



# Truly eccentric – II. When can two circular planets mimic a single eccentric orbit?

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## ABSTRACT

When, in the course of searching for exoplanets, sparse sampling and noisy data make it necessary to disentangle possible solutions to the observations, one must consider the possibility that what appears to be a single eccentric Keplerian signal may in reality be attributed to two planets in near-circular orbits. There is precedent in the literature for such outcomes, whereby further data or new analysis techniques reveal hitherto occulted signals. Here, we perform suites of simulations to explore the range of possible two-planet configurations that can result in such confusion. We find that a single Keplerian orbit with  $e \gtrsim 0.5$  can virtually never be mimicked by such deceptive system architectures. This result adds credibility to the most eccentric planets that have been found to date, and suggests that it could well be worth revisiting the catalogue of moderately eccentric ‘confirmed’ exoplanets in the coming years, as more data become available, to determine whether any such deceptive couplets are hidden in the observational data.

**Key words:** methods: numerical – techniques: radial velocities – planets and satellites: detection.

## 1 INTRODUCTION

Roughly 30 years ago, we saw the dawn of the Exoplanet Era, with the detection of the first planet-mass objects orbiting other stars (Campbell, Walker & Yang 1988; Latham et al. 1989; Wolszczan & Frail 1992; Mayor & Queloz 1995). In the decades since, we have become ever more adept at observing the minute variations in the behaviour of stars that hint at the presence of their planetary companions (Fischer et al. 2016).

Through the late 1990s and early 2000s, the radial velocity (RV) technique dominated exoplanetary science, revealing a variety of planets that was far greater than we had previously imagined. We found giant, Jupiter-mass planets orbiting perilously close to their host stars (e.g. Butler et al. 1997; Hellier et al. 2011; Wright et al. 2012), as well as planets moving on orbits so eccentric that they more closely resemble the orbits of comets than the planets in our own backyard (e.g. Wittenmyer et al. 2007; Tamuz et al. 2008; Wittenmyer et al. 2017b). We also found many systems with multiple planets moving on orbits locked in mutual mean-motion resonance (e.g. Robertson et al. 2012; Wittenmyer, Horner & Tinney

2012b; MacDonald et al. 2016; Wittenmyer et al. 2016a; Gillon et al. 2017).

Over the past decade, it has become clear that it is possible for a multiple-planet system containing resonant planets on near-circular orbits to masquerade as a single, moderately eccentric planet in typical sparsely sampled radial velocity data (e.g. Anglada-Escudé, López-Morales & Chambers 2010; Wittenmyer et al. 2013b; Boisvert, Nelson & Steffen 2018). Since researchers typically look for the simplest explanation for a given signal, such systems are often initially reported as single, moderately eccentric planets. The true multiplicity of these systems is then only revealed after further observations are carried out (e.g. Wittenmyer et al. 2012a; Kürster et al. 2015; Trifonov et al. 2017). For this reason, in 2013, we carried out a pilot study examining the likelihood that several moderately eccentric exoplanetary systems within the published literature might actually be such multiple planet systems, masquerading as single worlds (Wittenmyer et al. 2013b). On the other hand, a number of extremely eccentric planets have been found (e.g. Jones et al. 2006; Wittenmyer et al. 2007; Tamuz et al. 2008; Marmier et al. 2013; Wittenmyer et al. 2017b), whose best-fit orbits are so extreme that it seems unlikely that they could be reproduced by any given combination of two exoplanets moving on resonant, near-circular orbits.

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This, then, poses an obvious question – how eccentric can an orbit be before one can be truly confident that we are observing a single planet on a highly eccentric orbit, rather than poorly sampling a multiple planet system. Clearly, there exists a threshold for which no multiple planet solution can explain a highly eccentric orbit. Equally, there exists a range of single-planet orbital eccentricities that could readily be explained by invoking multiple planets on near-circular orbits.

In this work, we attempt to answer that question, in order to both strengthen confidence in the interpretation of extremely eccentric single-planet systems, and to identify the most dangerous regime of eccentricity space, for which the risk is greatest that a given multiple planet system will be misidentified as a single, moderately eccentric world.

In Section 2, we describe the approach we take to address this question, detailing how we created simulated radial velocity data sets to model the effects of multiplicity and sparse data sampling on the type of solutions found for a given planetary system. In Section 3, we present the results of our analysis, before moving on to discuss those results, and draw our conclusions, in Section 4.

## 2 SIMULATION APPROACH

In this section, we describe the procedures for producing the simulated radial velocity data sets. The stochasticity of real observational sampling in combination with the well-known sampling biases induced by telescope scheduling constraints (O’Toole et al. 2009) can result in poor detectability at certain orbital periods. Perversely, we wish to embrace this real-life pathology to provide the best possible assessment of the degree to which observers might be bamboozled by circular double systems masquerading as single eccentric planets.

### 2.1 Sampling

To create simulated observation times for this experiment, we attempted to reproduce the sampling properties of real radial velocity data. Following the simulation procedure in Wittenmyer et al. (2013a), we made the following assumptions: (1) one observation in a 10-night block every 30 d (bright-time scheduling), (2) the target is unobservable for four consecutive months every year, and (3) poor weather randomly prevents the observation 33 per cent of the time. These conditions were selected first because planet-search programs are usually allocated time in bright lunations owing to the brightness of the targets, and secondly, for a mid-latitude site such as the Anglo-Australian Telescope, with planet-survey targets distributed randomly in Right Ascension, the average target is unobservable for four months in a year.<sup>1</sup> This procedure generated a string of 50 observation epochs for each of 10 000 simulated stars, for each scenario tested herein.

We also explored more realistic sampling by drawing the observation times from real data sets. The 18-year Anglo-Australian Planet Search (AAPS; e.g. Carter et al. 2003; Tinney et al. 2011; Wittenmyer et al. 2016b, 2017a) has 90 stars for which more than 50 epochs were obtained. We generated strings of 50 epochs as follows: for each of the 90 real AAPS data files, we selected a 50-epoch window, then frame-shifted it by one until reaching the end of the data set. In this way, a real data set with, e.g. 55

<sup>1</sup>Here, we define ‘unobservable’ to mean that the target spends less than 1 h at an airmass less than 2.

**Table 1.** Recovery of two-planet solutions.

Scenario	Fraction Recovered
Ia	100.00%
Ib	100.00%
Ic	99.93%
Id	99.96%
IIa	100.00%
IIb	100.00%
IIc	99.98%
IId	99.90%
IIIa	98.69%
IIIb	98.94%
IIIc	91.73%
IIId	92.94%

epochs would generate six lists of 50-epoch samples, preserving the sampling characteristics of the real data. The result is 3871 lists of 50 observational epochs, each drawn from a real AAPS target and hence preserving all the associated idiosyncrasies of real data. Sets of 50 observation times for each of the 10 000 simulated data sets were then drawn at random (with replacement) from this pool. While the replacement means that some simulated data sets had identical observation times, we emphasize that the simulated velocity measurements are different, as described below.

### 2.2 Noise model

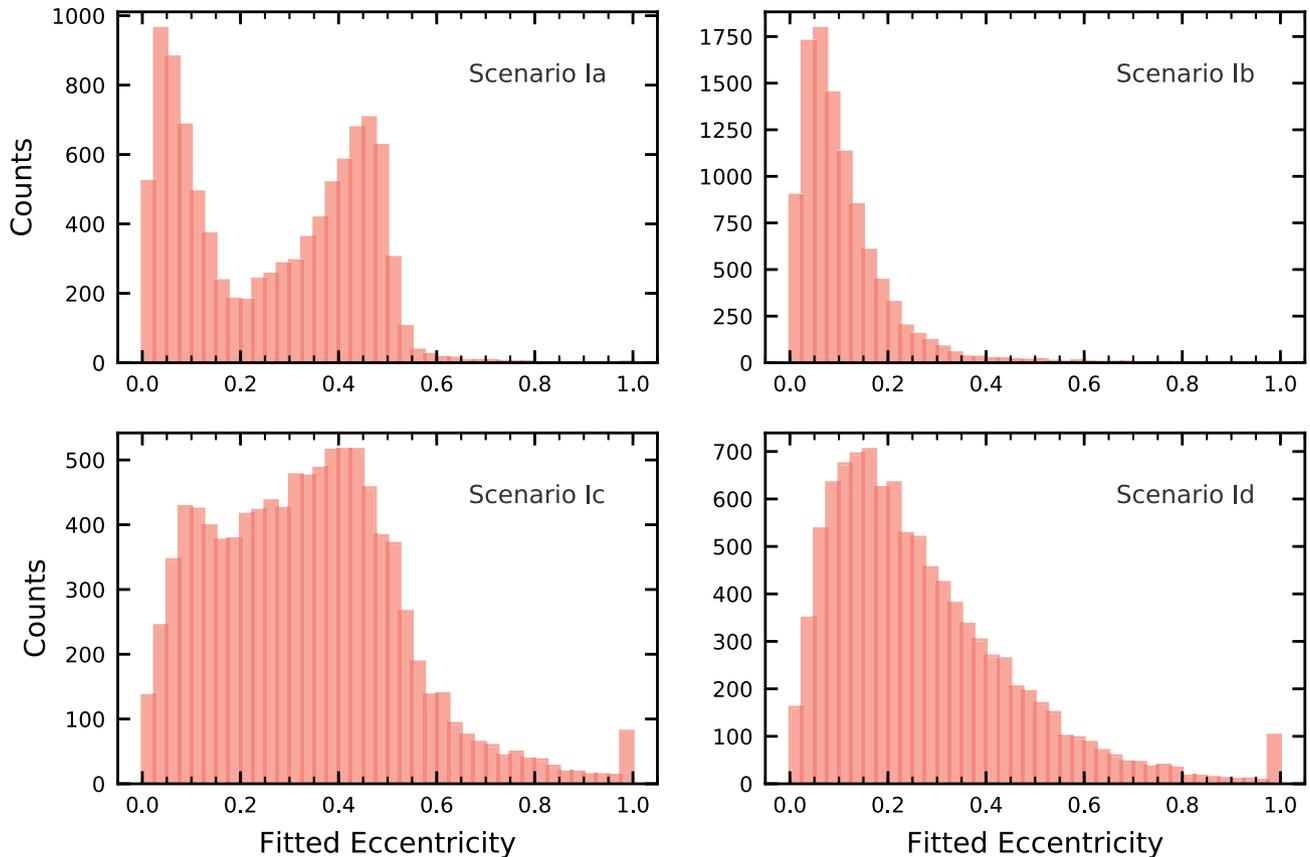
We simulated stellar velocity noise by choosing the radial velocity values and their uncertainties by a random draw from the AAPS data for six stable solar-type stars (531 epochs). These velocities have a mean of zero and an rms scatter of  $2.99 \text{ m s}^{-1}$ . In this way, we assume the input data are purely noise containing no planetary signals, and we have made no assumptions about the noise distribution (e.g. Gaussianity). The uncertainties, derived only from photon statistics, have a mean of  $1.1 \text{ m s}^{-1}$ . We then add  $3 \text{ m s}^{-1}$  of stellar jitter in quadrature to the individual measurement uncertainties. From our experience in least-squares fitting, this treatment reduces the possibility of the fitting routine getting stuck in local minima due to high-leverage points with small uncertainties. In the next subsection, we describe the process for generating the simulated Keplerian signals, which are added to the noise to produce the final simulated data.

### 2.3 Simulated planetary signals

For all simulated two-planet systems, the velocity amplitude  $K$  for each planet was assigned a random value between 20 and  $100 \text{ m s}^{-1}$ . This is largely arbitrary, but reflects values typical of securely detected radial velocity exoplanets. That is, the amplitudes are large enough to be unambiguously detectable in the presence of noise, yet small enough to remain in the planetary regime. Each scenario resulted in 10 000 synthetic radial velocity data sets.

*Scenario I* – First, we considered a simple circular-double configuration, with orbits whose periods are in a 2:1 commensurability (Scenario Ia) at arbitrarily chosen periods of 100 and 200 d. In light of the observed pile up of planets at the 2.17:1 period ratio (Steffen & Hwang 2015), we repeated this with planets at periods of 217 and 100 days, respectively (Scenario Ib).

*Scenario II* – Next, we considered the combination of planets moving on slightly eccentric orbits. Scenario IIa consists of two planets, both fixed at  $e = 0.1$ , and at periods of 100 and 200 d



**Figure 1.** Distribution of fitted eccentricities for Scenario I (two circular planets).

as above. Likewise, Scenario IIb considers planets on periods of 217 and 100 d. For these eccentric orbits, we set the periastron arguments at  $\omega_1 = 0$  and  $\omega_2$  as a random value on  $[0, 2\pi]$ . Hence, the relative apsidal alignment of the two planets is randomized.

*Scenario III* – Finally, we investigated a more realistic choice of orbital periods for two circular planets. In Scenario IIIa, we drew the period of the outer planet at random (with replacement) from a set of 673 planets detected by radial velocity, obtained from the NASA Exoplanet Archive. The inner planet signal was then chosen to be exactly half this period. As above, Scenario IIIb sets the inner planet period such that the two are in a 2.17:1 period ratio.

For each of the three scenarios above, we also created ‘c’ and ‘d’ subsets. The ‘c’ and ‘d’ subsets of the above scenarios (I, II, and III) are analogous to the ‘a’ and ‘b’ setups (2:1 and 2.17:1 respectively), but using the more realistic observation epoch times drawn from real data as described in Section 2.1.

## 2.4 Orbit fitting

Armed with synthetic data sets, we proceeded to fit each with a single Keplerian signal using the IDL package *RVLIN* (Wright & Howard 2009), which employs the Levenberg–Marquardt method for non-linear  $\chi^2$ -minimization. We obtained uncertainty estimates for the orbital parameters with the bootstrapping algorithm from the *BOOTTRAN* package (Wang et al. 2012). As pointed out by Wang et al. (2012), we note that bootstrapping is not an ideal method to determine parameter uncertainties in cases of sparsely sampled data, but our philosophy was to perform the fitting as blindly as possible. In order to save computational time, we limited our calculations

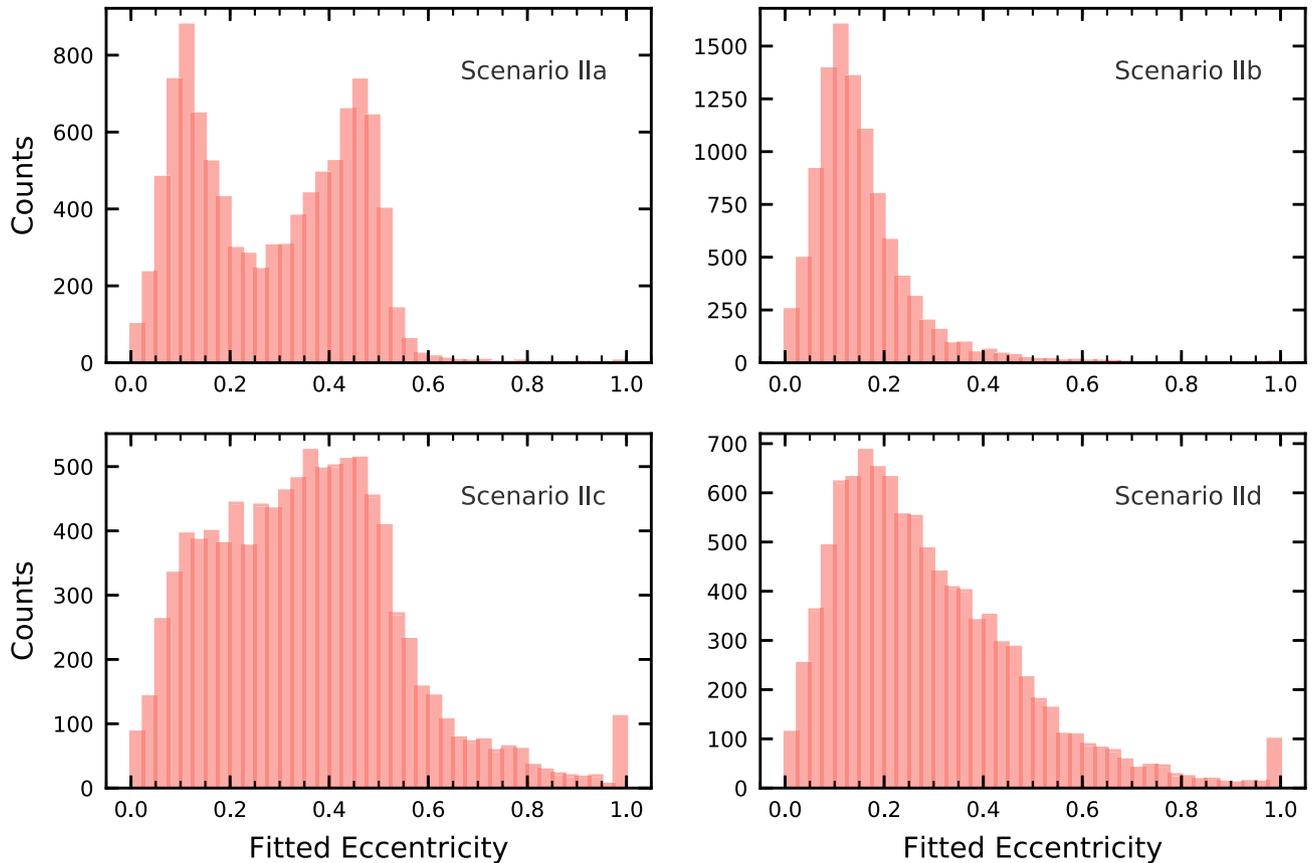
to 1000 bootstrap realizations; comparison of a small number of data sets for which we also obtained uncertainty estimates using 100 000 steps showed only marginal differences. The initial guess for the orbital period comes from the highest peak in a Lomb–Scargle periodogram (Lomb 1976; Scargle 1982; Horne & Baliunas 1986), and we used a default initial value of 0.3 for the eccentricity. We also employed an upper limit of 10 000 days for the orbital period. The only deviation from a completely blind fit we pursued was to prevent the fitting routine from getting stuck at  $e = 0$ , which sometimes happened especially if the fit was quite poor, and which we interpret as a peculiarity of the specific fitting package used. While not excluding circular orbits entirely, we automatically stepped through initial guesses of the periastron passage time if the initial fit returned an eccentricity of zero.

As a sanity check, we also attempted to fit each synthetic data set with a two-planet model, keeping the eccentricities fixed at either zero or 0.1, depending on the scenarios described above. Table 1 shows the degree to which this test successfully recovered the input periods of both planets to within 10 per cent.

## 3 RESULTS

Not surprisingly, the act of fitting a single planet to data containing two signals produced a wide range of results. In this work, we are most concerned with the eccentricity; Figs. 1–3 show the distributions of fitted eccentricities resulting from all scenarios.

The primary aim of this work is to determine the frequency and conditions which cause two circular (or nearly-circular) planets to masquerade as a single eccentric planet in realistically sampled



**Figure 2.** Distribution of fitted eccentricities for Scenario II (Two slightly eccentric planets).

radial velocity data. To quantify this, we must first set some criteria for a ‘plausible’ single-eccentric fit. As the input data consist of two Keplerian signals at two different periods, which we fit with a single planet, it is clear that a significant number of the resulting fits will be very poor. In defining ‘plausible’ fit results, we must quantify what humans have been intuitively doing across three decades of radial velocity fitting. We therefore establish five criteria that must be passed for a given eccentric-planet fit to be deemed plausible.

- (i) The fitted eccentricity must be at least  $3\sigma$  from zero.
- (ii) The fitted velocity amplitude  $K$  must be at least four times its own uncertainty:  $K/\sigma_K > 4.0$ . This is derived from the NASA Exoplanet Archive, in which 95 per cent of confirmed radial-velocity-detected planets satisfy this criterion.
- (iii) The fitted velocity amplitude must be at least 1.23 times larger than the rms scatter about the fit:  $K/\text{rms} > 1.23$ . As above, we choose this limit as that which holds true for 95 per cent of NASA Exoplanet Archive confirmed planets.
- (iv) The fitted period must be less than 1.5 times the total duration of the observations. This is derived from noting that virtually no radial velocity planet discoveries are published with less than about 0.7 orbital cycles of observations (cf. fig. 4 of Wittenmyer et al. (2011)).
- (v) The rms of the fit must be less than three times the mean measurement uncertainty:  $\text{rms}/\bar{\sigma} < 3.0$ .

While the final criterion is admittedly somewhat arbitrary, it eliminates obviously bad fits characterized by large residual scatter. Given that the input data always contain two signals, and we fit for only one, it is reasonable to expect a large number of instances in

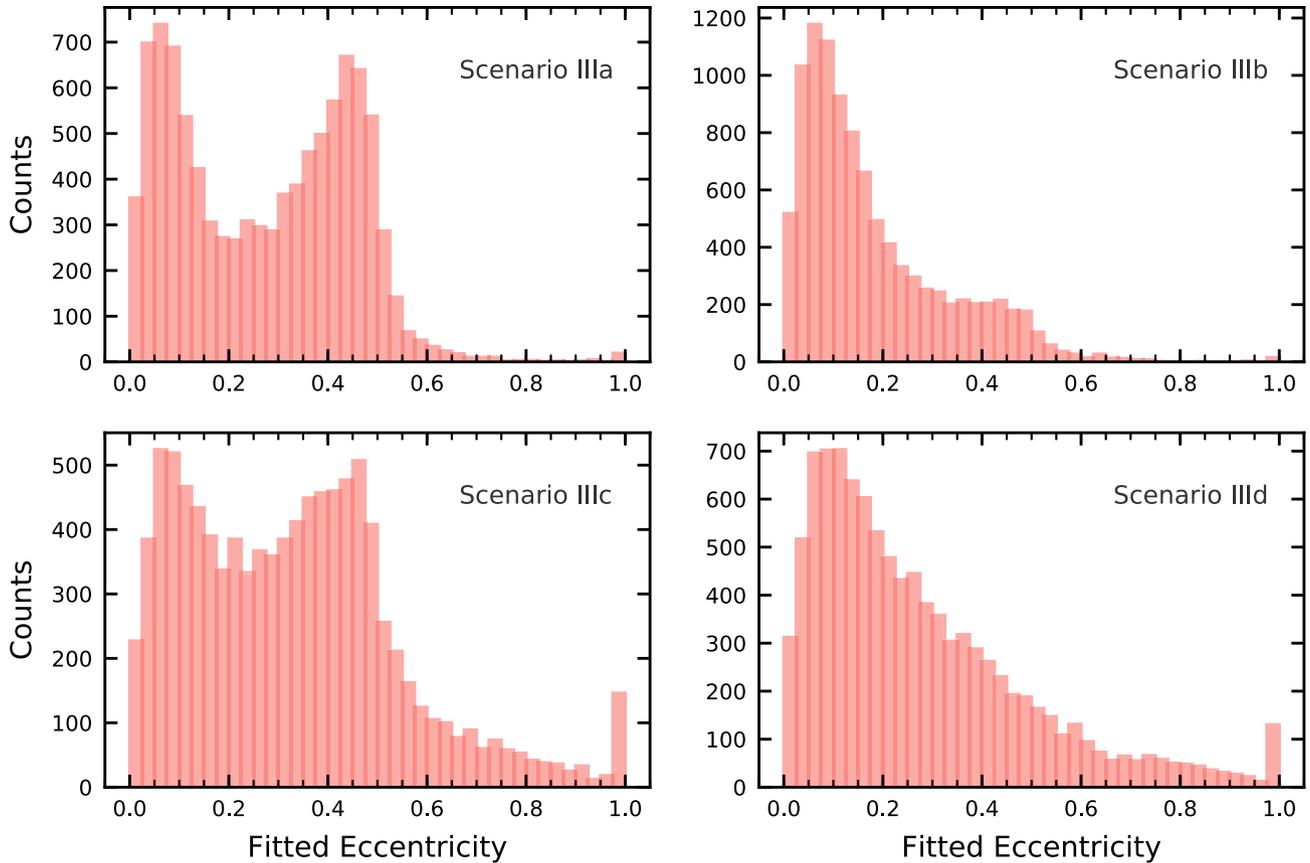
which RVLIN chooses one periodicity and arrives at a ‘best fit’ with an abominably large scatter.

Table 2 briefly summarizes the plausible single-eccentric fits resulting from the 12 trial scenarios described above, after applying these criteria. Scenario I, the double-circular configuration, resulted in 13.08 per cent plausible single-eccentric fits for the 2:1 period ratio ( $P_2 = 200$  d), but none for the 2.17:1 period ratio. Increasing the realism by drawing timestamps from real observations produced 19.04 per cent and 4.94 per cent such plausible fits for the 2:1 and 2.17:1 period ratios, respectively. Similar results were achieved in Scenario II, using the ‘slightly eccentric’  $e = 0.1$  configuration.

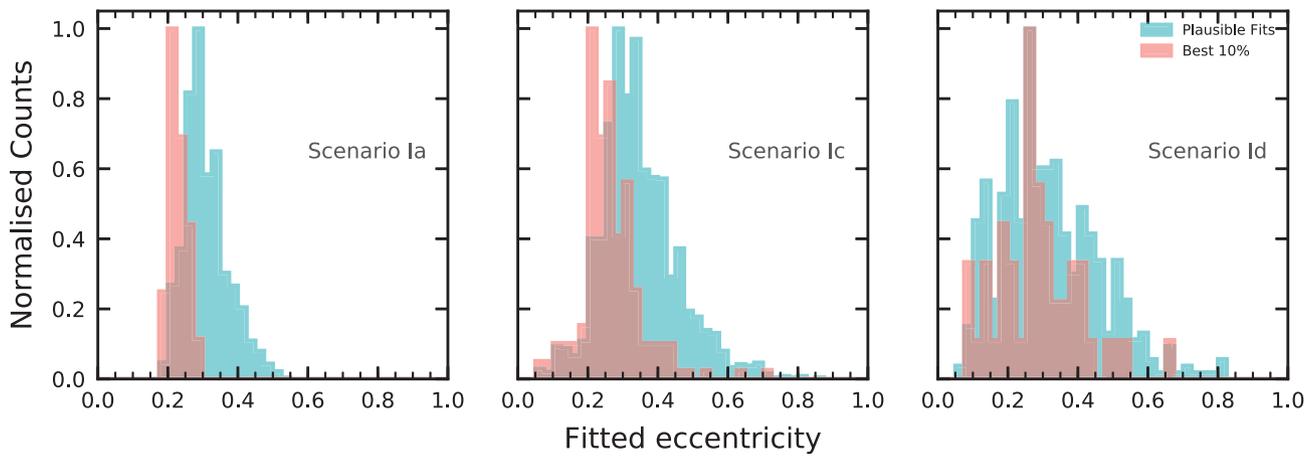
Scenario III, where the outer period  $P_2$  was drawn from the set of real planets, yielded the largest number of plausible single-eccentric fits, with 15.27 per cent (IIIa), 2.89 per cent (IIIb), 20.49 per cent (IIIc), and 9.30 per cent (IIId). The detailed characteristics of these plausible fits are discussed further in the next subsection.

### 3.1 Plausible eccentric single solutions

In this subsection, we investigate the properties of the plausible single-eccentric fits in more detail. Figs. 4–6 show as green histograms the distribution of fitted eccentricities obtained by the RVLIN single-planet fits which passed all five criteria described above. Overplotted as red histograms are the ‘best’ 10 per cent of plausible fits, those with the smallest ratio of rms scatter to mean measurement uncertainty. This selection resulted in fits with rms values of 2.9–5.1  $\text{m s}^{-1}$  (where the mean measurement uncertainty is  $\sim 3.2 \text{ m s}^{-1}$ ). These are fits which are most likely to convince the



**Figure 3.** Distribution of fitted eccentricities for Scenario III (two circular planets with ‘realistic’ orbital periods).



**Figure 4.** Distribution of fitted eccentricities for the ‘plausible’ single-eccentric fits found in Scenario Ia, Ic, and Id.

observer that a single eccentric planet is a good fit to the data, when in reality two planetary signals are ensconced within.

In Scenario I (periods fixed at 200:100 or 217:100 d), the bulk of the best fits were generally restricted to eccentricities between 0.2 and 0.4, with a handful of examples out to  $e \sim 0.6$  for Ic and Id, which sampled the simulated velocities using time series from real stars. Notably, the 2.17:1 arrangement (Scenario Ib; no histogram possible) never produced a reasonable single-eccentric solution. Scenario II, in which slightly eccentric ( $e = 0.1$ ) pairs

of input planets were tested, gave very similar results to Scenario I. Scenario III was the most realistic, with periods drawn from real radial velocity planets, and resulted in the highest fraction of plausible single-eccentric fits. As in the previous trials, the best fits clustered at lower eccentricities ( $e < 0.2$ ) but developed a long tail extending even to  $e > 0.9$  for Scenarios IIIc and IIIId (Fig. 6). Closer inspection revealed that the extreme examples were characterized by pathologies such as (1) reaching the maximum allowed period (10000 d) with an error bar of more than 200 per cent, or (2)

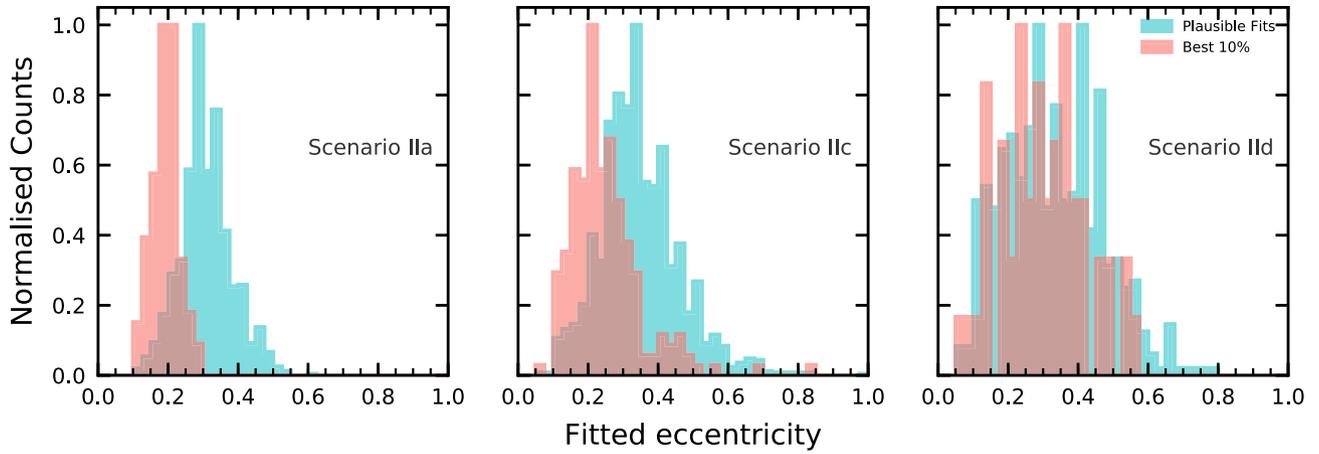


Figure 5. Same as Fig. 4, but for Scenario IIa, IIc, and IId.

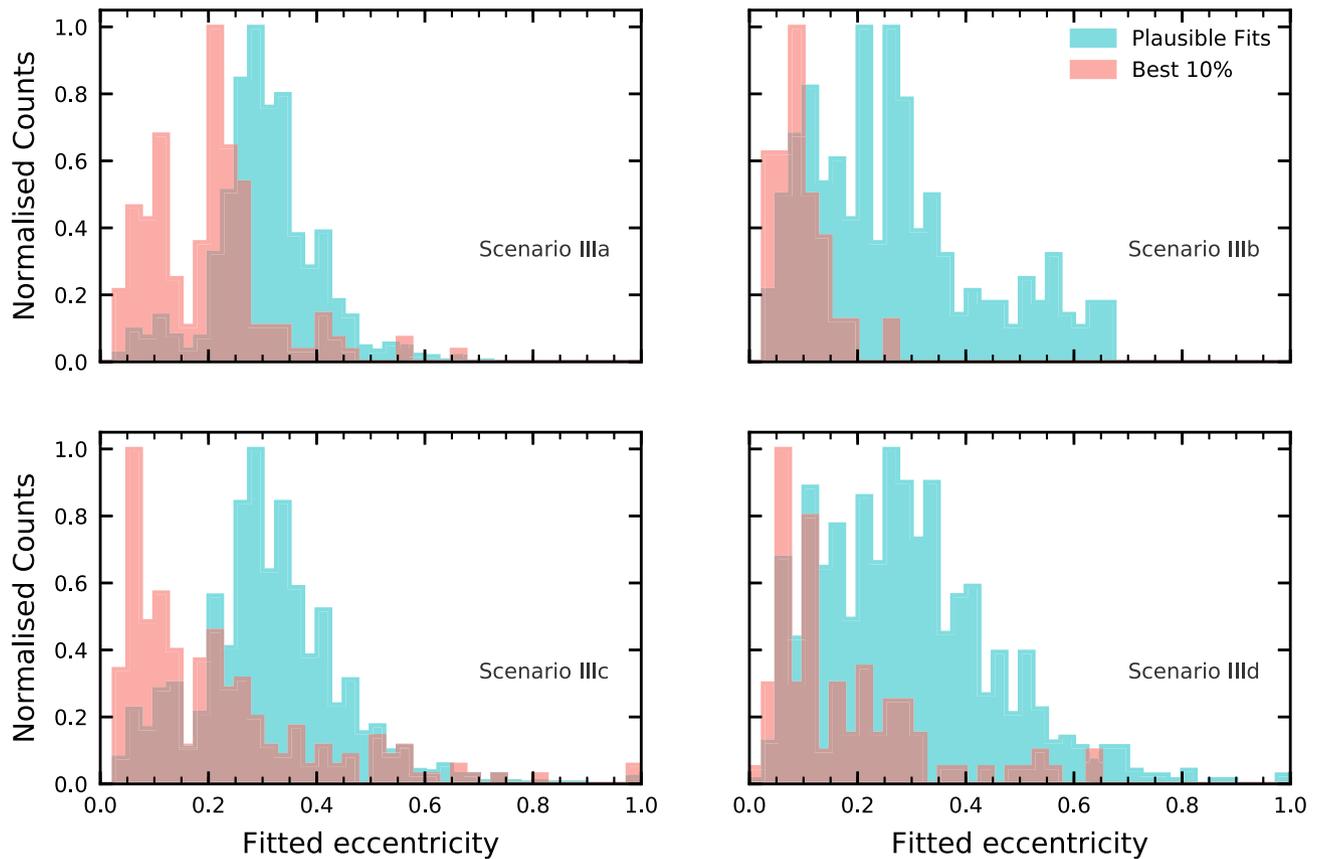


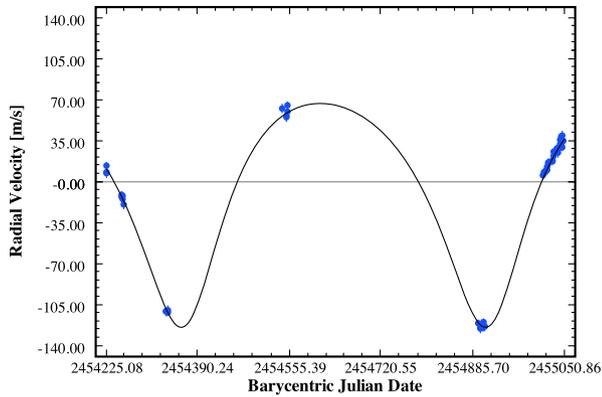
Figure 6. Same as Fig. 4, but for Scenario IIIa, IIIb, IIIc, and IIId.

amplitudes  $K > 500 \text{ m s}^{-1}$  with large uncertainties, likely driven by phase gaps in the time series (as the observation times were taken from real data). These ‘best’ fits were selected strictly by the lowest rms scatter, with no human intervention as yet, and so such oddities are to be expected. We explore the cause of these peculiarities further in the next section.

Fig. 7 shows an example of a simulated data set that yielded a ‘good’ single-eccentric fit, with  $e = 0.31 \pm 0.03$ ,  $K = 95.6 \pm 0.9 \text{ m s}^{-1}$ , and an rms of  $2.56 \text{ m s}^{-1}$ .

#### 4 DISCUSSION AND CONCLUSIONS

In this work, we consider the problematical false positive single-planet solutions that can arise from poorly sampled radial velocity observations of systems containing two exoplanets moving on near-circular orbits. With poor sampling and noisy data, the analysis of such systems can often return a convincing single-planet solution, with that planet moving on an orbit with moderate eccentricity. It is likely that a number of such near-circular exoplanetary



**Figure 7.** Example of a simulated data set for which the circular-double combination can be plausibly fit with a single eccentric model. Shown is an example from Scenario IIIId, chosen from the ‘best’ 10 per cent subset, to illustrate the pathology we investigate here. This fit has a period of 548.8 days,  $e = 0.31 \pm 0.03$ , and an rms of  $2.56 \text{ m s}^{-1}$ . The true injected data are two circular orbits with periods of 580.7 and 267.6 days.

couplets remain undiscovered among the many ‘confirmed’ single, moderately eccentric exoplanets.

The results of our analysis demonstrate the range of eccentricities for which it is conceivable that a circular two-planet system can masquerade as an eccentric single planet. Particularly for the most realistic trials (IIIc and IIIId), there was no obvious cutoff eccentricity beyond which circular-double systems failed to produce plausible single-eccentric fits. However, in reality, it seems reasonable to expect such a threshold to exist. This is due to the shape of the radial velocity orbit departing so far from sinusoidal that it realistically cannot be reproduced by only two circular components, which is the problem at the heart of this degeneracy (as noted in equation (1) of Boisvert et al. 2018). We see evidence of such a limit for the ‘a’ and ‘b’ subsets, in which the sampling turned out to be less realistic than for the ‘c’ and ‘d’ scenarios.

The grand mean-fitted eccentricity across all scenarios (Table 2) is  $e = 0.32 \pm 0.12$ , with a 95 per cent confidence interval of [0.09, 0.59]. We take this to be the ‘sweet spot’ (or danger zone) for deceptive configurations. If we consider only the ‘best’ 10 per cent of the plausible single-eccentric fits, we derive a grand mean  $e = 0.23 \pm 0.12$ , with a 95 per cent confidence interval of [0.05, 0.53].

**Table 2.** ‘Plausible’ single-eccentric fit results from 10 000 trials.

Scenario	Number	Mean eccentricity
Ia	1308	$0.31 \pm 0.06$
Ib	0	–
Ic	1904	$0.35 \pm 0.11$
Id	494	$0.32 \pm 0.15$
IIa	1238	$0.31 \pm 0.07$
IIb	0	–
IIc	2004	$0.35 \pm 0.11$
IId	548	$0.33 \pm 0.14$
IIIa	1527	$0.31 \pm 0.10$
IIIb	289	$0.28 \pm 0.16$
IIIc	2049	$0.31 \pm 0.13$
IIId	930	$0.29 \pm 0.16$

Our results therefore show that more eccentric planet candidates ( $e \gtrsim 0.5$ ) are exceedingly unlikely to be mimics due to this degeneracy.

As noted above, the best such fits (in an rms sense) sometimes exhibited pathologies such as extreme amplitudes brought on by phase gaps, or fitted periods at the upper boundary and with outsized error bars. In Scenario III, up to 9 per cent of trials failed even to recover the two injected planets when subjected to a proper double-Keplerian fit (Table 1). We repeated the five tests described in Section 3 for the subset of trials which passed the sanity checking (see Section 2.4). We show the results for this new set of ‘plausible’ results in Fig. 8. For Scenarios IIIc and IIIId, which were most affected, we indeed see that all but 2–3 of the  $e > 0.6$  fits are eliminated. A handful remain, but the vast majority of ‘good’ single-eccentric fits have  $e \lesssim 0.4$ . This is again consistent with our result that the ‘danger zone’ lies generally in the range between  $e \sim 0.2$  and  $e \sim 0.4$ .

The results presented herein are also relevant to exoplanet searches that utilize the transit method. The relative lack of orbital phase coverage from transiting planets without accompanying RV data makes the reliable extraction of orbital eccentricity from transit light curves a challenging endeavour (e.g. Kipping 2008; Van Eylen & Albrecht 2015). Statistical studies of orbital eccentricities derived from *Kepler* discoveries have found the eccentricity distribution of transiting planets to be consistent with that from the RV exoplanet population (Kane et al. 2012). It was further concluded by Kane et al. (2012) and later by Van Eylen & Albrecht (2015) that there is a negative correlation of orbital eccentricity with planet size, particularly for those planets in compact systems. These factors underline both the need for complementary RV observations of transiting planets to reliably extract orbital eccentricities, and the potential degeneracy of circular orbits in the terrestrial regime. Exoplanet discoveries from the Transiting Exoplanet Survey Satellite (*TESS*) are expected to primarily be in relation to relatively bright host stars where such a complement of precision photometry and RV data will be far more accessible than for the *Kepler* systems (Ricker et al. 2015). In particular, extended mission scenarios for the *TESS* mission, such as those described by Sullivan et al. (2015), allow for the detection of longer period planets that will be more likely to have larger eccentricities.

Through the course of this work, we have demonstrated the impact of poor sampling on the veracity of convincing radial velocity exoplanet detections. Our work reveals the regime for which most caution should be exercised when considering whether a single-planet fit to observational data is a reflection of the true reality of the system in question. At the same time, however, our results add credibility to the detections of exoplanets moving on highly eccentric orbits, such as HD 80606b ( $e = 0.933 \pm 0.001$ ; Wittenmyer et al. 2007), HD 4113b ( $e = 0.903 \pm 0.005$ ; Tamuz et al. 2008) and HD 76920b ( $e = 0.856 \pm 0.009$ ; Wittenmyer et al. 2017b). While such planets are clearly oddities, our results suggest that they are not false-positive, ghost planets. Indeed, our results suggest that deceptive planetary couplets will rarely, if ever, masquerade as single planets with orbital eccentricities greater than  $e \sim 0.6$ , consistent with the results of Kürster et al. (2015) as shown in their fig. 4.

In future work, we intend to extend this analysis still further, examining a wider range of planetary couplet period ratios, orbital eccentricities, and masses. In addition, it is interesting to consider whether similar effects would result from deceptive triplets or quadruplets – in other words, whether it might be possible to mimic a signal of arbitrarily large eccentricity through the superposition

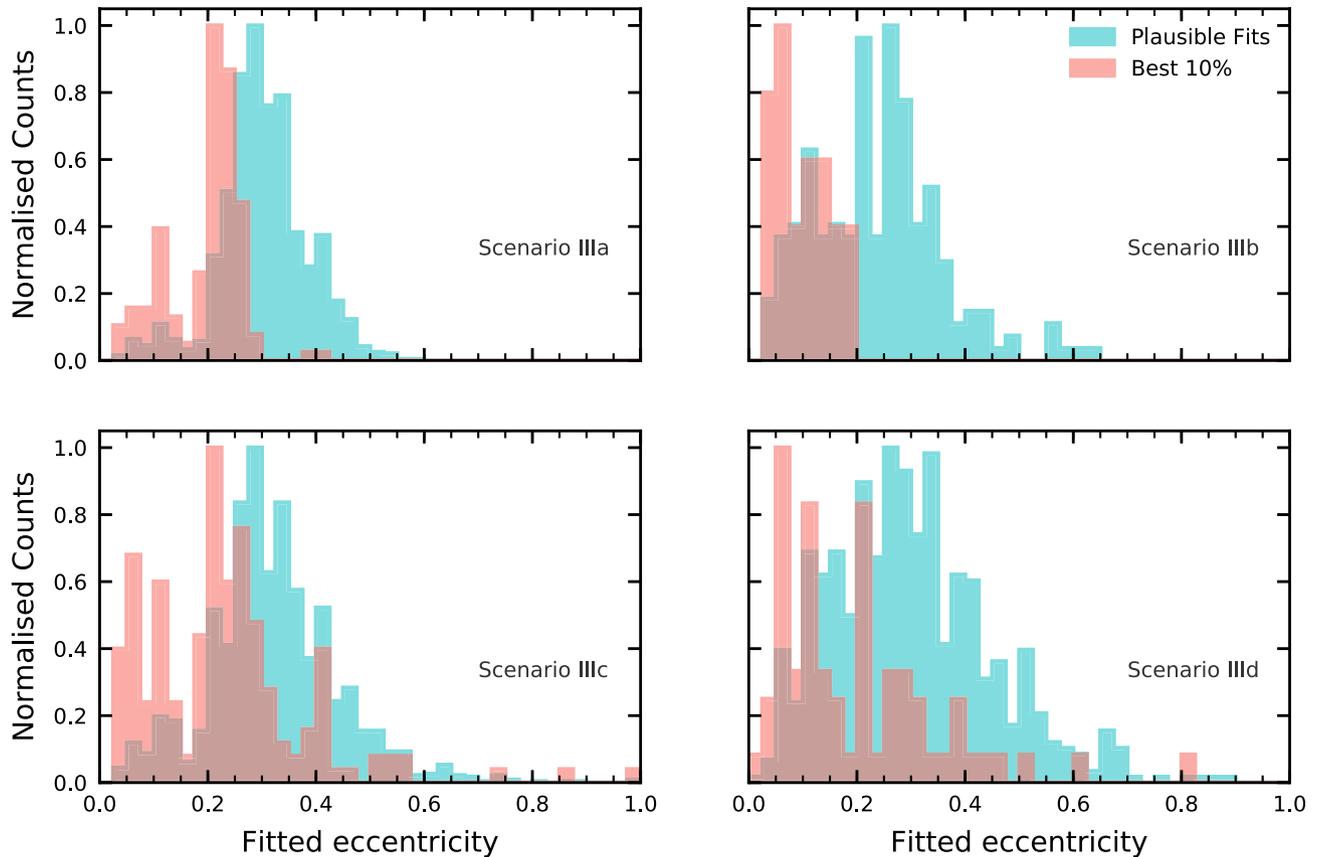


Figure 8. Same as Fig. 6, but with the failed data sets excluded.

of a given number ( $N > 2$ ) of signals resulting from planets moving on near-circular orbits.

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