsecond away from the reported coordinates of the candidate. The spectrum from MWF11 of N300-S16 is shown in Fig. 26.

In the HST H $\alpha$  image of N300-S16 (Fig. 10) there was little if any evidence of a disturbance in the H gas in the region of the location of N300-S16. While the equivalent images of the radio SNR candidates (see Figs. 5, 6, 7 and 8) show strong evidence of such, even though the measured [ S II ]:H $\alpha$  ratio does not confirm them as optical SNRs. There is also no known X-ray emission associated with N300-S16.

Fig. 27 shows a 3-D plot of the BL97 CCD H $\alpha$  (top) and [ S II ] (middle) image pixel counts of N300-S16. The pixels plotted (a  $27 \times 27$  array centered on the candidate's coordinates) are shown by the green box in the image tile on the left. The plot is shown on the right. In both cases the emission is only slightly above the background but clearly discernible. The size of the candidate (50 pc) implies it to be at the end of its PDS stage. The 3-D plot on the bottom of the figure is a plot of the result of dividing the [ S II ] pixel value by the H $\alpha$  pixel value. The apparent SNR disappears in the high-level DIG noise.

There are spectral line characteristics other than [ S II ]:H $\alpha$  to support the SNR candidacy (e.g. the presence of [ O I ] in the spectrum) and there is

a selection effect biased toward finding SNRs away from H regions in optical surveys (P04). Type Ia SN which are away from any H region are then located in less dense ISM. As a result, there is a lower abundance of shocked material to produce the [ S II ] (and other metallic) spectral lines in the SNR remnant. Such SNR spectra may be dominated by the Balmer lines (e.g. SN1006, Tycho's and Kelper's SN) and are generally missed by optical surveys (Pannuti *et al.* 2000). A check of the data from MWF11 shows no Balmer dominance in the spectrum of N300-S16, particularly with a [ S II ]:H $\alpha$  of 0.6 and a high level of [ O I ] usually associated with shocked ISM. High-resolution imaging should be used to confirm the existence of this SNR. This imaging could be done with space-based telescopes (HST) or with four-meter class telescopes. Better results may be obtained with ground-based telescopes using adaptive optics.

## 9. CONCLUSION

A careful investigation of the data collected with the telescopes and instruments typically used for the discovery and confirmation of extra galactic supernova remnants reveals that the reliability of these techniques may be questionable. A strict flux density measurement error analysis shows that large errors in the [ S II ]:H $\alpha$  ratio occur when the ratio is based on low flux density levels (MWF11). Seeing conditions lead to blurring of telescope positioning and thus pointing errors which may impinge in the reliability of the flux measurements and on the confidence of exactly what object is being measured (Fig. 17).

The seeing conditions also introduce error into the astrometry of the sources which is generally as large as the telescope pointing error. If these errors add in the same direction there is a possibility that the spectrograph slit is completely off the candidate. In cases where the target galaxy has high level emission from its diffuse ionized gas (such as NGC 300), the spectroscopic signal may not be sufficient to confirm the presence of a supernova remnant. High resolution optical images from space-based telescopes or from ground-based telescopes using adaptive optics may be necessary for confirming the existence of these extra-galactic supernova remnants.

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ОПТИЧКЕ ТЕХНИКЕ ПОСМАТРАЊА ВАНГАЛАКТИЧКИХ  
ОСТАТАКА СУПЕРНОВИХ: СЛУЧАЈ NGC 300

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*Оригинални научни рад*

У овој студији представљамо посматрачке технике коришћене за спектроскопску идентификацију кандидата за остатке супернових у NGC 300 из Скулптор Групе галаксија. Главни циљ ове студије односи се на испитивање сврсисходности коришћења параметра  $[Si]:H\alpha \geq 0.4$  у идентификацији вангалактичких остатака супернових. Посебна пажња посвећена је чињеници да посматрачки услови као што су видљивост и позициона пре-

цизност телескопа имају доминантну улогу у коришћењу ове технике. Ова студија базира се на архивским посматрањима са Hubble Space Telescope на којима су идентификовани остаци супернових у Скулптор галаксији – NGC 300. Наши резултати показују да је применљивост ове технике веома непоуздана и да су неопходна посматрања у вишој резолуцији приликом идентификације вангалактичких остатака супернових.